

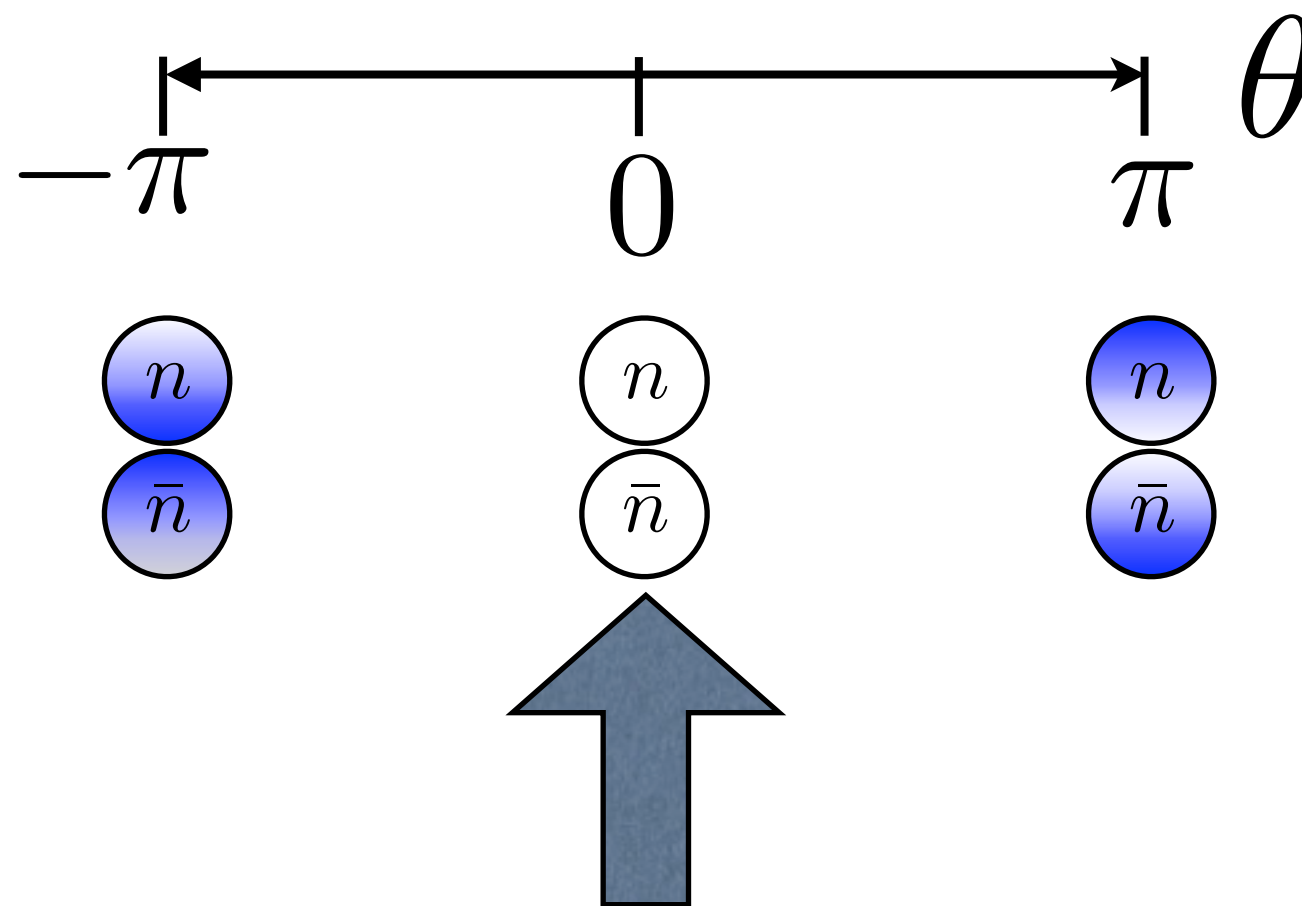
Axion Dark matter

Experimental review

Javier Redondo
(Zaragoza U. & MPP)

The theta angle of the strong interactions

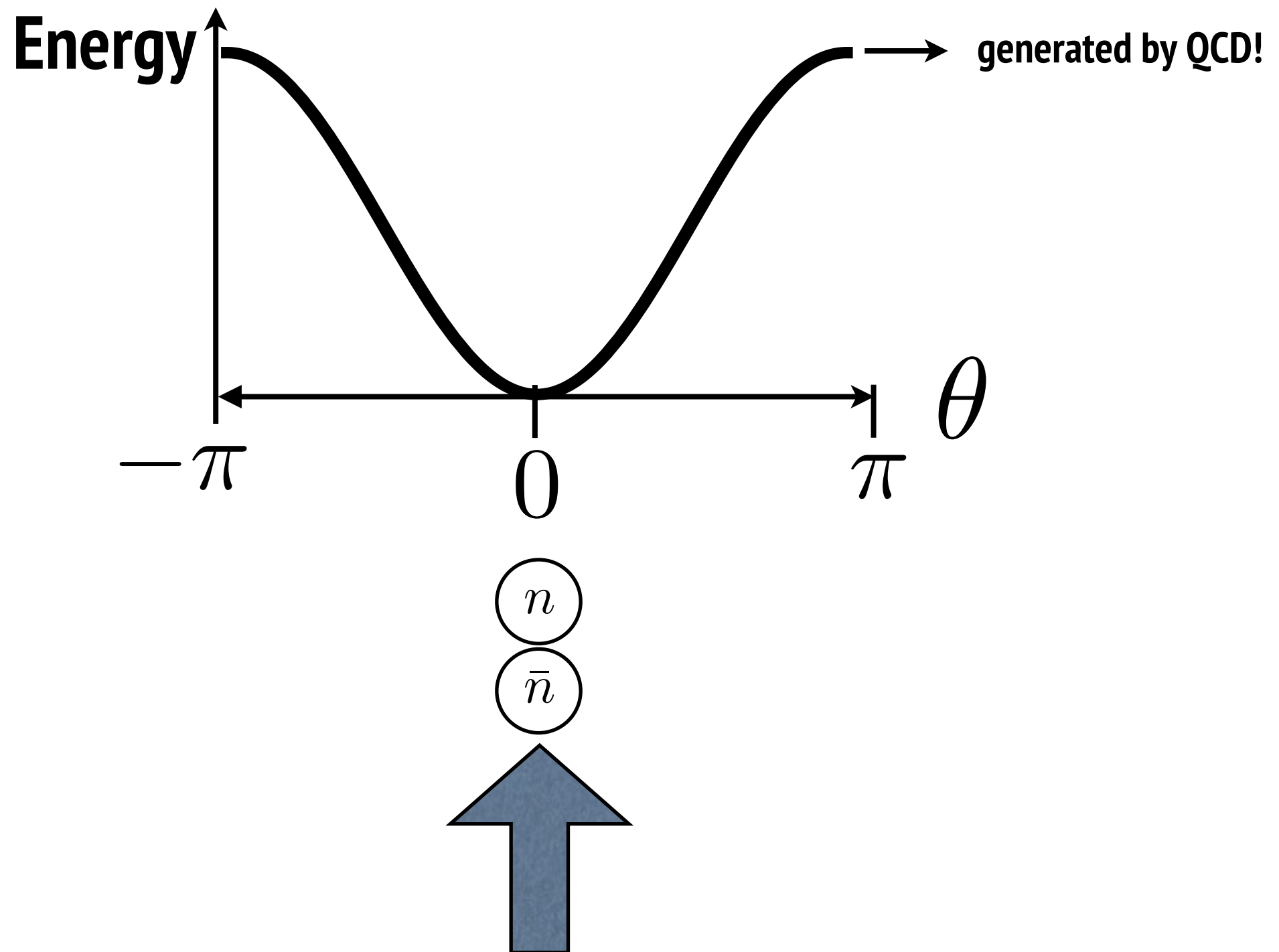
- The value of θ controls matter-antimatter differences in QCD
- In particular neutron (and proton) Electric dipole moment



Measured today $|\theta| < 10^{-10}$ (strong CP problem)

Axions

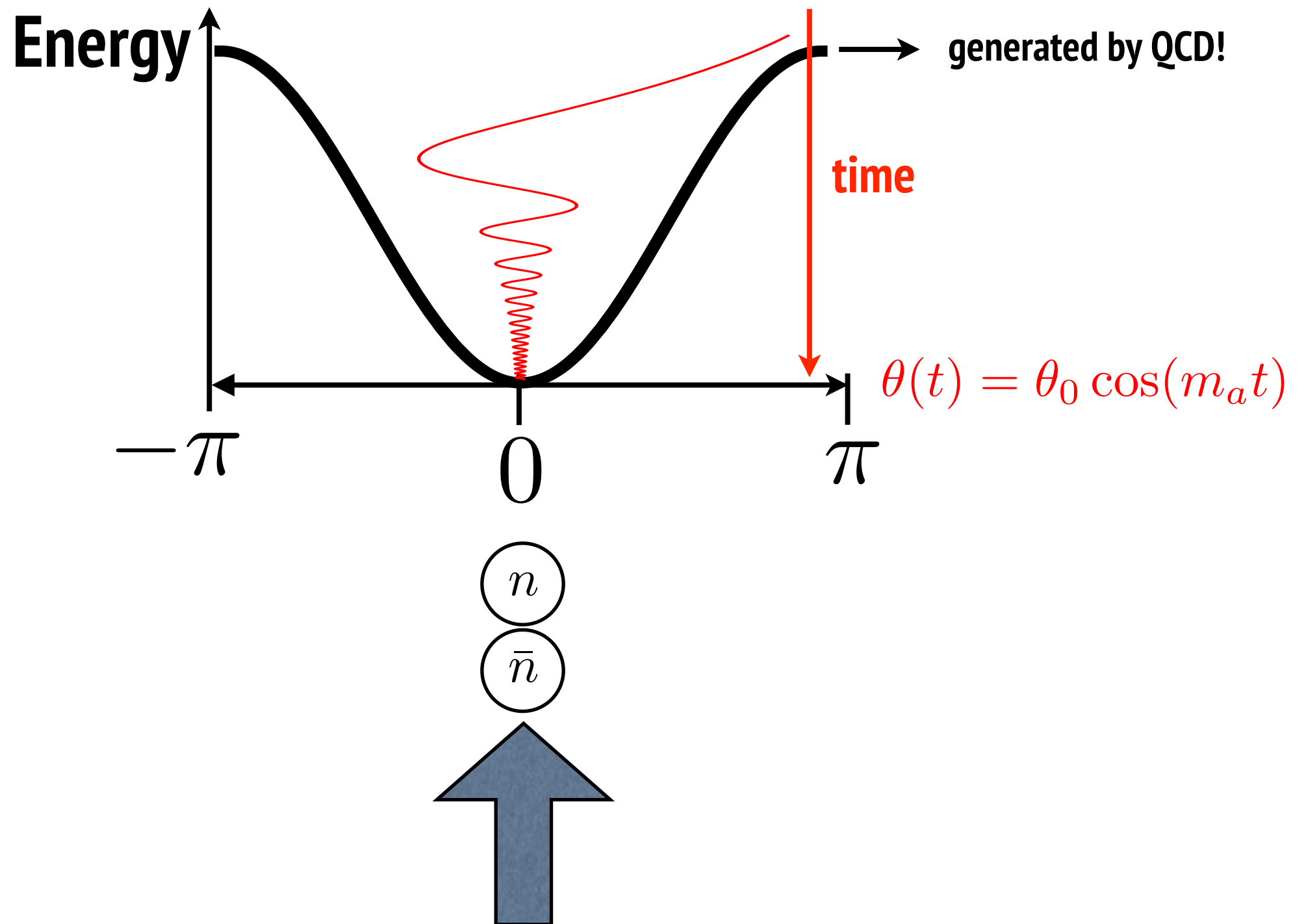
- is it a dynamical field? $\theta(t, \mathbf{x})$



Measured today $|\theta| < 10^{-10}$ (strong CP problem)

Axions

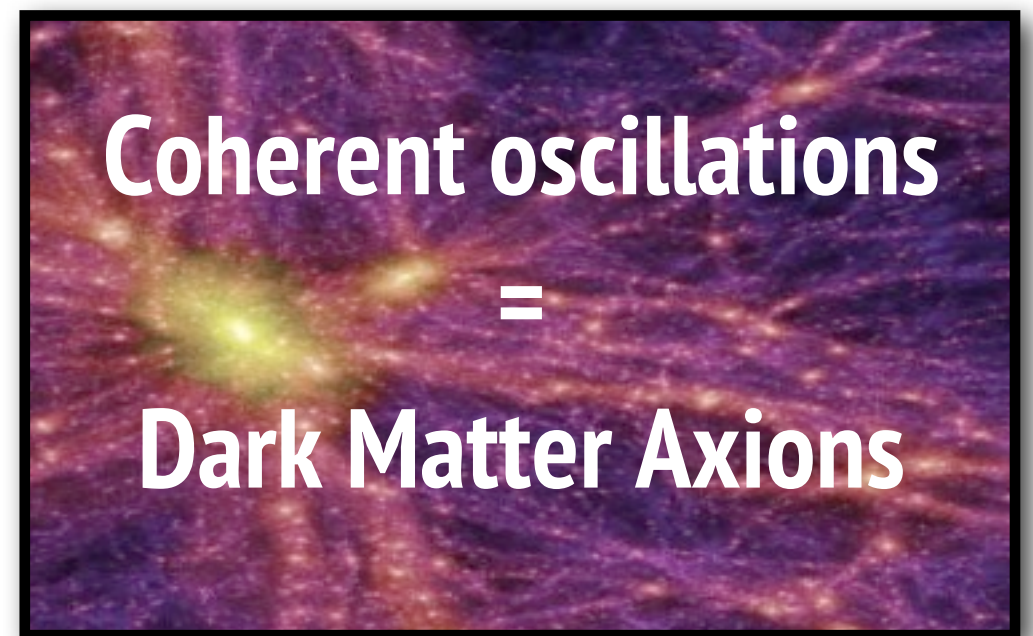
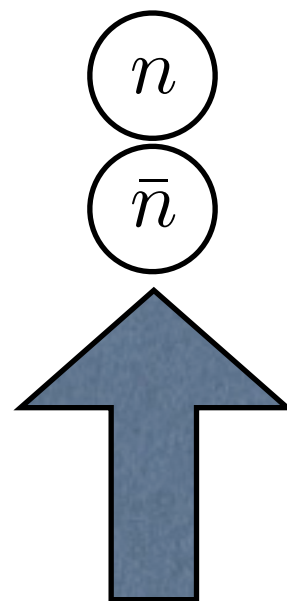
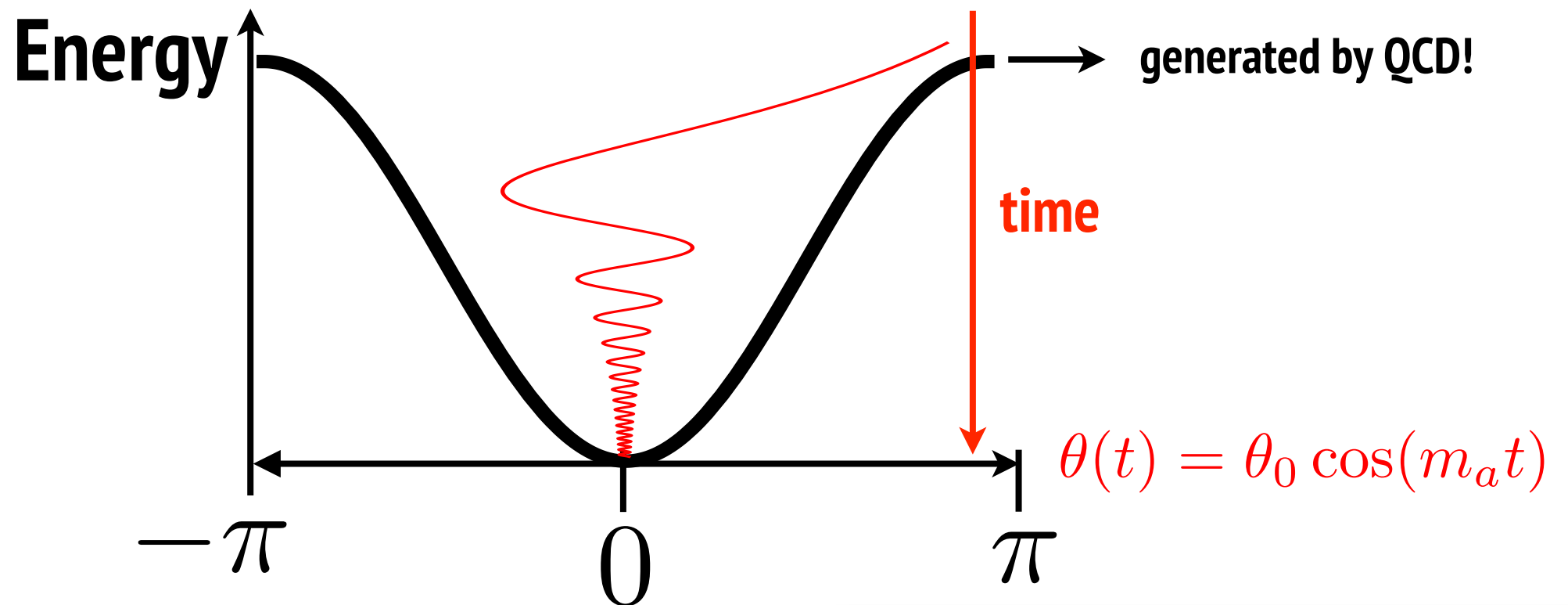
- is it a dynamical field? $\theta(t, \mathbf{x})$



Measured today $|\theta| < 10^{-10}$ (strong CP problem)

Axions

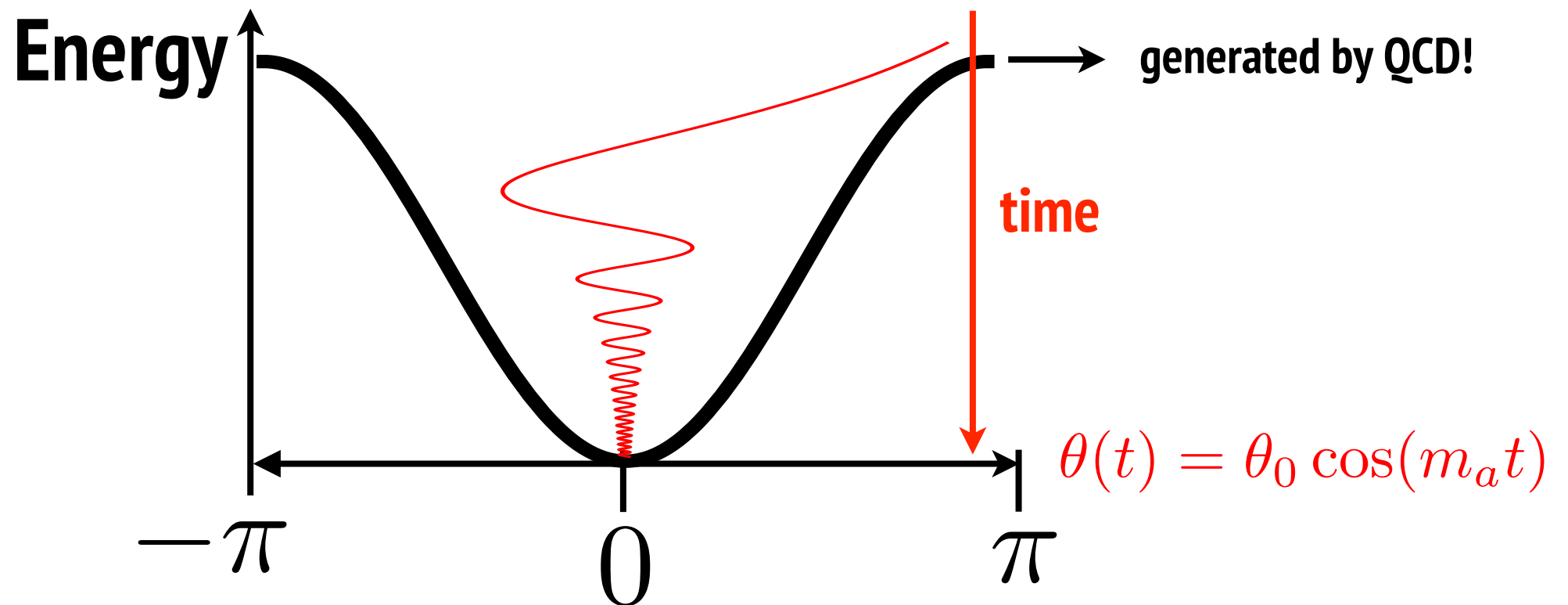
- is it a dynamical field? $\theta(t, \mathbf{x})$



Measured today $|\theta| < 10^{-10}$ (strong CP problem)

Axions

- is it a dynamical field? $\theta(t, \mathbf{x})$

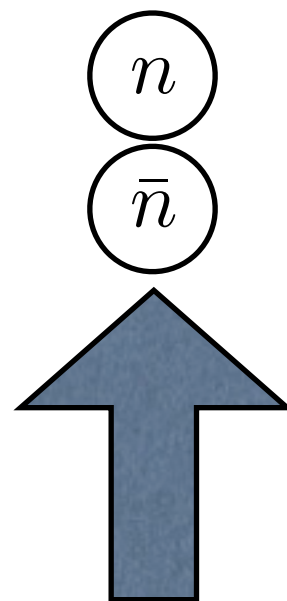


~ One parameter theory

$$\theta(t, x) = a(t, x)/f_a$$

axion mass

$$m_a = 6 \text{ meV} \frac{10^9 \text{ GeV}}{f_a}$$



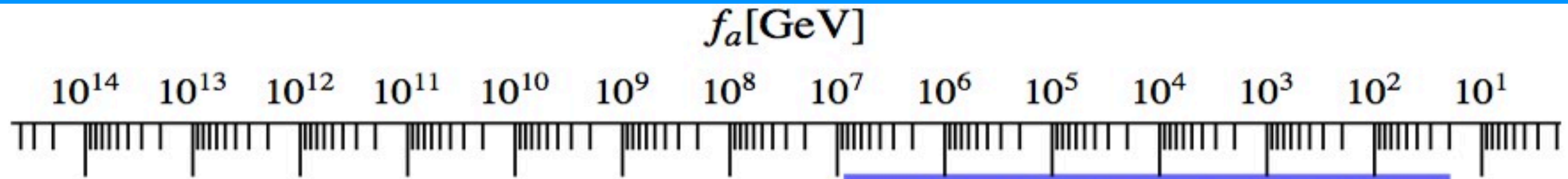
Coherent oscillations

=

Dark Matter Axions

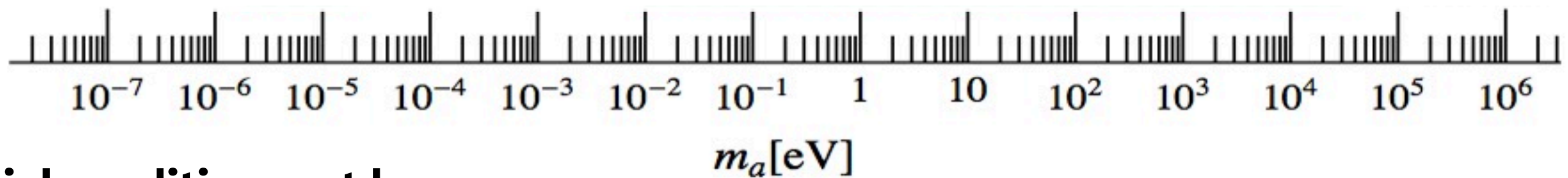
Measured today $|\theta| < 10^{-10}$ (strong CP problem)

Axion dark matter scenarios



- Axion DM scenarios

tuned (anthropic?) ok tuned

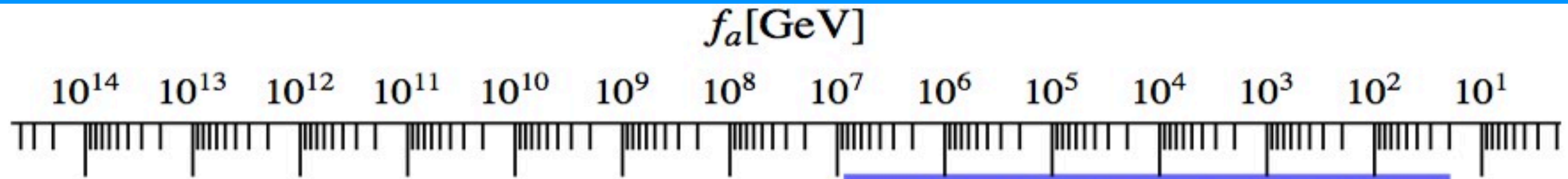


Initial conditions set by :

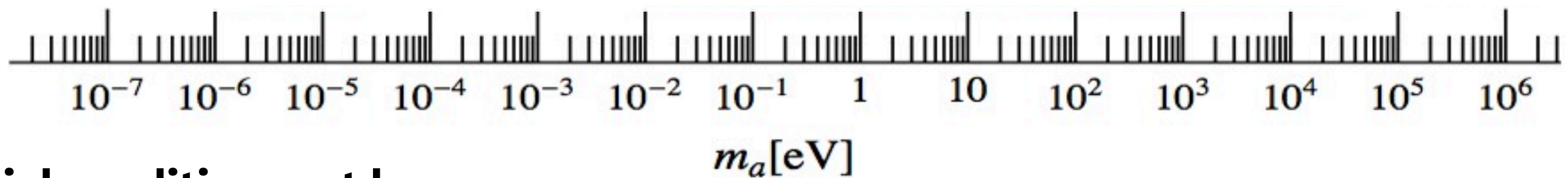
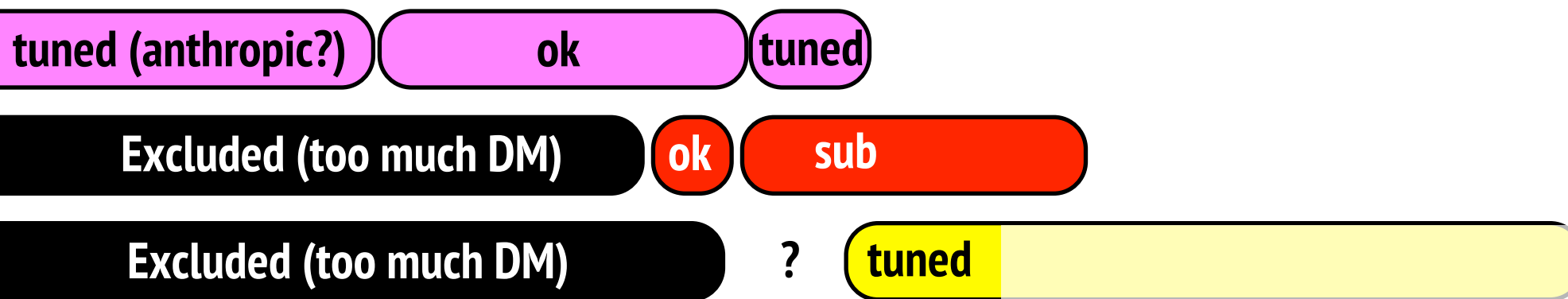
Inflation smooth

$$\Omega_{\text{aDM}} h^2 \simeq \theta_I^2 \left(\frac{80 \mu\text{eV}}{m_a} \right)^{1.19}$$

Axion dark matter scenarios



- Axion DM scenarios



Initial conditions set by :

Inflation smooth

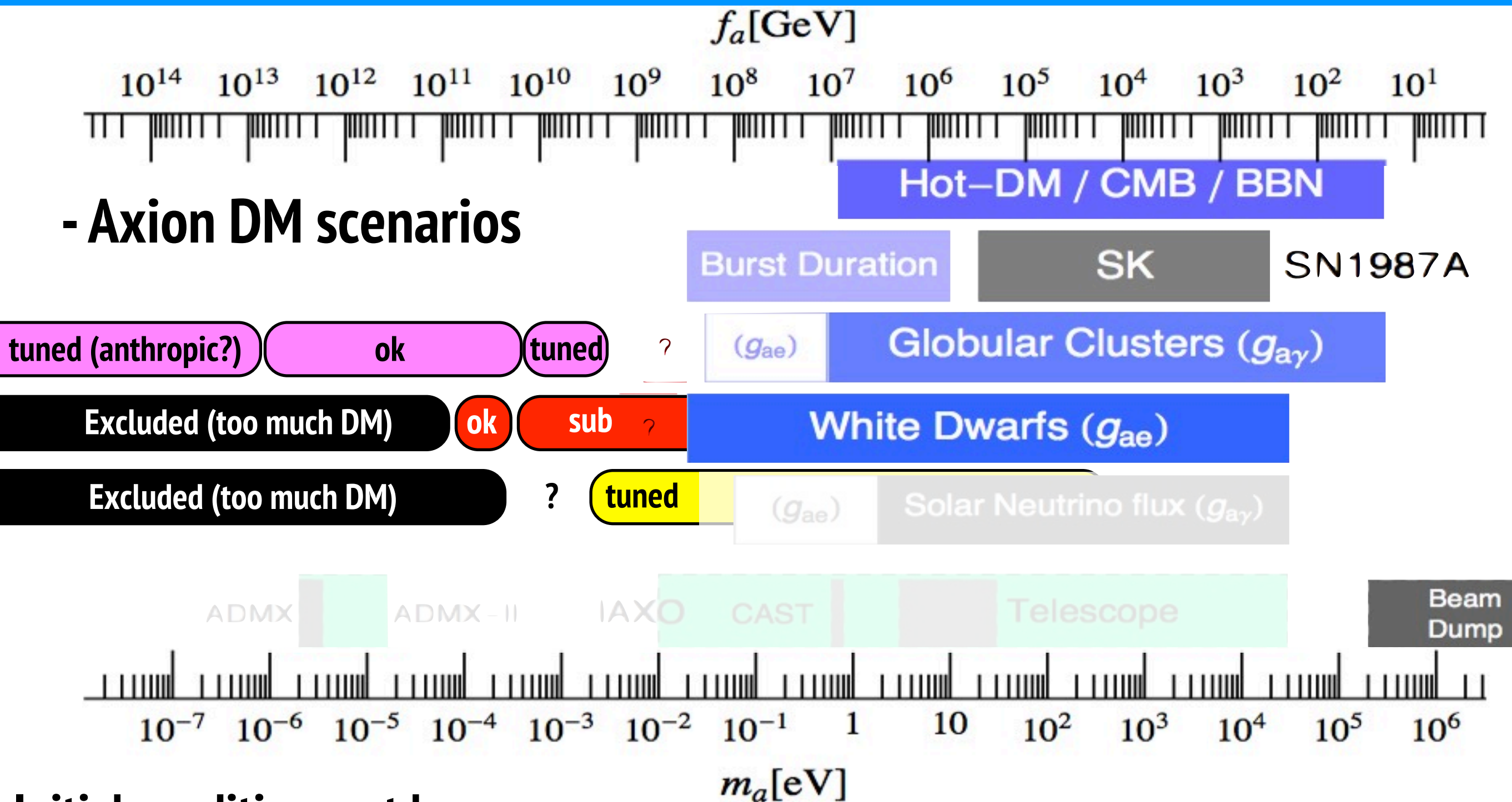
$$\Omega_{\text{aDM}} h^2 \simeq \theta_I^2 \left(\frac{80 \mu\text{eV}}{m_a} \right)^{1.19}$$

Phase transition (N=1)
strings+unstable DW's

Phase transition (N>1)
strings+long-lived DWs

Axion dark matter scenarios

- Axion DM scenarios



Initial conditions set by :

Inflation smooth

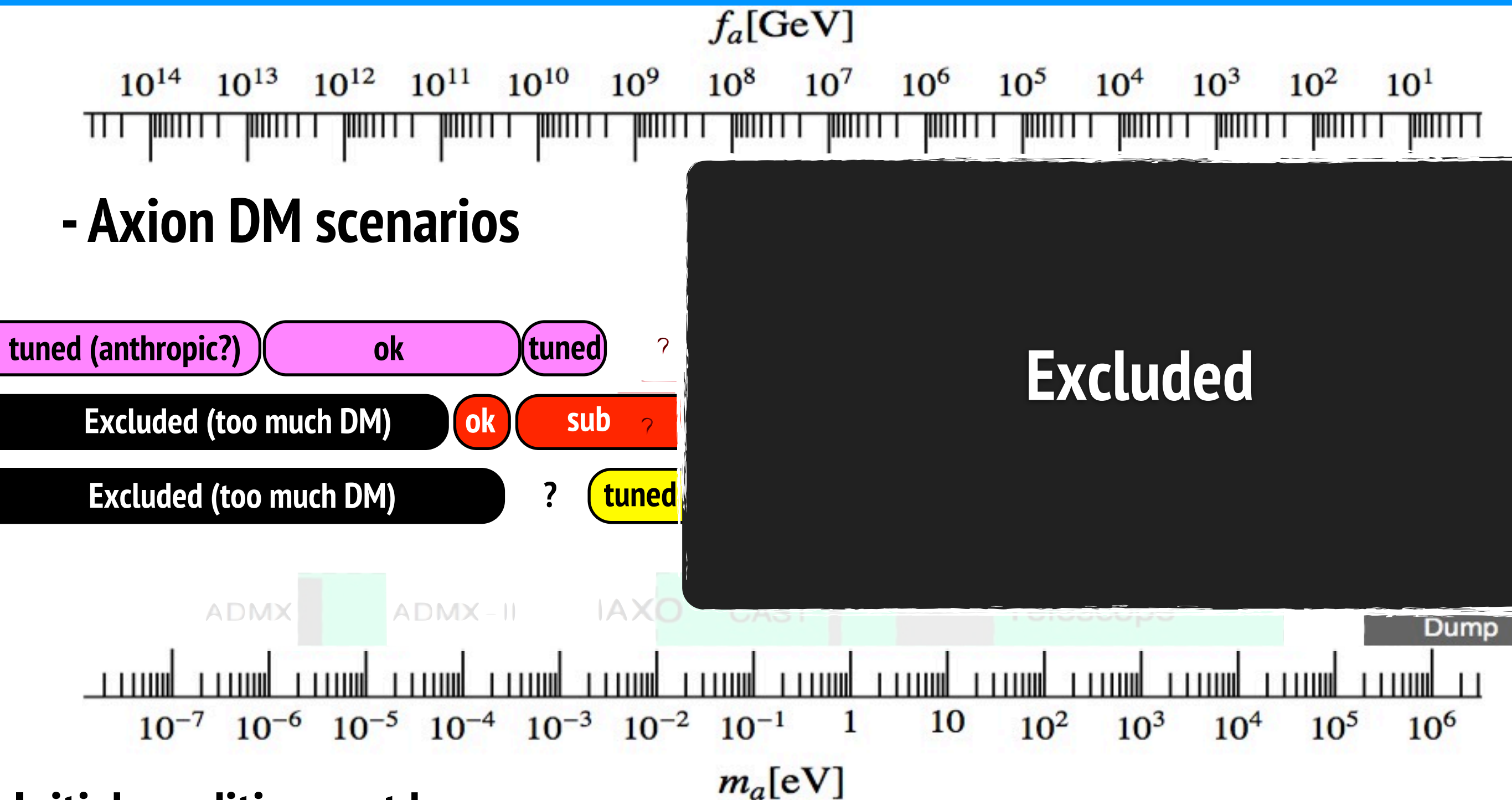
$$\Omega_{\text{aDM}} h^2 \simeq \theta_I^2 \left(\frac{80 \mu\text{eV}}{m_a} \right)^{1.19}$$

Phase transition (N=1)
strings+unstable DW's

Phase transition (N>1)
strings+long-lived DWs

Axion dark matter scenarios

- Axion DM scenarios



Initial conditions set by :

Inflation smooth

$$\Omega_{\text{aDM}} h^2 \simeq \theta_I^2 \left(\frac{80 \mu\text{eV}}{m_a} \right)^{1.19}$$

Phase transition (N=1)
strings+unstable DW's

Phase transition (N>1)
strings+long-lived DWs

Axion dark matter scenarios

**Dark Matter
huge parameter space!**

$f_a[\text{GeV}]$

$10^8 \quad 10^7 \quad 10^6 \quad 10^5 \quad 10^4 \quad 10^3 \quad 10^2 \quad 10^1$

tuned (anthropic?)

ok

tuned

?

Excluded (too much DM)

ok

sub

?

Excluded (too much DM)

?

tuned

Excluded

ADMX

ADMX-II

IAXO

CAS

PERSEUS

Dump

$10^{-7} \quad 10^{-6} \quad 10^{-5} \quad 10^{-4} \quad 10^{-3} \quad 10^{-2} \quad 10^{-1} \quad 1 \quad 10 \quad 10^2 \quad 10^3 \quad 10^4 \quad 10^5 \quad 10^6$

$m_a[\text{eV}]$

Initial conditions set by :

Inflation smooth

$$\Omega_{\text{aDM}} h^2 \simeq \theta_I^2 \left(\frac{80 \mu\text{eV}}{m_a} \right)^{1.19}$$

**Phase transition (N=1)
strings+unstable DW's**

**Phase transition (N>1)
strings+long-lived DWs**

Axions in the galactic Halo

Energy density $\theta(t) = \theta_0 \cos(m_a t)$

$$\rho_{\text{CDM}} = 0.3 \frac{\text{GeV}}{\text{cm}^3} \equiv \frac{1}{2}(\dot{a})^2 + \frac{1}{2}m_a^2 a^2 = \frac{1}{2}m_a^2 f_a^2 \theta_0^2 \longrightarrow \theta_0 \sim 3.6 \times 10^{-19}$$

Axions in the galactic Halo

Energy density $\theta(t) = \theta_0 \cos(m_a t)$

$$\rho_{\text{CDM}} = 0.3 \frac{\text{GeV}}{\text{cm}^3} \equiv \frac{1}{2}(\dot{a})^2 + \frac{1}{2}m_a^2 a^2 = \frac{1}{2}m_a^2 f_a^2 \theta_0^2 \longrightarrow \theta_0 \sim 3.6 \times 10^{-19}$$

Fourier-transform $a(x)$

$$\omega \simeq m_a(1 + v^2/2 + \dots)$$

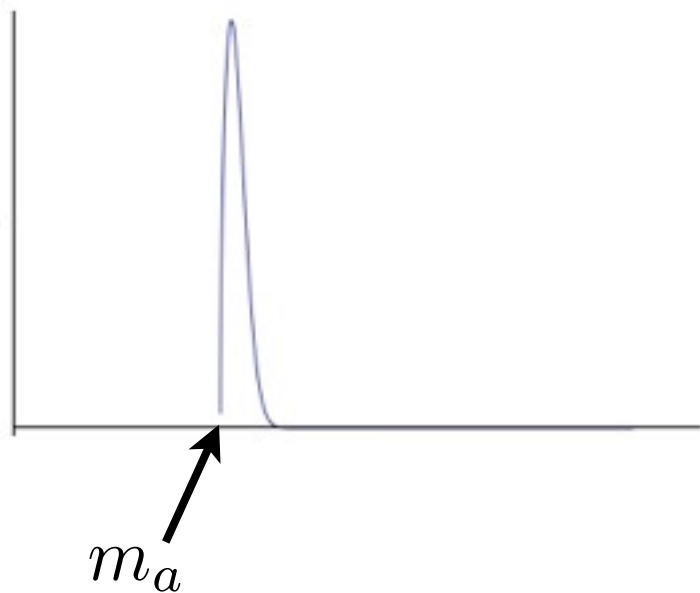
Axions in the galactic Halo

Energy density $\theta(t) = \theta_0 \cos(m_a t)$

$$\rho_{\text{CDM}} = 0.3 \frac{\text{GeV}}{\text{cm}^3} \equiv \frac{1}{2}(\dot{a})^2 + \frac{1}{2}m_a^2 a^2 = \frac{1}{2}m_a^2 f_a^2 \theta_0^2 \longrightarrow \theta_0 \sim 3.6 \times 10^{-19}$$

Fourier-transform $a(x)$

$$\omega \simeq m_a(1 + v^2/2 + \dots)$$



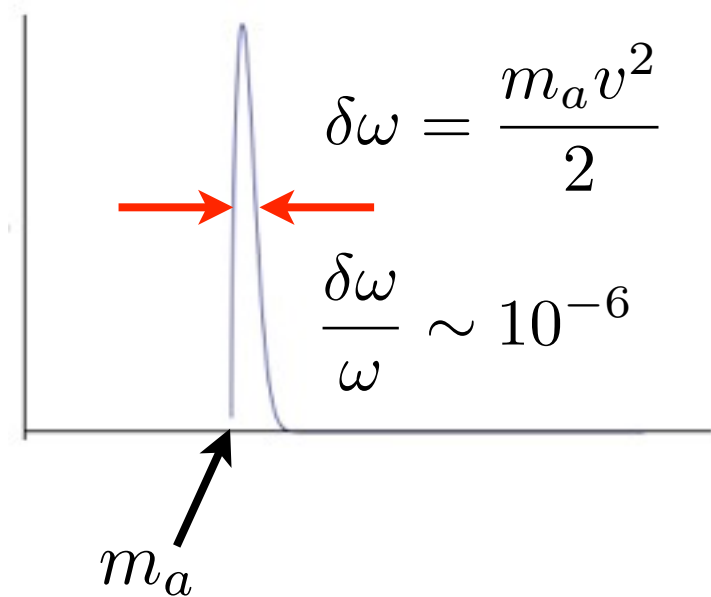
Axions in the galactic Halo

Energy density $\theta(t) = \theta_0 \cos(m_a t)$

$$\rho_{\text{CDM}} = 0.3 \frac{\text{GeV}}{\text{cm}^3} \equiv \frac{1}{2}(\dot{a})^2 + \frac{1}{2}m_a^2 a^2 = \frac{1}{2}m_a^2 f_a^2 \theta_0^2 \longrightarrow \theta_0 \sim 3.6 \times 10^{-19}$$

Fourier-transform $a(x)$

$$\omega \simeq m_a(1 + v^2/2 + \dots)$$



Axions in the galactic Halo

Energy density $\theta(t) = \theta_0 \cos(m_a t)$

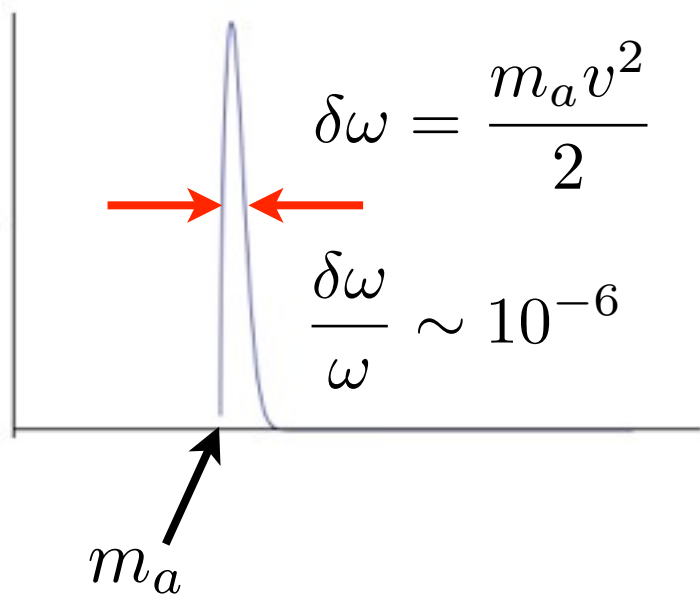
$$\rho_{\text{CDM}} = 0.3 \frac{\text{GeV}}{\text{cm}^3} \equiv \frac{1}{2}(\dot{a})^2 + \frac{1}{2}m_a^2 a^2 = \frac{1}{2}m_a^2 f_a^2 \theta_0^2 \longrightarrow \theta_0 \sim 3.6 \times 10^{-19}$$

Fourier-transform $a(x)$

$$\omega \simeq m_a(1 + v^2/2 + \dots)$$

coherence time

$$\delta t \sim \frac{1}{\delta \omega} \sim 0.13 \text{ms} \left(\frac{10^{-5} \text{eV}}{m_a} \right)$$



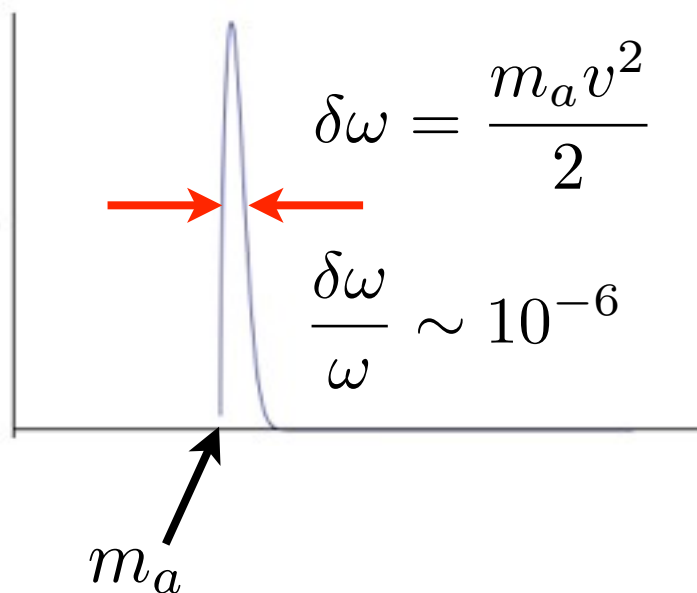
Axions in the galactic Halo

Energy density $\theta(t) = \theta_0 \cos(m_a t)$

$$\rho_{\text{CDM}} = 0.3 \frac{\text{GeV}}{\text{cm}^3} \equiv \frac{1}{2}(\dot{a})^2 + \frac{1}{2}m_a^2 a^2 = \frac{1}{2}m_a^2 f_a^2 \theta_0^2 \longrightarrow \theta_0 \sim 3.6 \times 10^{-19}$$

Fourier-transform $a(x)$

$$\omega \simeq m_a(1 + v^2/2 + \dots)$$



coherence time

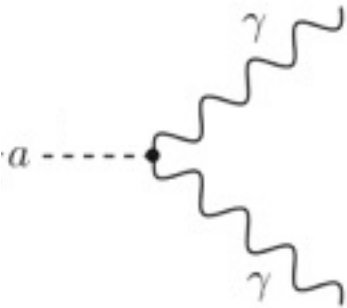
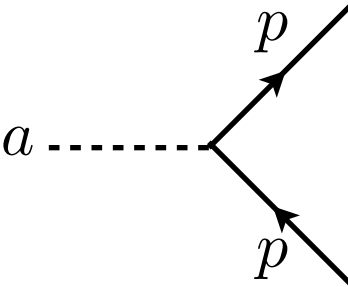
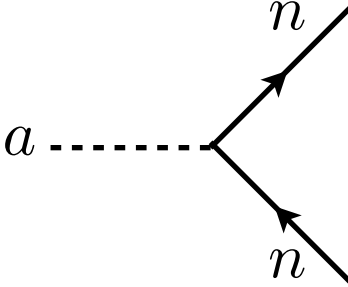
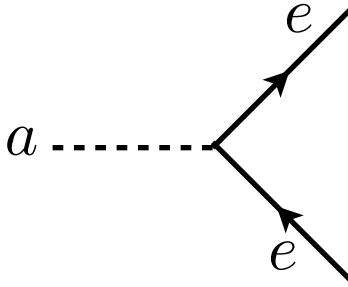
$$\delta t \sim \frac{1}{\delta\omega} \sim 0.13\text{ms} \left(\frac{10^{-5}\text{eV}}{m_a} \right)$$

coherence length

$$\delta L \sim \frac{1}{\delta p} \sim 20\text{m} \left(\frac{10^{-5}\text{eV}}{m_a} \right)$$

Detection channels, Axion Couplings

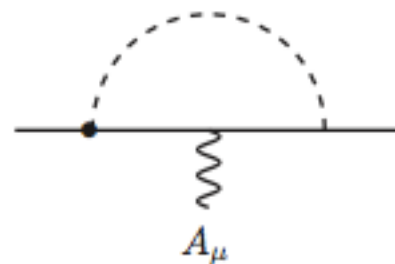
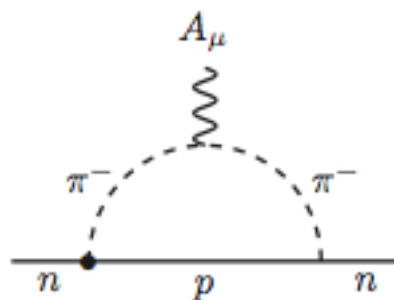
$$\frac{\alpha_s}{8\pi} \theta G_{\mu\nu} \tilde{G}^{\mu\nu} + \text{m.d.} \rightarrow \frac{\alpha C_{a\gamma}}{2\pi} \frac{a}{f_a} \frac{F_{\mu\nu} \tilde{F}^{\mu\nu}}{4} + C_{ap} m_p \frac{a}{f_a} [i\bar{p}\gamma_5 p] + C_{an} m_n \frac{a}{f_a} [i\bar{n}\gamma_5 n] + C_{ae} m_e \frac{a}{f_a} [i\bar{e}\gamma_5 e] + \dots$$

2 photon	proton	neutron	electron
			

Pions, etc...

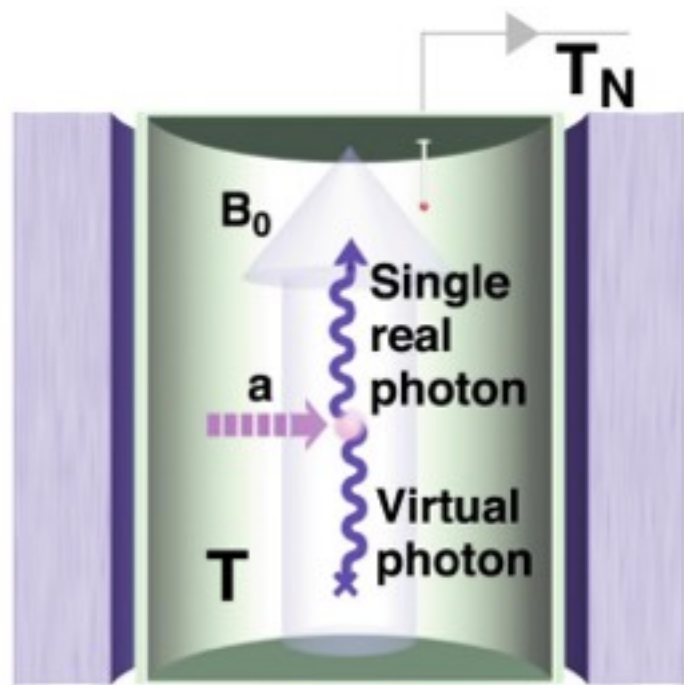
CP conserved at $\theta = 0$, but otherwise, CP violation $\propto \theta$

Neutron EDM

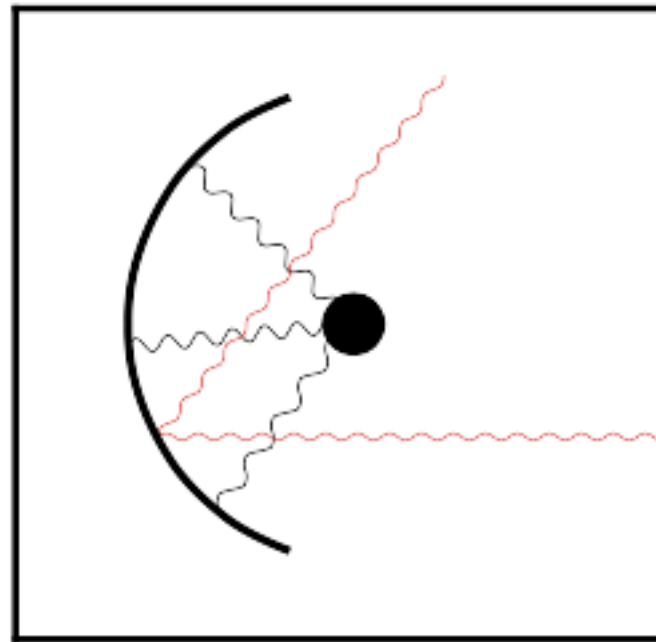


$$d_n \sim \theta \times \mathcal{O}(10^{-15}) [e \text{ cm}]$$

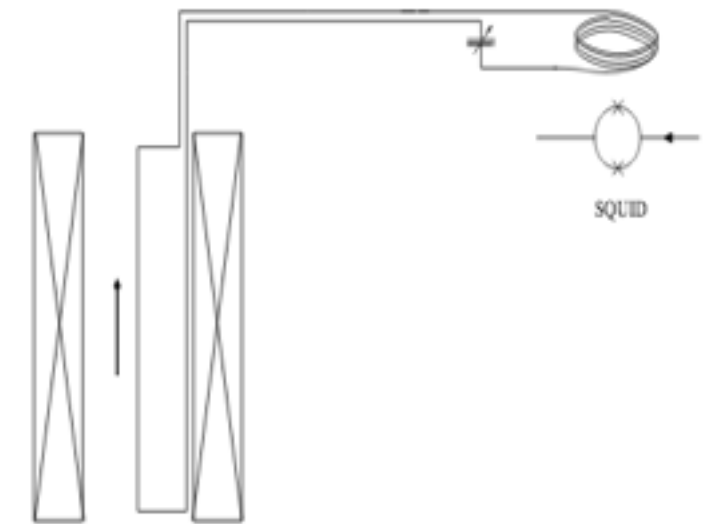
Cavities



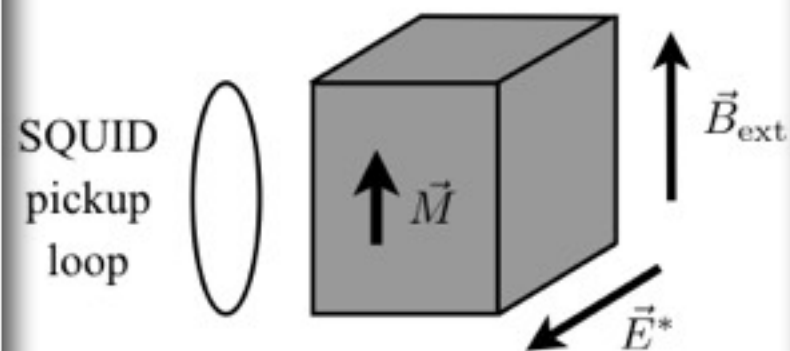
Mirrors



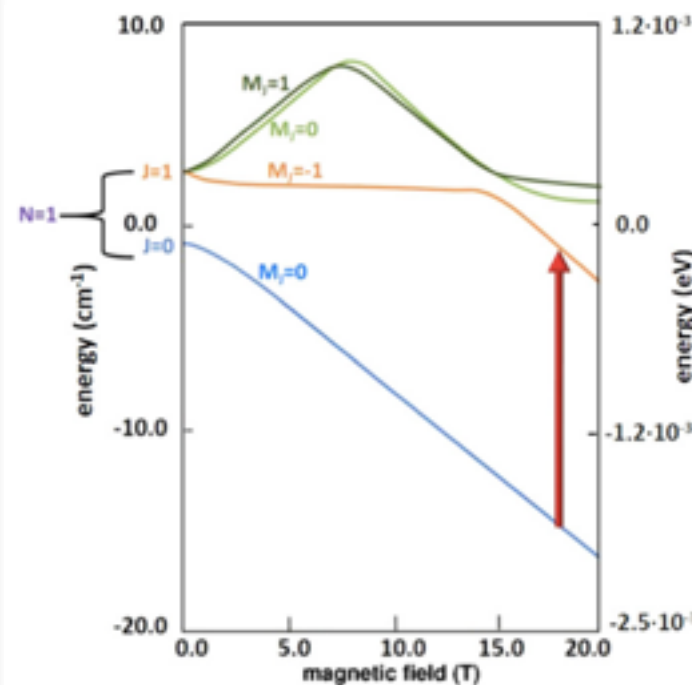
LC-circuit



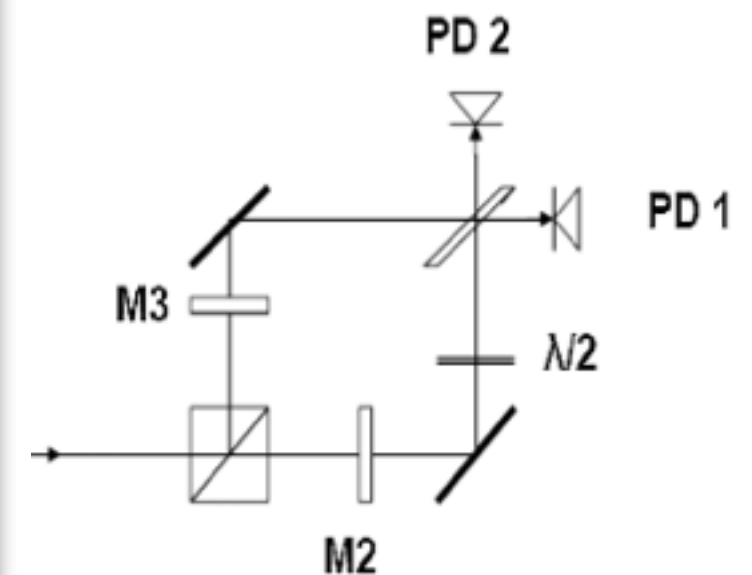
Spin precession



Atomic transitions



Optical

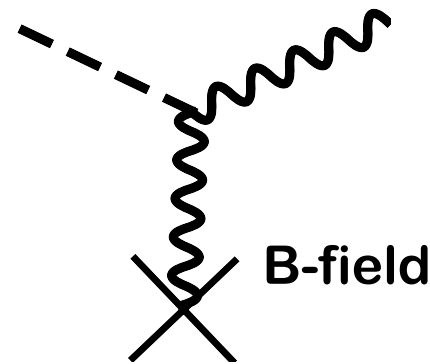


Axion DM in a B-field

$$\mathcal{L}_I = -C_{a\gamma} \frac{\alpha}{2\pi} \frac{a}{f_a} \mathbf{B} \cdot \mathbf{E}$$

- In a static magnetic field, the oscillating axion field generates EM-fields

$$\mathcal{L}_I = - \underbrace{C_{a\gamma} \frac{\alpha}{2\pi} \theta(t)}_{\text{source}} \mathbf{B}_{\text{ext}} \cdot \mathbf{E}$$



- Electric fields $\mathbf{E}_a = C_{a\gamma} \frac{\alpha \mathbf{B}_{\text{ext}}}{2\pi} \theta_0 \cos(m_a t)$ (amp independent of mass!)

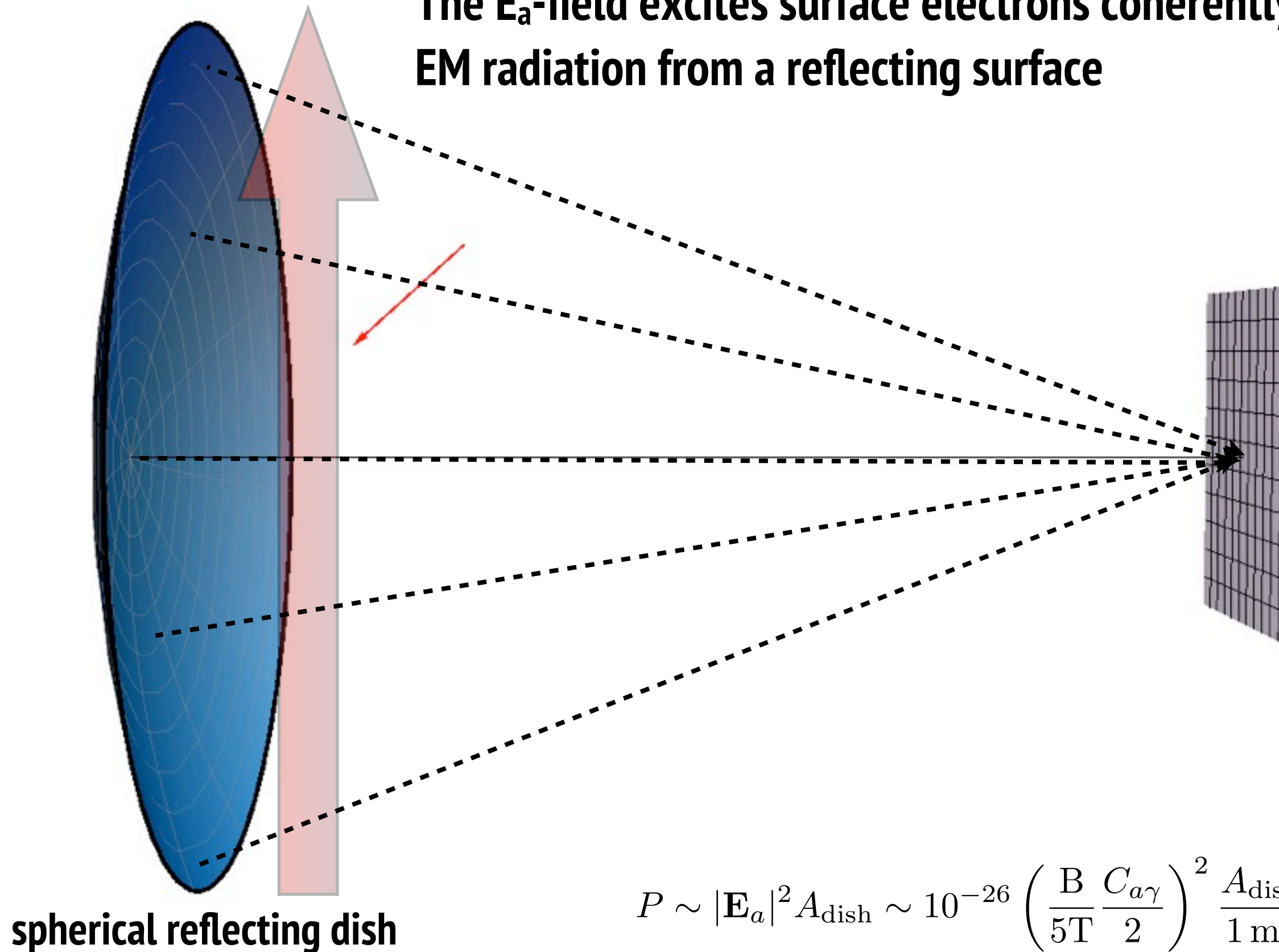
- Oscillating at a frequency $\omega \simeq m_a$

- B-fields $\propto \nabla \theta$ $|\mathbf{B}_a| \sim \langle v \rangle |\mathbf{E}_a|$

Dish antenna experiment?

Horns 2012

The E_a -field excites surface electrons coherently
EM radiation from a reflecting surface



$$P \sim |\mathbf{E}_a|^2 A_{\text{dish}} \sim 10^{-26} \left(\frac{B}{5T} \frac{C_{a\gamma}}{2} \right)^2 \frac{A_{\text{dish}}}{1 \text{ m}^2} \text{Watt}$$

Cavity experiments

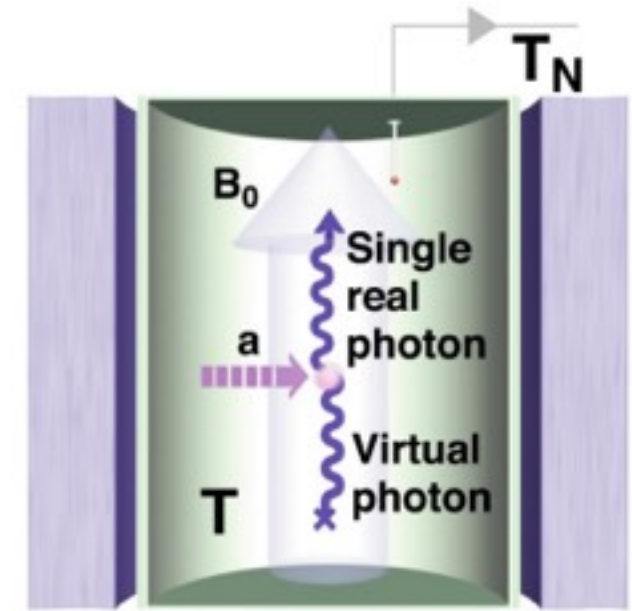
Sikivie 2013

- Haloscope (Sikivie 83)

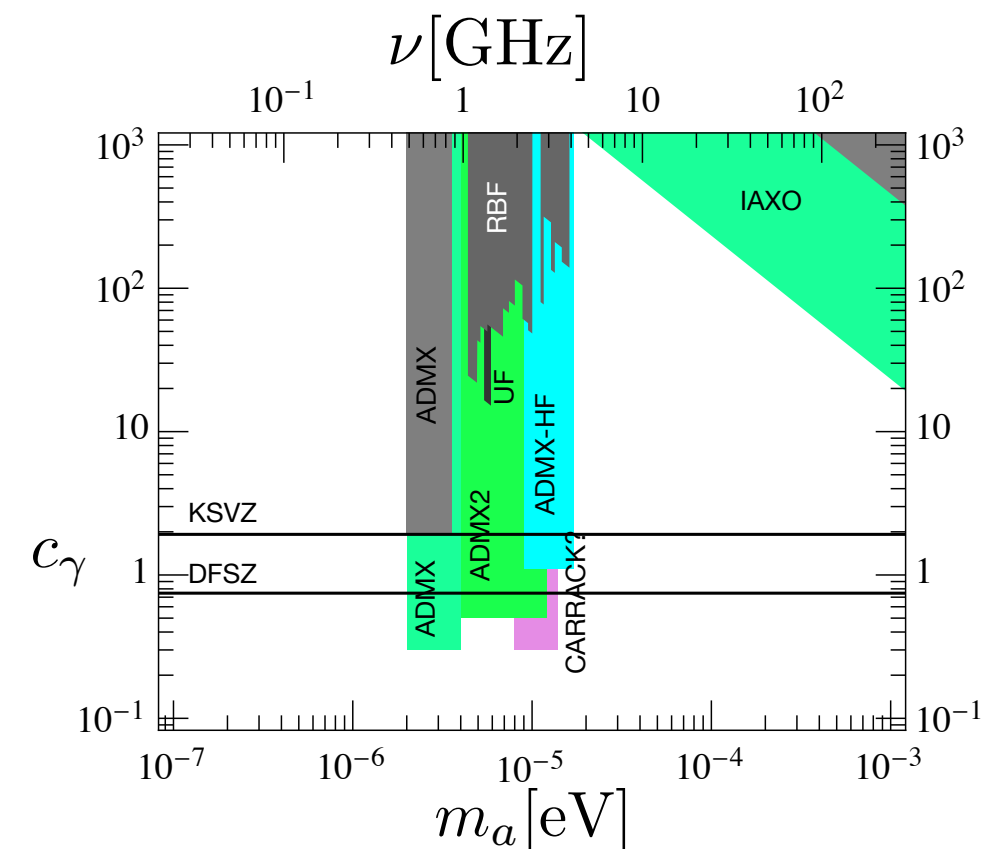
“Amplify resonantly the EM field in a cavity”

$$P \sim Q |\mathbf{E}_a|^2 (V m_a) \mathcal{G} \kappa \quad (\text{on resonance})$$

(integrate the power in a coherent time)



- Past experiments Florida U., RBF, ADMX, CARRACK
- Future endeavors: ADMX, ADMX-HF, YMCE, CAPP



Cavity experiments

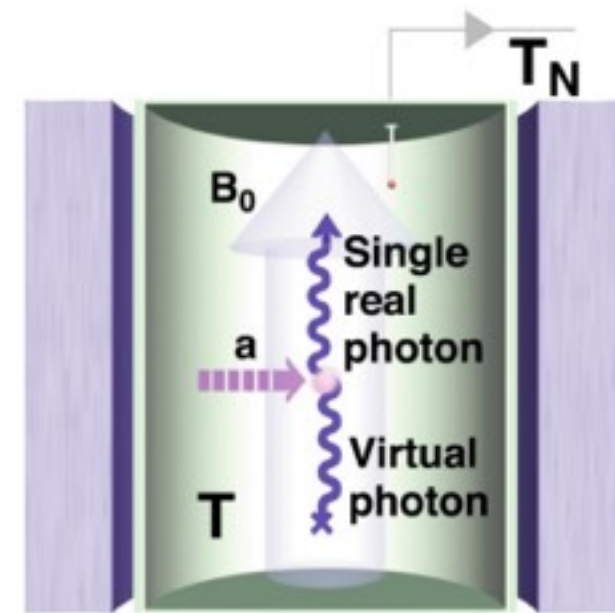
Sikivie 2013

- Haloscope (Sikivie 83)

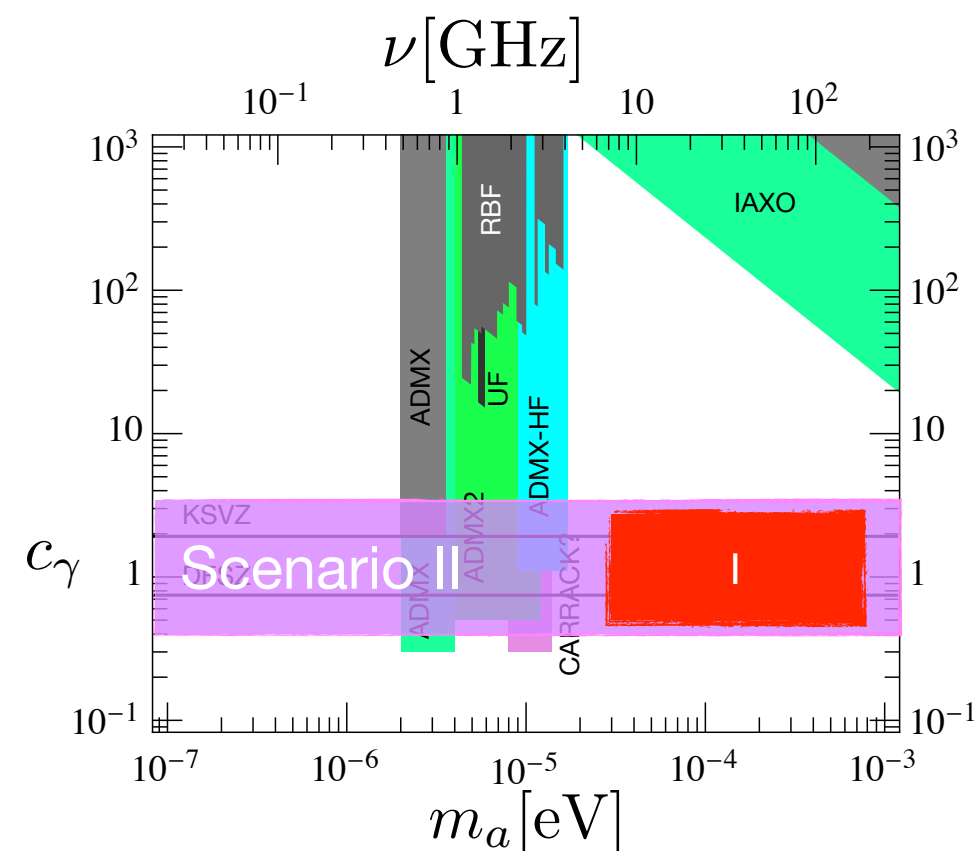
“Amplify resonantly the EM field in a cavity”

$$P \sim Q |\mathbf{E}_a|^2 (V m_a) \mathcal{G} \kappa \quad (\text{on resonance})$$

(integrate the power in a coherent time)



- Past experiments Florida U., RBF, ADMX, CARRACK
- Future endeavors: ADMX, ADMX-HF, YMCE, CAPP
- Parameters unexplored at low and high masses: WHY?



Cavity experiments

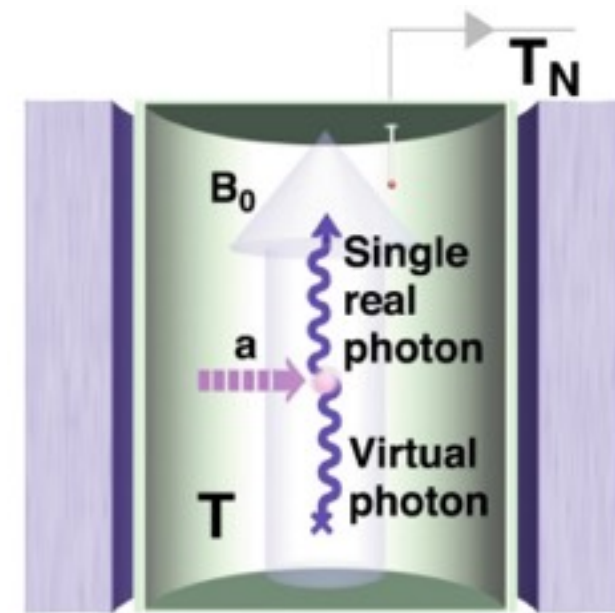
Sikivie 2013

- Haloscope (Sikivie 83)

“Amplify resonantly the EM field in a cavity”

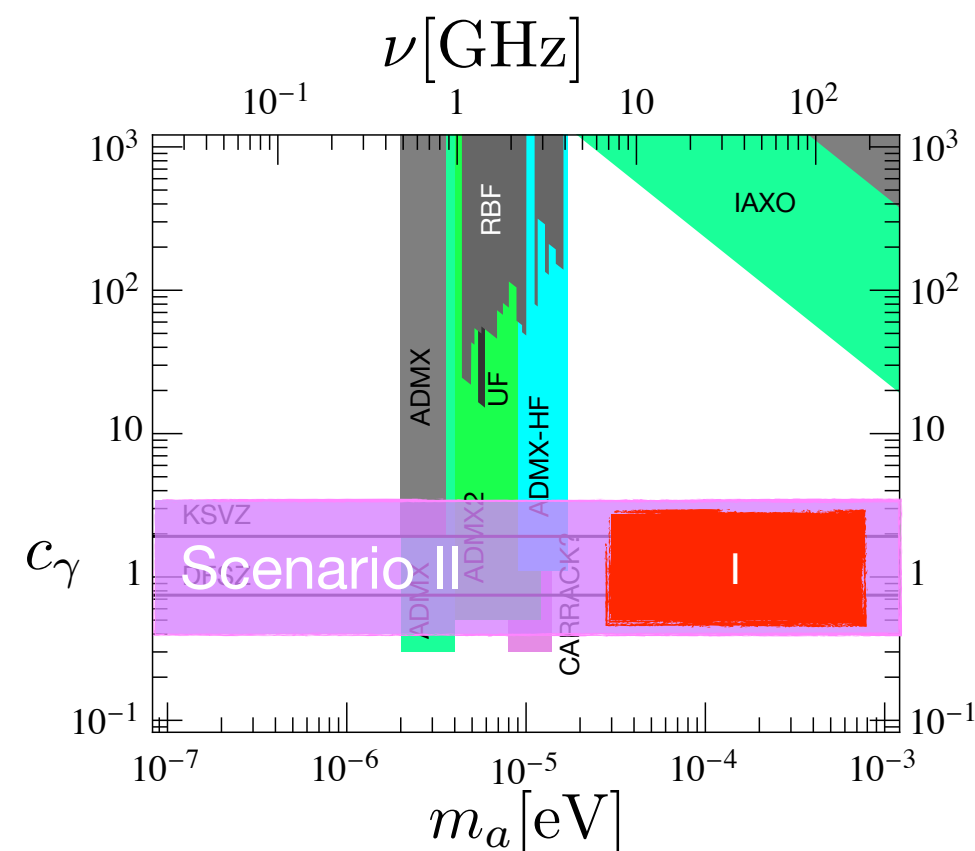
$$P \sim Q |\mathbf{E}_a|^2 (V m_a) \mathcal{G} \kappa \quad (\text{on resonance})$$

(integrate the power in a coherent time)



- Past experiments Florida U., RBF, ADMX, CARRACK
- Future endeavors: ADMX, ADMX-HF, YMCE, CAPP
- Parameters unexplored at low and high masses: WHY?

Cylindrical cavity ($h/r=b$) like ADMX but scaled



Cavity experiments

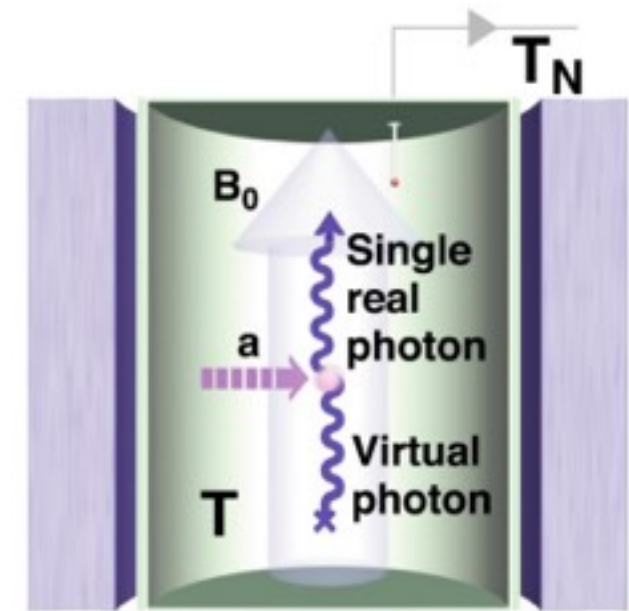
Sikivie 2013

- Haloscope (Sikivie 83)

“Amplify resonantly the EM field in a cavity”

$$P \sim Q |\mathbf{E}_a|^2 (V m_a) \mathcal{G} \kappa \quad (\text{on resonance})$$

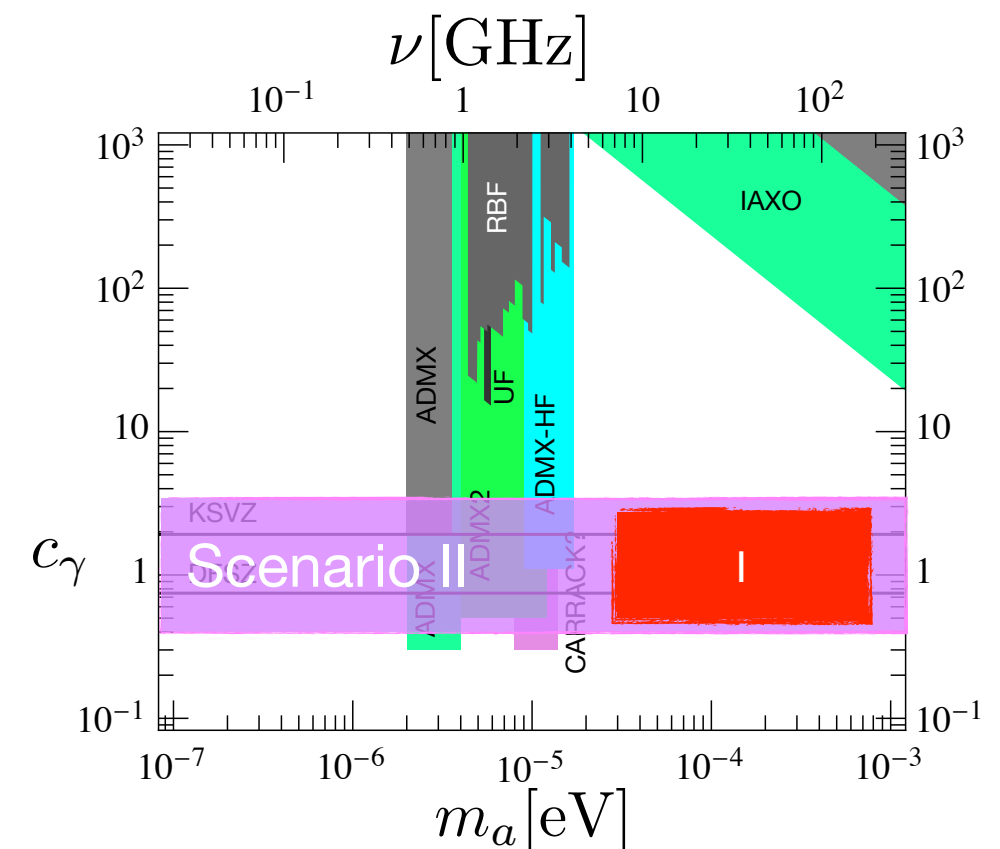
(integrate the power in a coherent time)



- Past experiments Florida U., RBF, ADMX, CARRACK
- Future endeavors: ADMX, ADMX-HF, YMCE, CAPP
- Parameters unexplored at low and high masses: WHY?

Cylindrical cavity ($h/r=b$) like ADMX but scaled

- Signal $(V \propto m_a^{-3}) \quad P_{\text{out}} \propto V m_a \sim \frac{1}{m_a^2}$



Cavity experiments

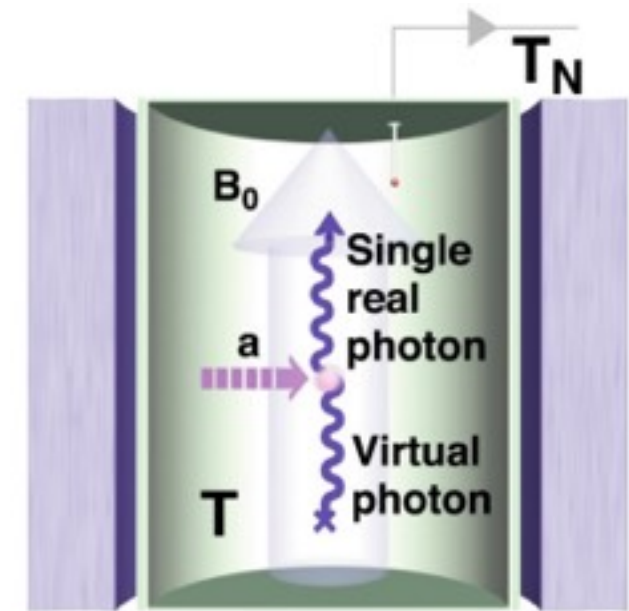
Sikivie 2013

- Haloscope (Sikivie 83)

“Amplify resonantly the EM field in a cavity”

$$P \sim Q |\mathbf{E}_a|^2 (V m_a) \mathcal{G} \kappa \quad (\text{on resonance})$$

(integrate the power in a coherent time)

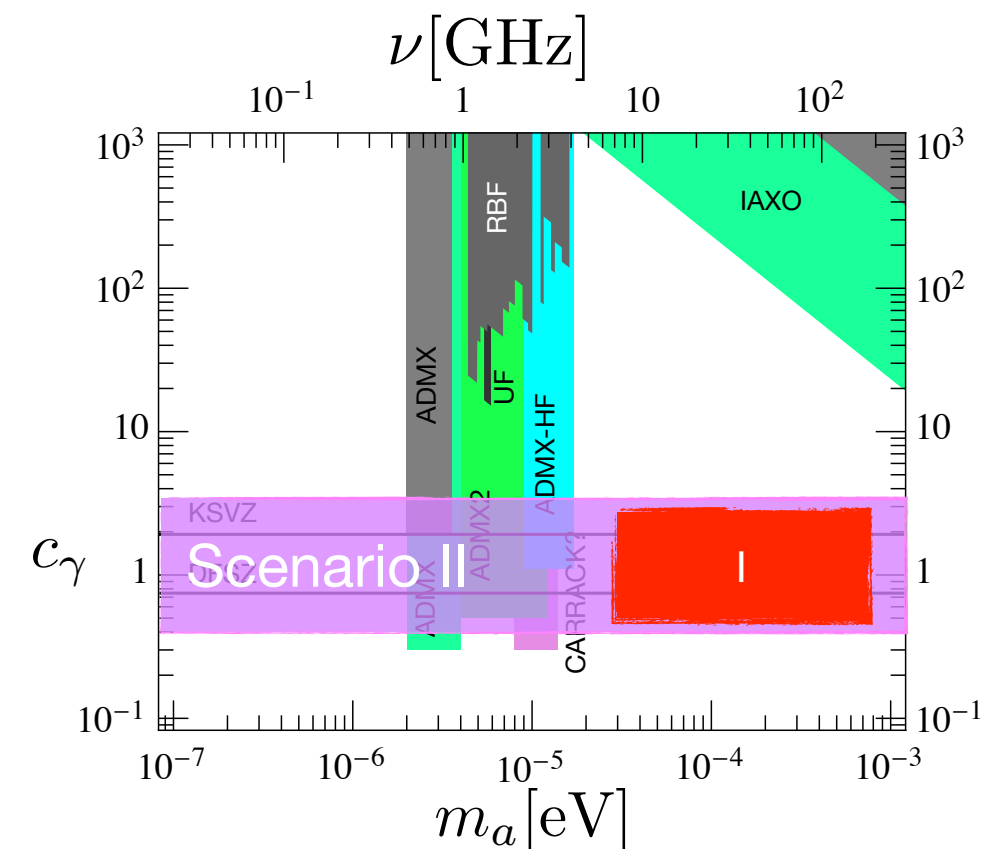


- Past experiments Florida U., RBF, ADMX, CARRACK
- Future endeavors: ADMX, ADMX-HF, YMCE, CAPP
- Parameters unexplored at low and high masses: WHY?

Cylindrical cavity ($h/r=b$) like ADMX but scaled

- Signal $(V \propto m_a^{-3}) \quad P_{\text{out}} \propto V m_a \sim \frac{1}{m_a^2}$

- Noise $P_{\text{noise}} = T_{\text{sys}} \Delta \nu_a \propto m_a^2$

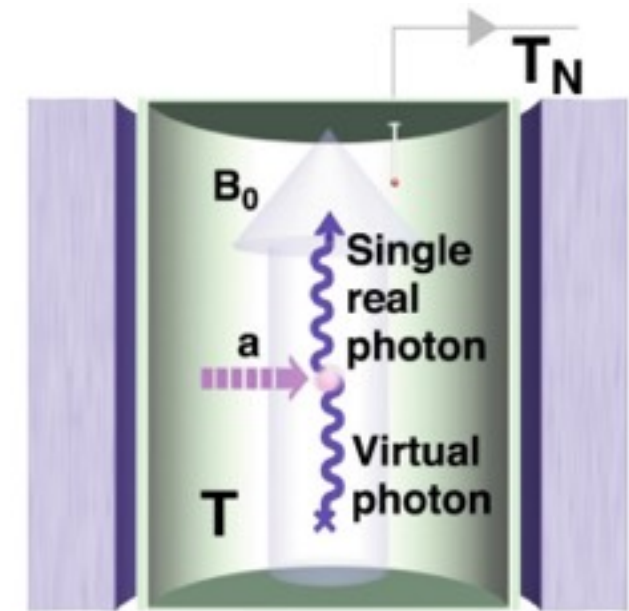


- Haloscope (Sikivie 83)

“Amplify resonantly the EM field in a cavity”

$$P \sim Q |\mathbf{E}_a|^2 (V m_a) \mathcal{G} \kappa \quad (\text{on resonance})$$

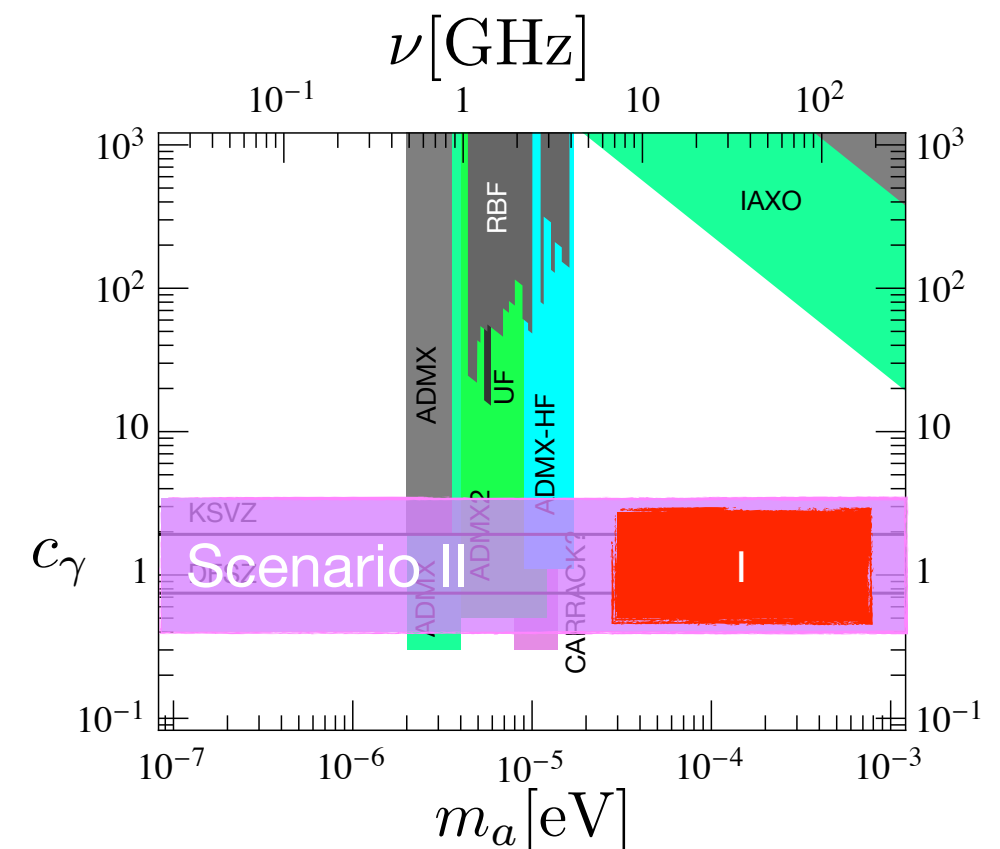
(integrate the power in a coherent time)



- Past experiments Florida U., RBF, ADMX, CARRACK
- Future endeavors: ADMX, ADMX-HF, YMCE, CAPP
- Parameters unexplored at low and high masses: WHY?

Cylindrical cavity ($h/r=b$) like ADMX but scaled

- Signal $(V \propto m_a^{-3}) \quad P_{\text{out}} \propto V m_a \sim \frac{1}{m_a^2}$
- Noise $P_{\text{noise}} = T_{\text{sys}} \Delta \nu_a \propto m_a^2$
- Signal/noise in $\Delta \nu_a$ of time, t , $\frac{S}{N} = \frac{P_{\text{out}}}{P_{\text{noise}}} \sqrt{\Delta \nu_a t}$



Cavity experiments

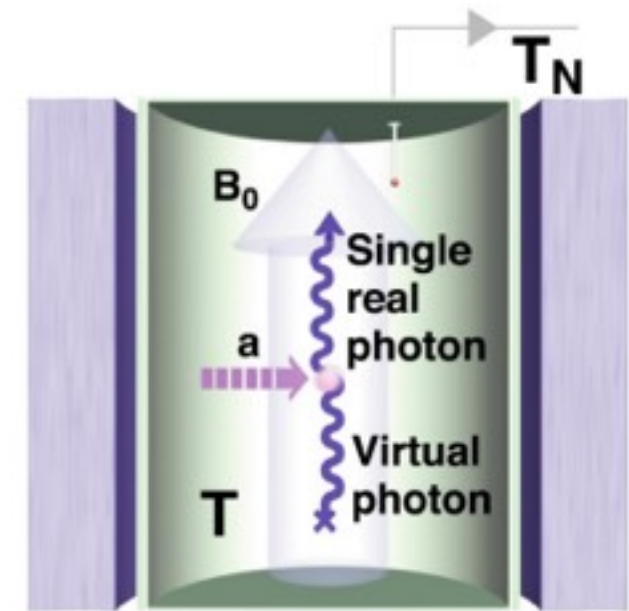
Sikivie 2013

- Haloscope (Sikivie 83)

“Amplify resonantly the EM field in a cavity”

$$P \sim Q |\mathbf{E}_a|^2 (V m_a) \mathcal{G} \kappa \quad (\text{on resonance})$$

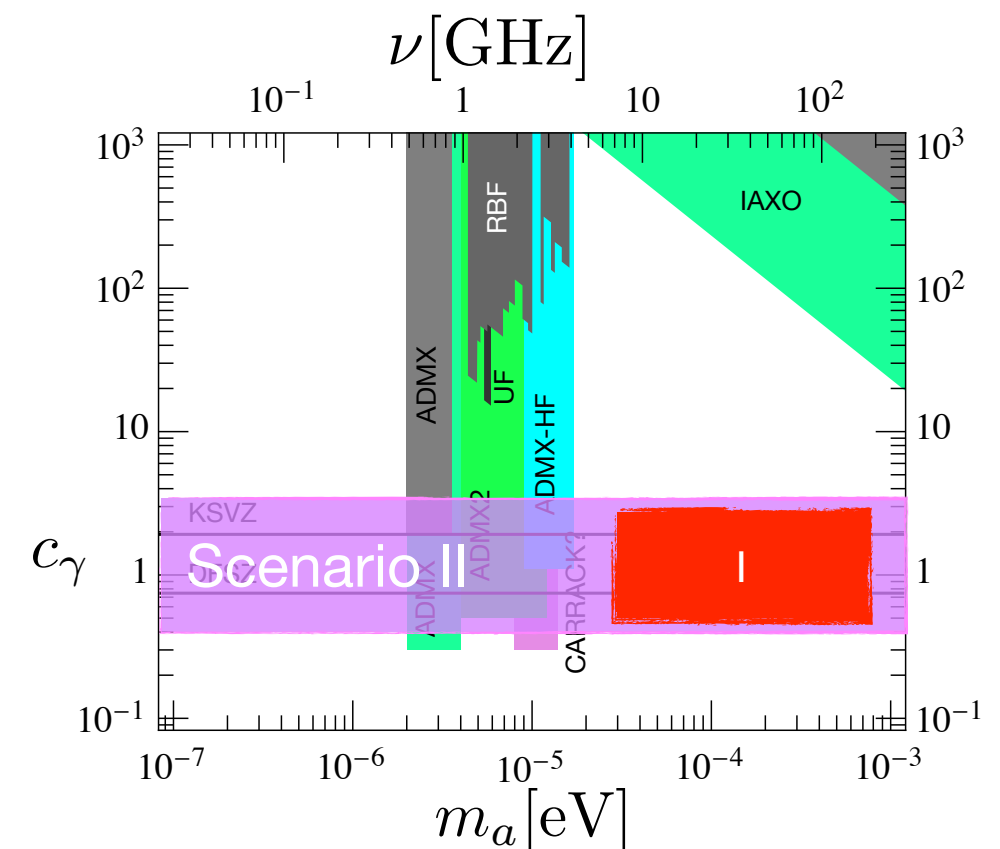
(integrate the power in a coherent time)



- Past experiments Florida U., RBF, ADMX, CARRACK
- Future endeavors: ADMX, ADMX-HF, YMCE, CAPP
- Parameters unexplored at low and high masses: WHY?

Cylindrical cavity ($h/r=b$) like ADMX but scaled

- Signal $(V \propto m_a^{-3}) \quad P_{\text{out}} \propto V m_a \sim \frac{1}{m_a^2}$
- Noise $P_{\text{noise}} = T_{\text{sys}} \Delta \nu_a \propto m_a^2$
- Signal/noise in $\Delta \nu_a$ of time, t , $\frac{S}{N} = \frac{P_{\text{out}}}{P_{\text{noise}}} \sqrt{\Delta \nu_a t}$
- Scanning rate $\frac{1}{m_a} \frac{d\Delta m_a}{dt} \propto \frac{C_{a\gamma}^4}{m_a^9}$



Cavity experiments

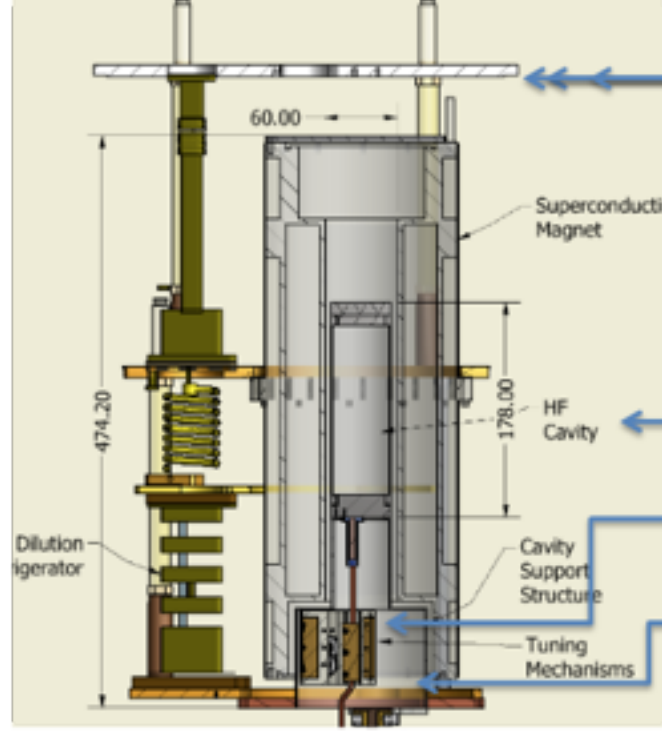
ADMX



CARRACK (discontinued)



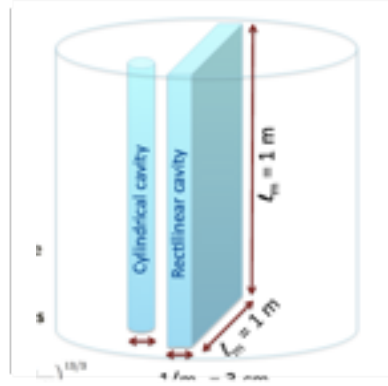
CULTASK - CAPP - Korea



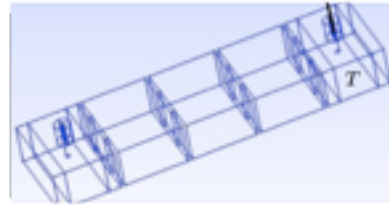
ADMX-HF



ADMX-Fermilab

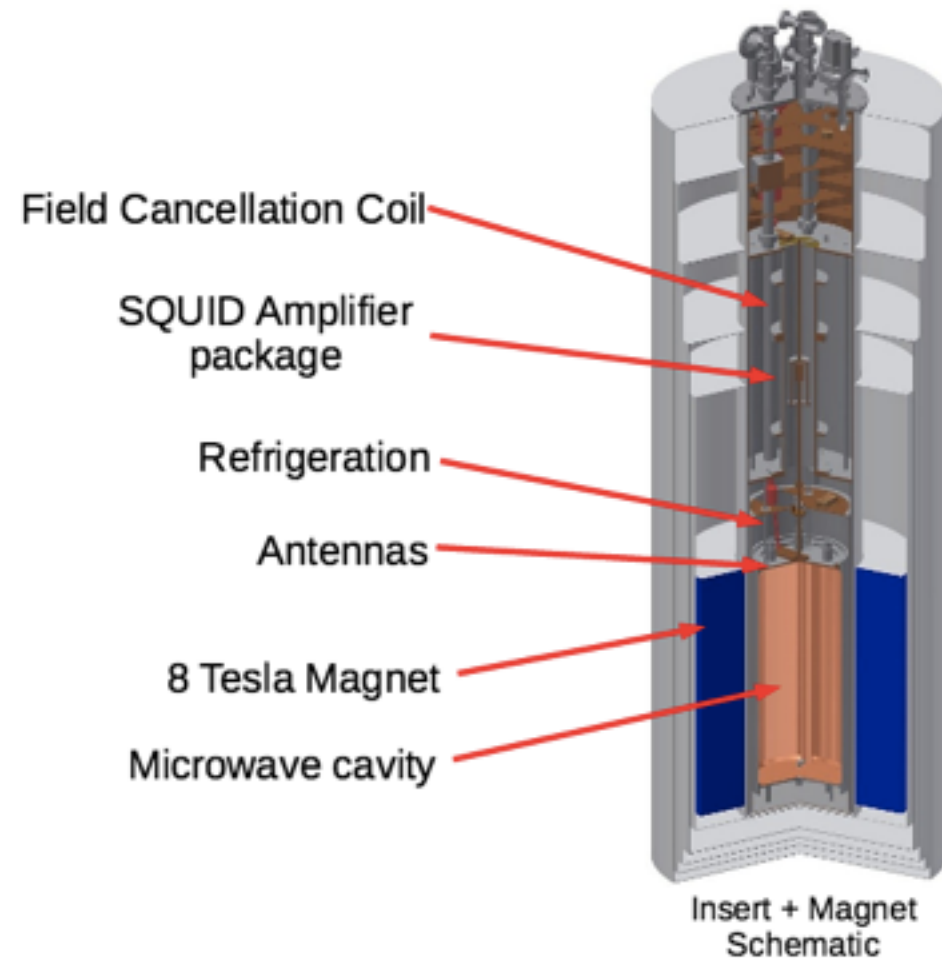


RADES



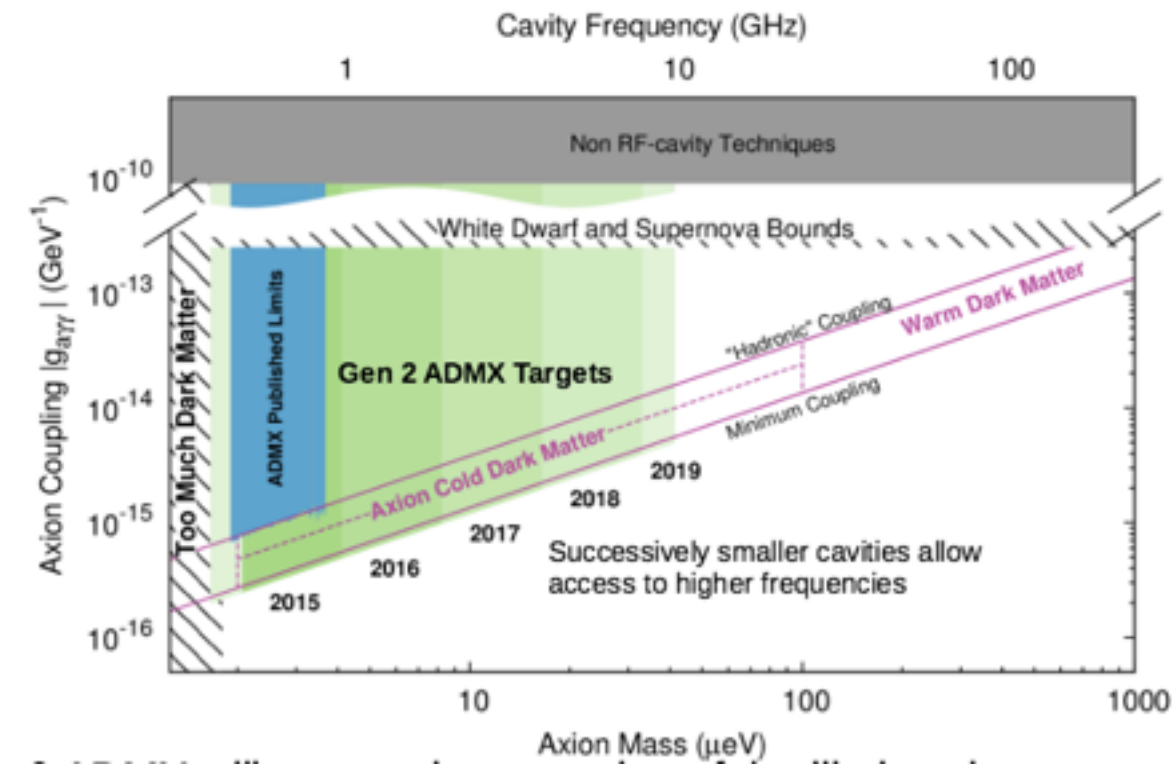
CAST-CAPP



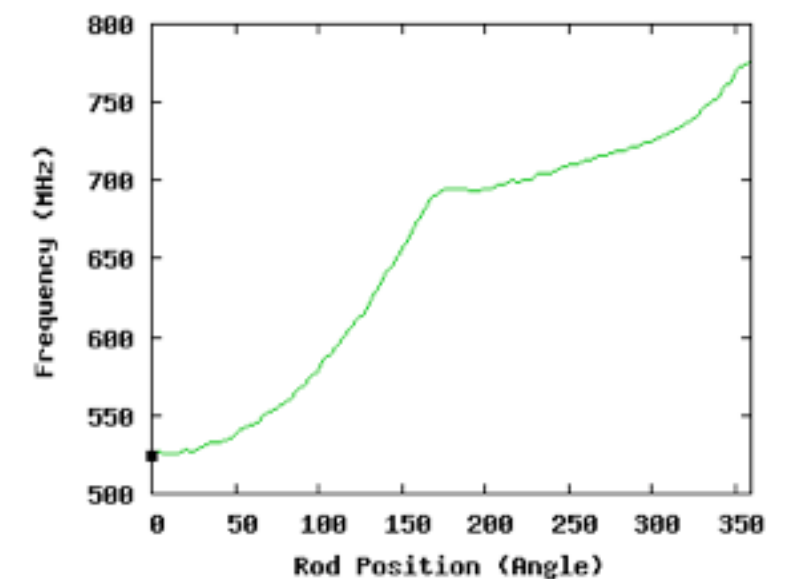
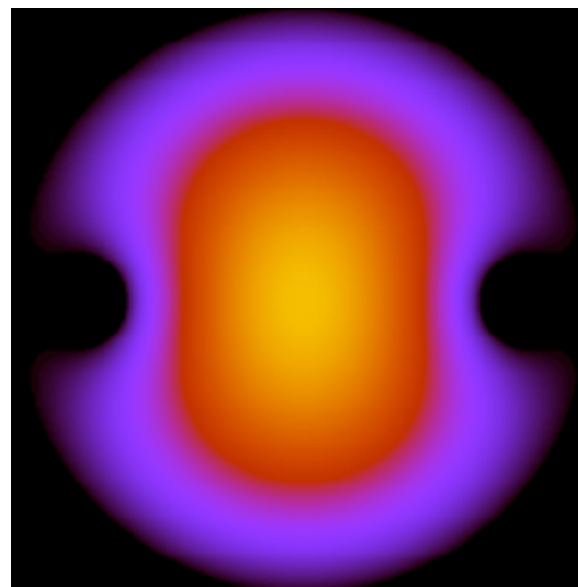
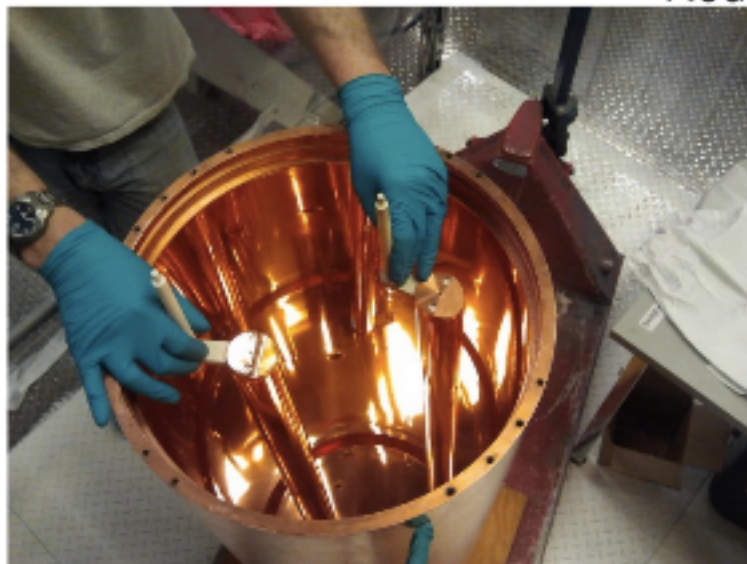


Insert extraction from magnet

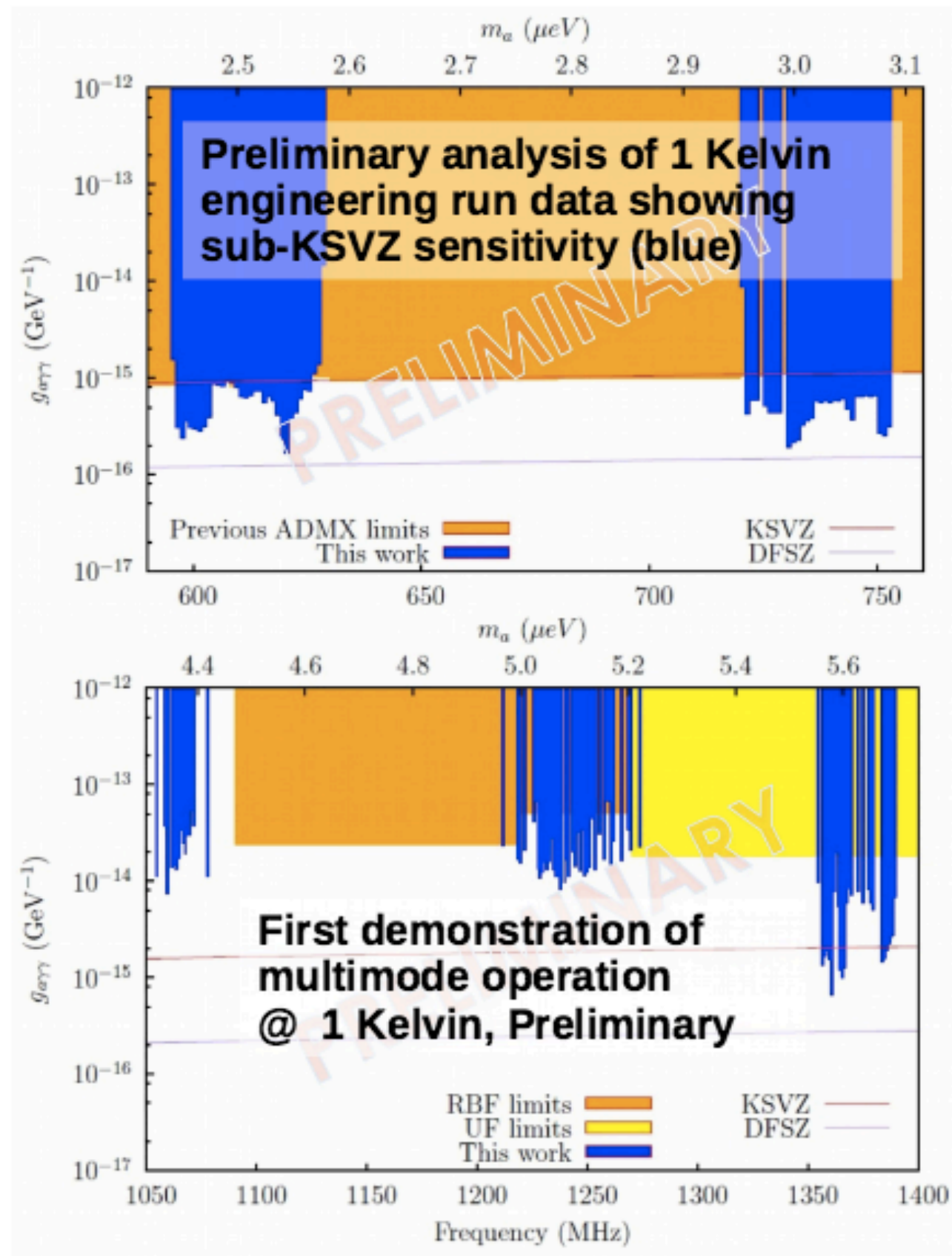
Goals



Scanning over frequencies

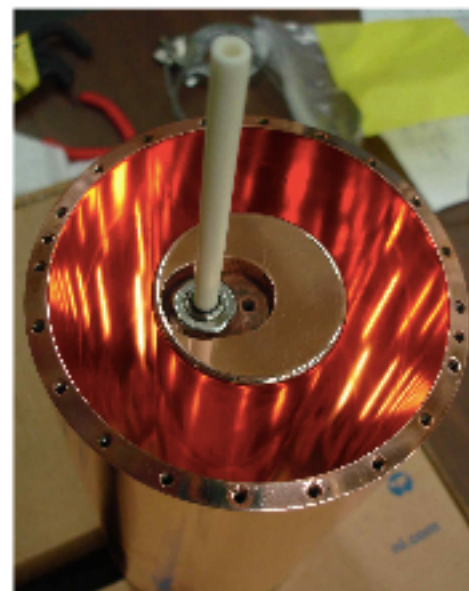


ADMX 2015



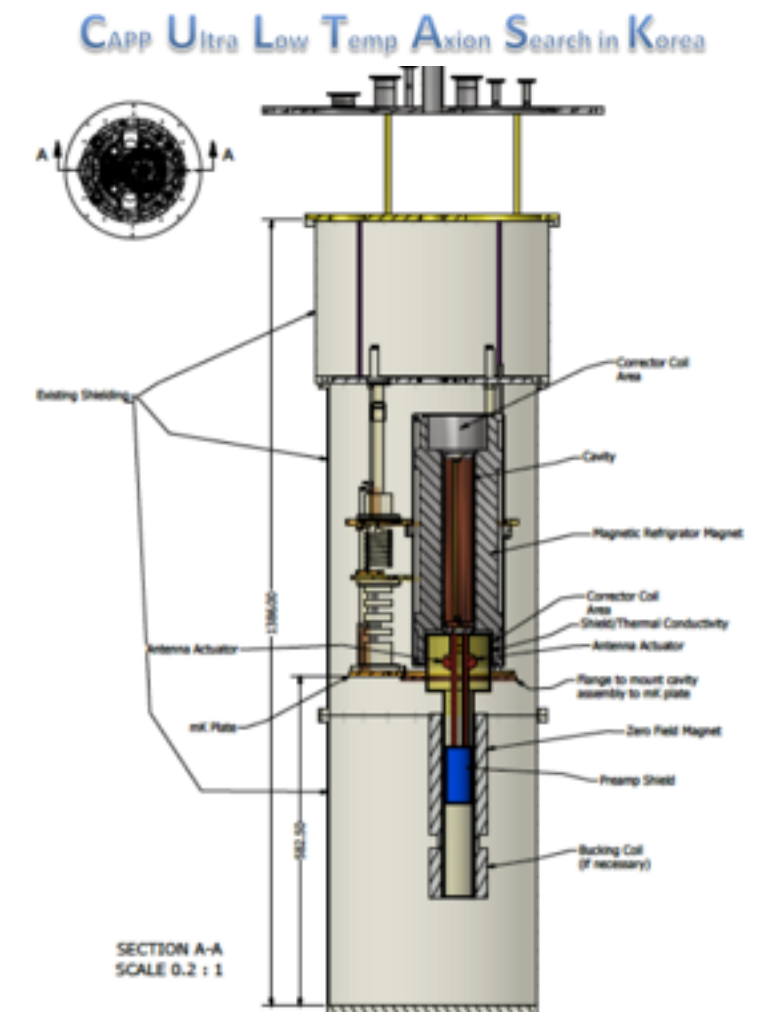
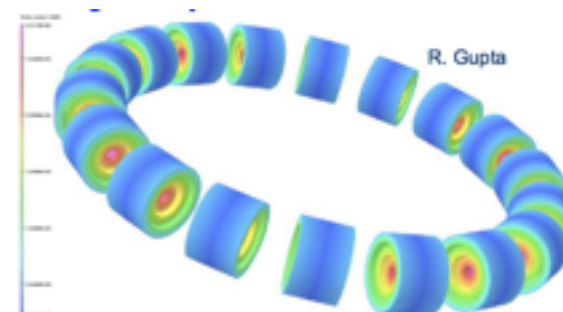
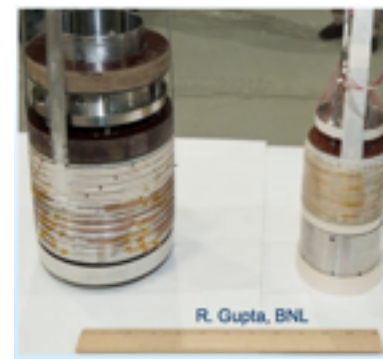
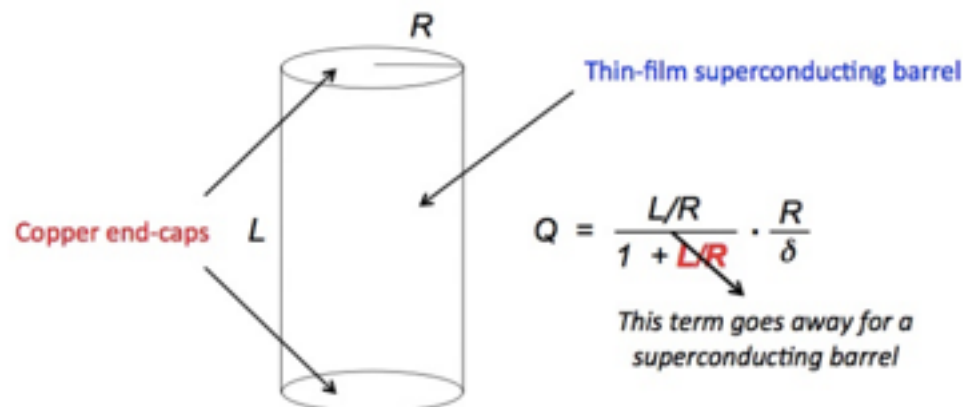
- Test bench for higher frequencies
- Hybrid Superconducting cavities
- Josephson parametric amps
- Single photon detectors at > 10 GHz

Set up running at $f \sim 5.6$ GHz, 25 mK temperatures, 9T

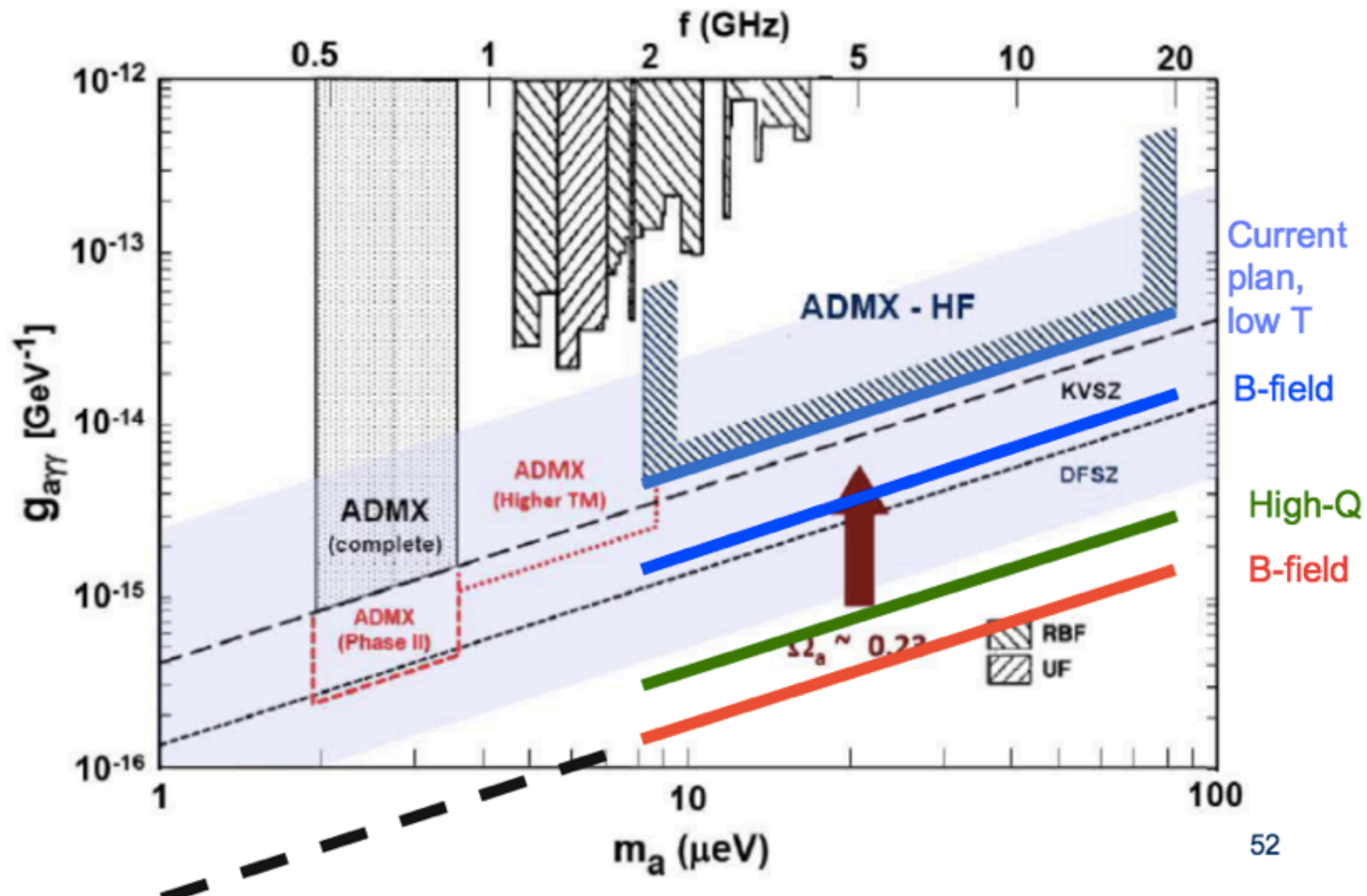


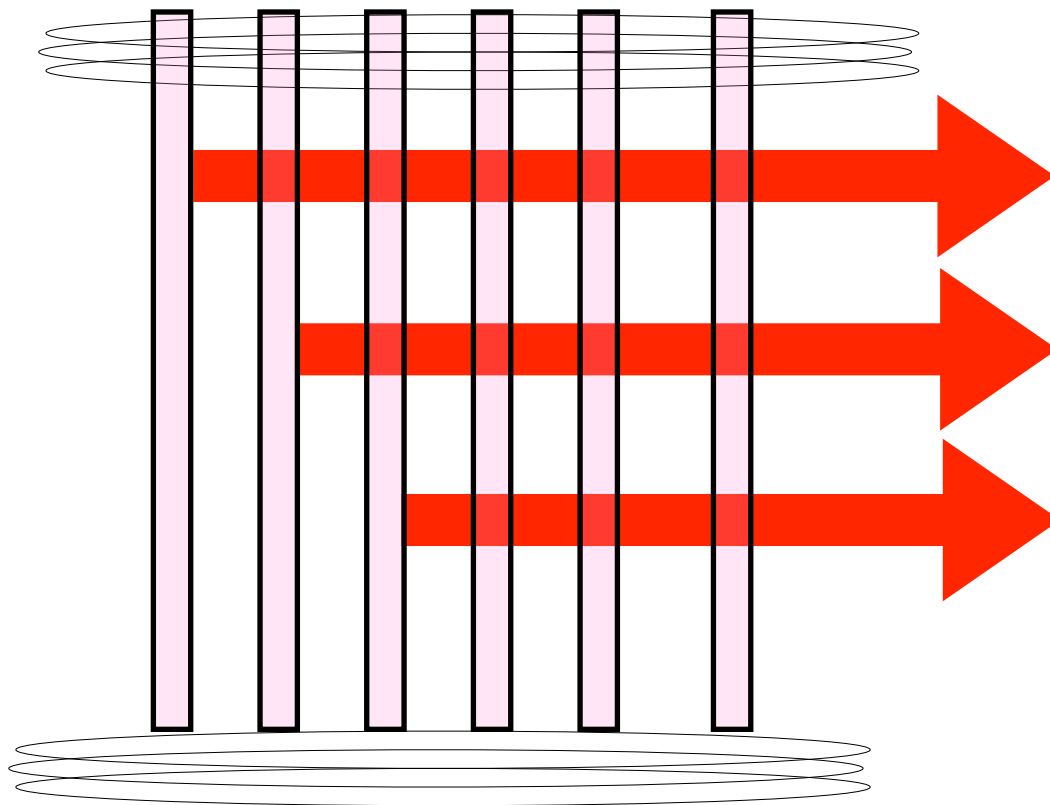
- Korean IBS Center for Axion and Precision Physics
- + proton EDM experiment (with Fermilab)
- + Axion DM experiments
 - * CULTASK at ~ 6-10 GHz?
 - * Lots of R+D
- High Tc superconducting magnets (25,35 T !)
- hybrid resonant cavities, toroidal?
- SQUIDS at HF
- Add up cavity outputs

	2015	2016	2017	2018
Lab Space	Munji Campus Design & Renovation	Occupation		Occupation
Magnet	Prototype, testing of SuperC cables	25T, 10cm bore SuperC Magnet design	Work on 35T, 10 cm bore SuperC magnet	Magnet Delivery
Cavity Development	Procure Equip. Study res. and geom.	Development of high Q SC resonator	Production of high Q resonator	
Amplifier	Design and production of prototype SQUID for 1-10 GHz Acquire JPA and test		SQUID delivery from KRISS Develop higher freq. amplifier	
Axion Cavity Experiment	Building infrastructure. Engineering Run at KAIST	Experimental Setup at Munji Test Runs		High Field Magnet + SQUID + SC Cavity



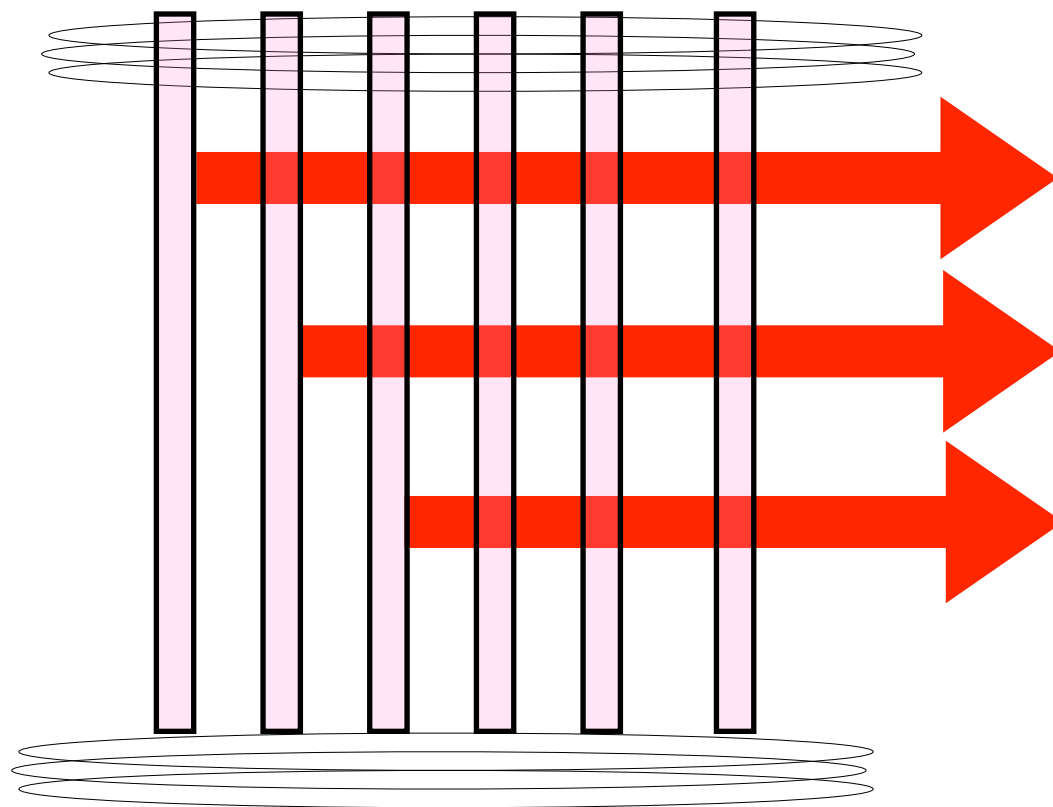
ADMX goals and CAPP plan





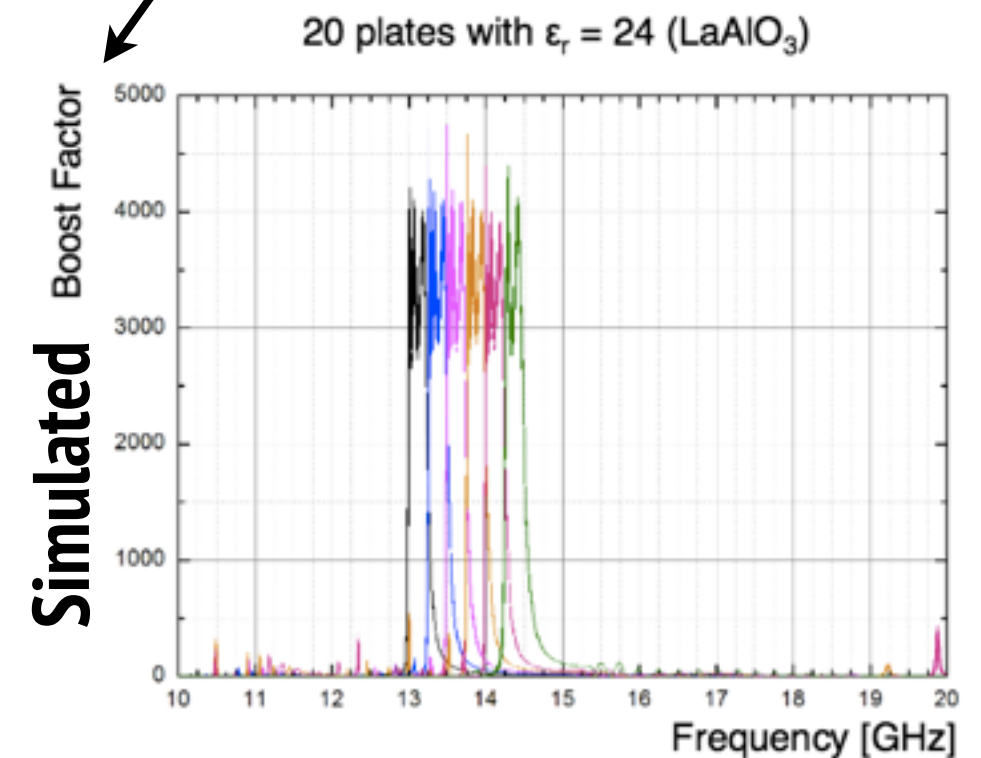
**Emitted EM-waves from each interface
+ internal reflections ...**

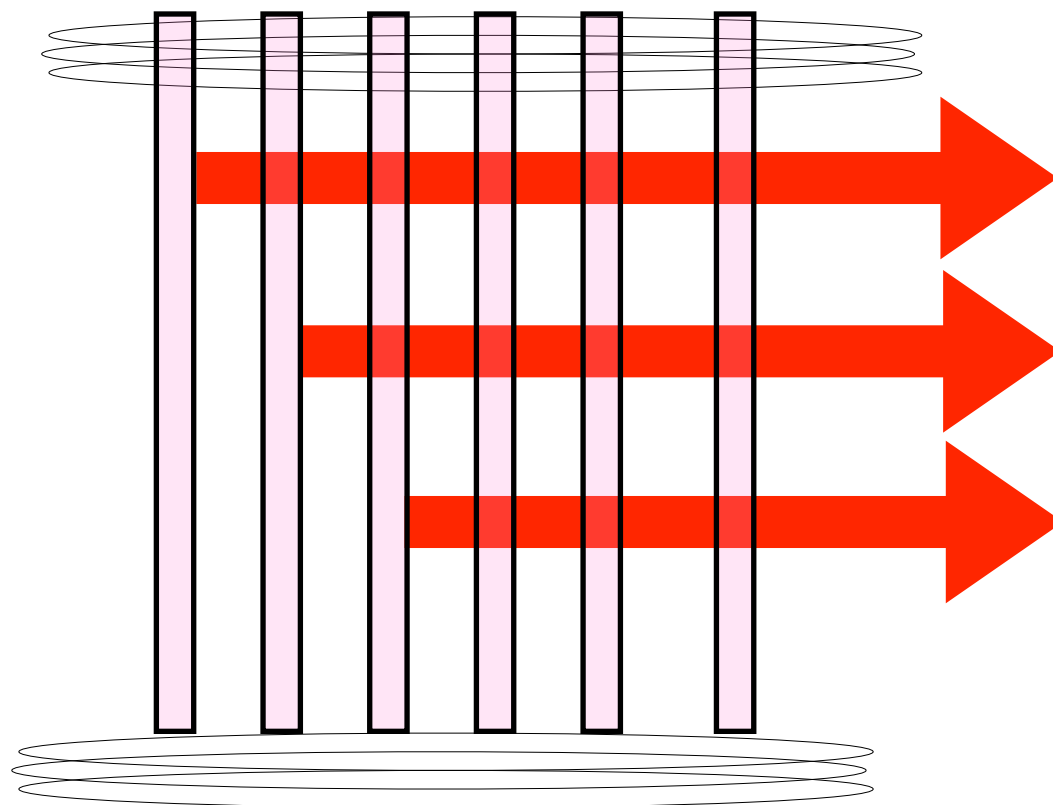
$$P \sim |\mathbf{E}_a|^2 \text{Area} \times \mathcal{O}(N^2)$$



**Emitted EM-waves from each interface
+ internal reflections ...**

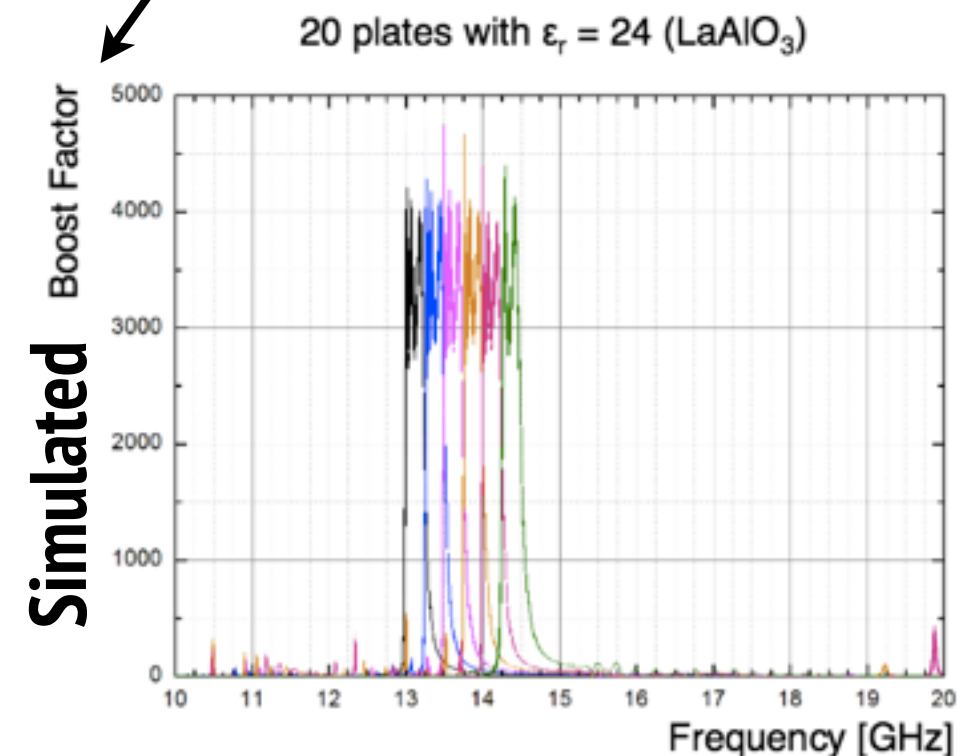
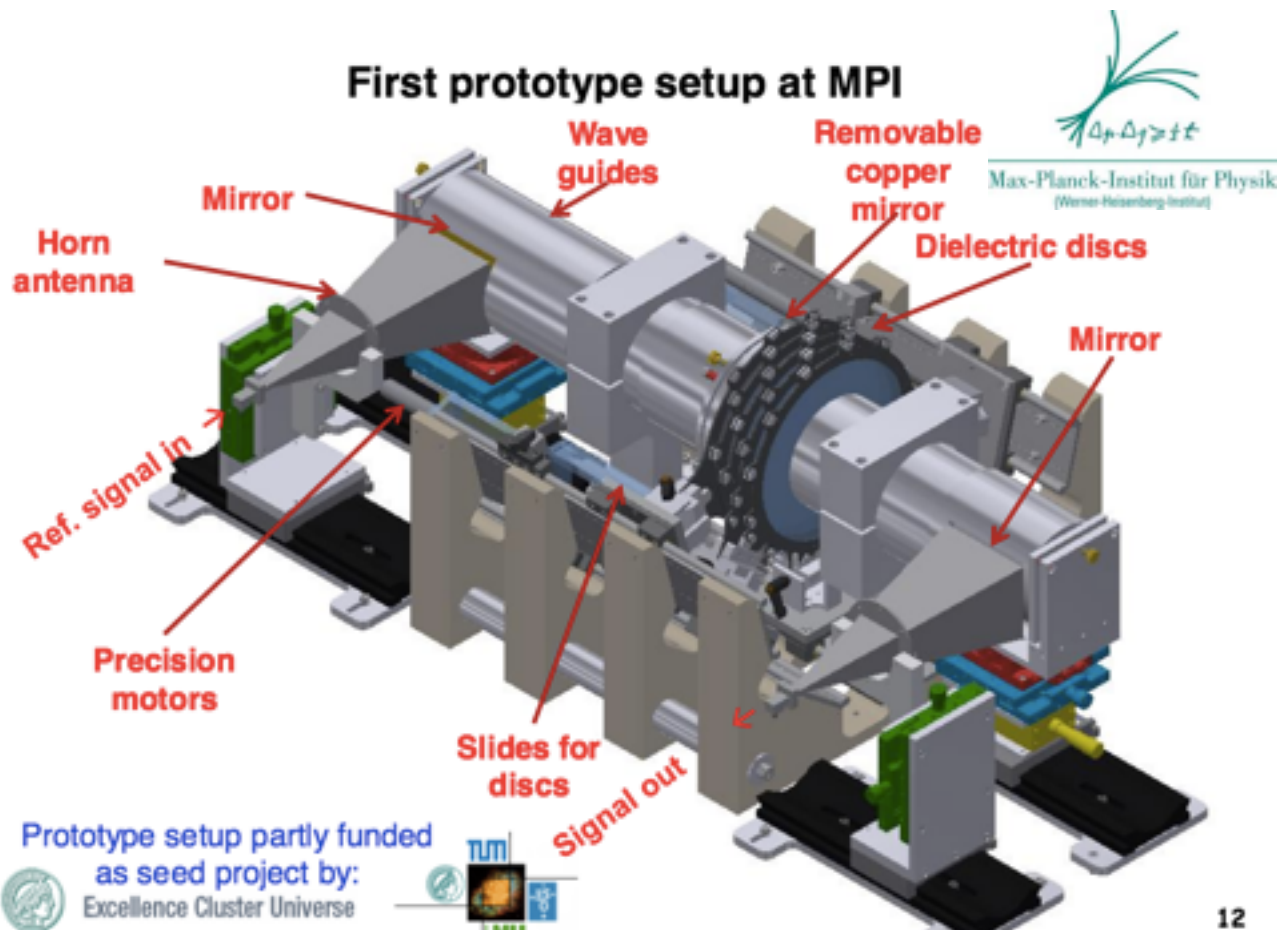
$$P \sim |\mathbf{E}_a|^2 \text{Area} \times \mathcal{O}(N^2)$$



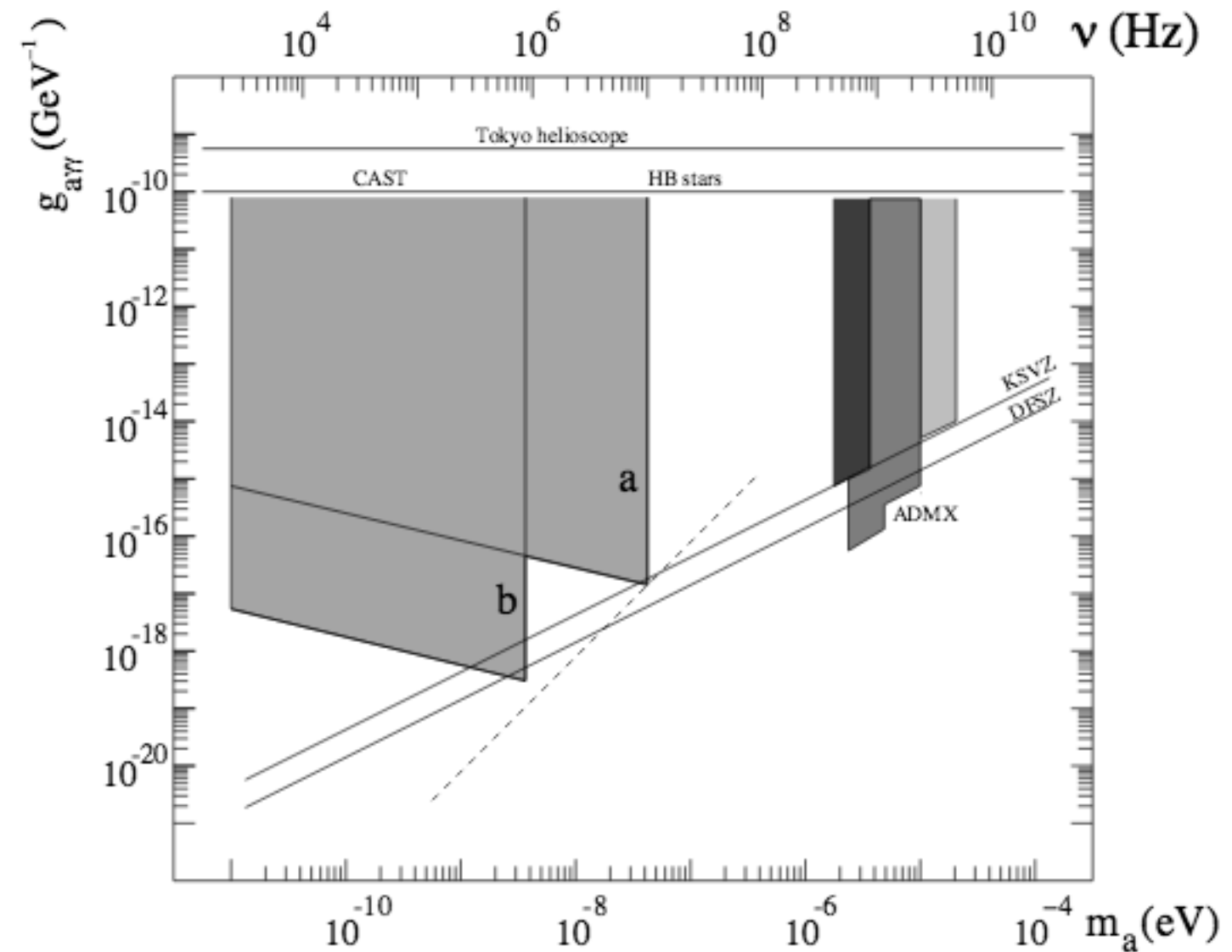
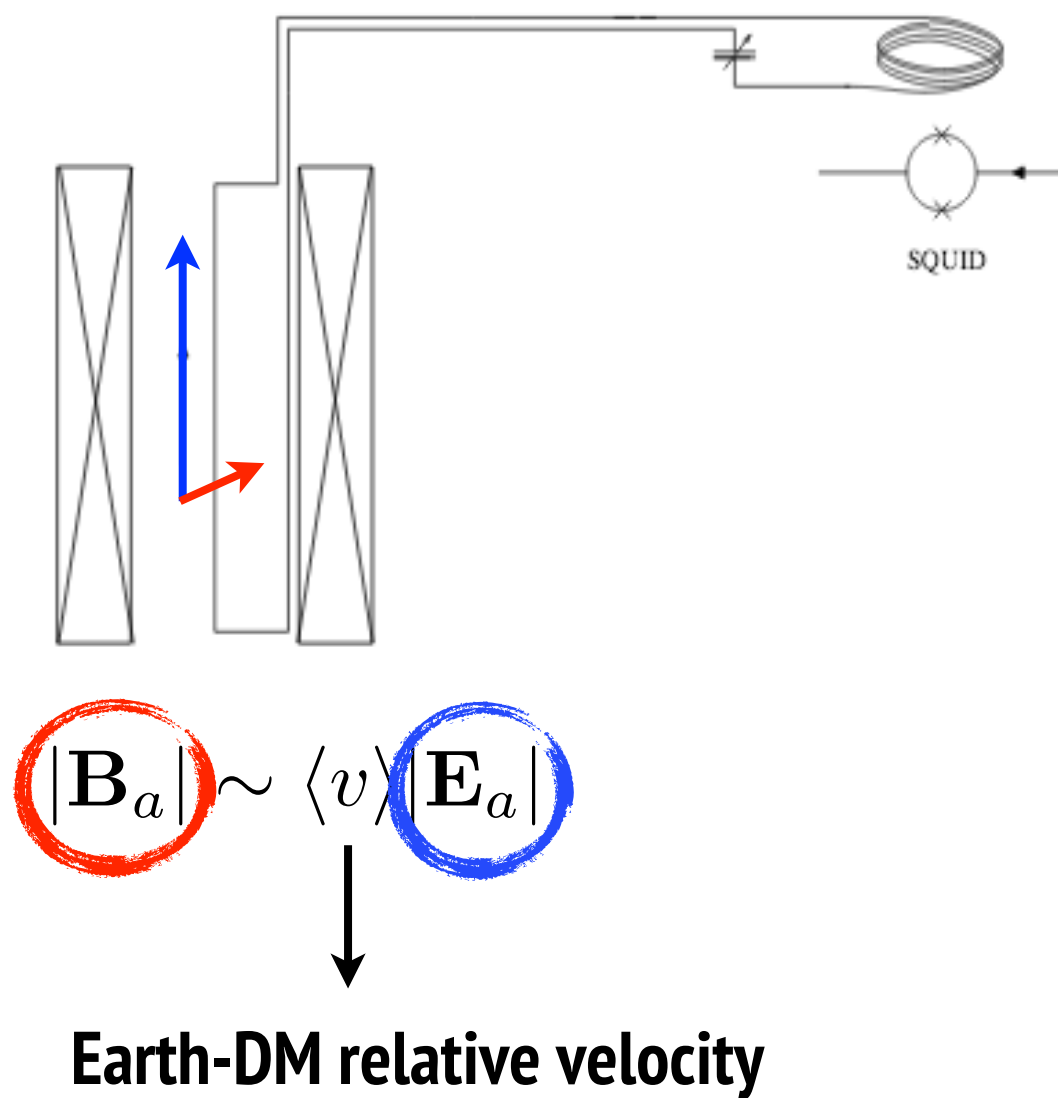


Emitted EM-waves from each interface
+ internal reflections ...

$$P \sim |\mathbf{E}_a|^2 \text{Area} \times \mathcal{O}(N^2)$$

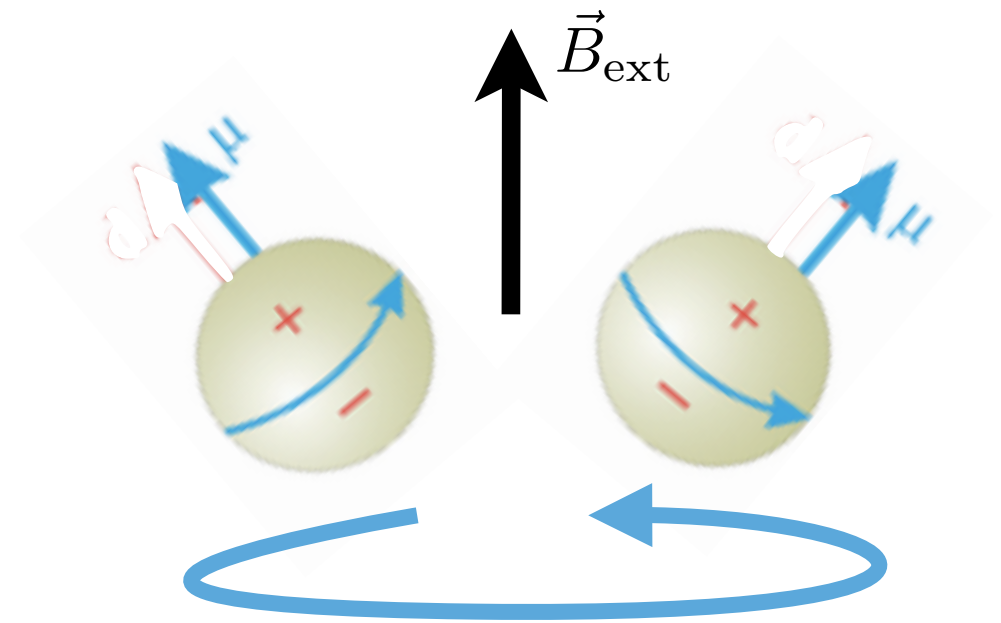
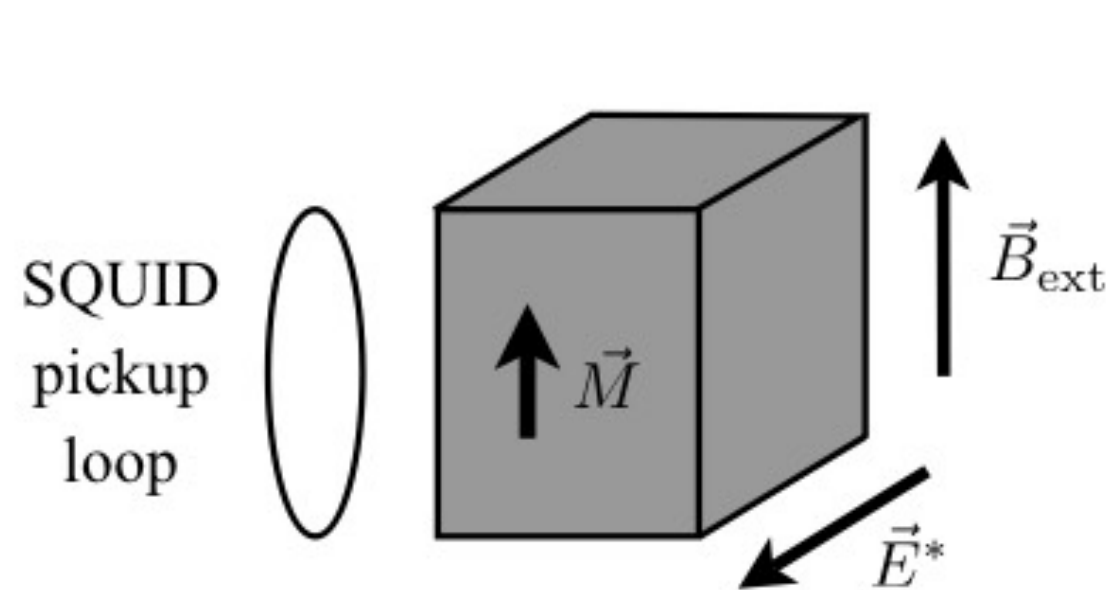


- Detect low-frequency B-field with a tunable LC
- First moves in Florida U.



CASPER : Spin precession

Graham 2012

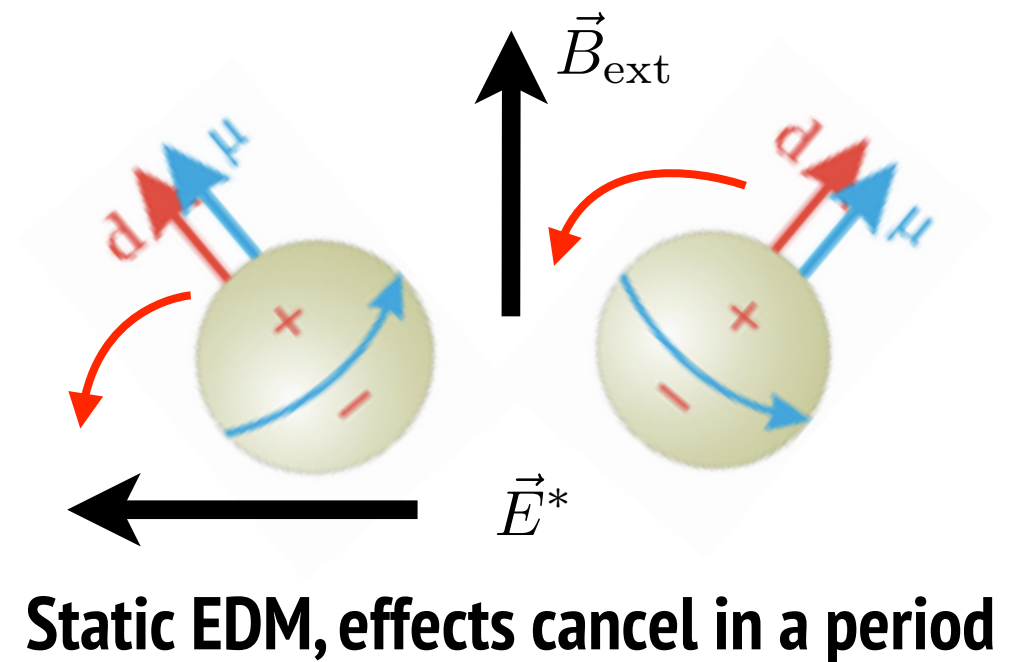
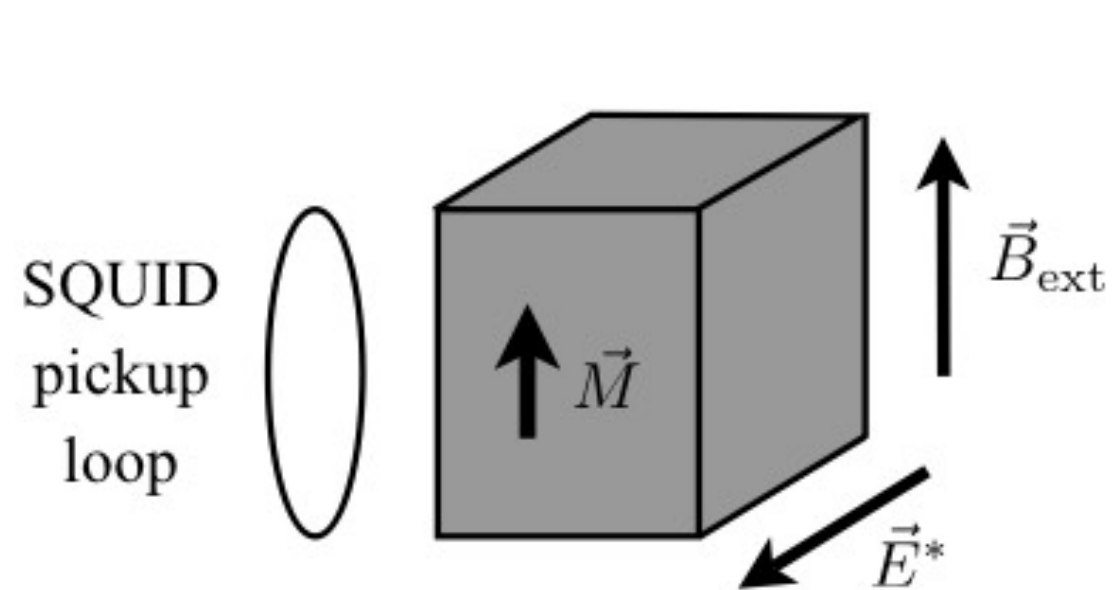


Spin precession

$$\omega = \mu |\vec{B}_{\text{ext}}|$$

CASPER : Spin precession

Mainz, Berkeley



Static EDM, effects cancel in a period

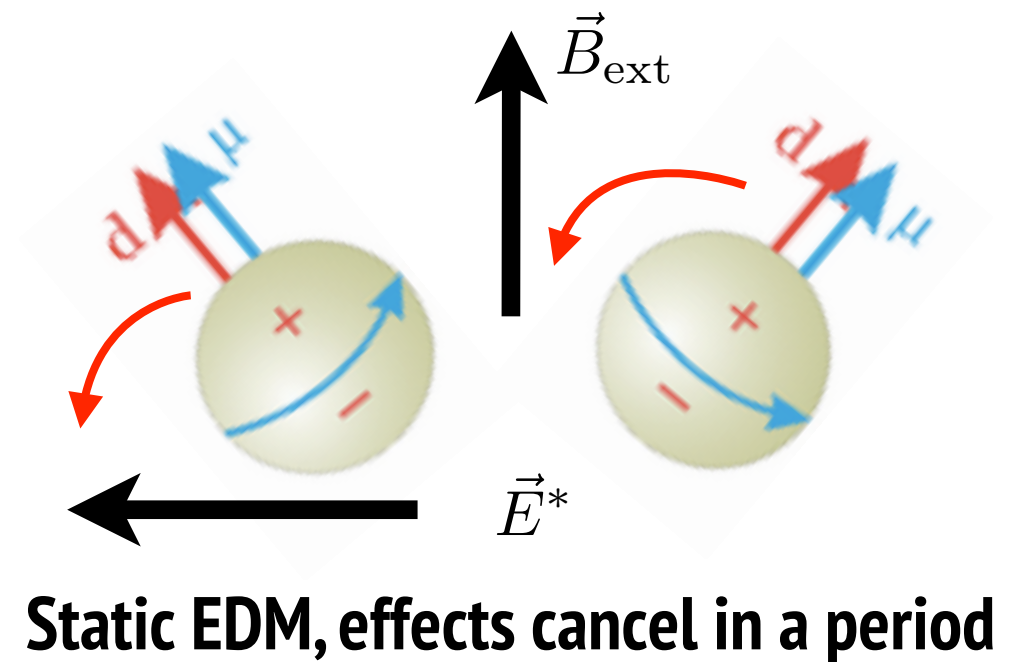
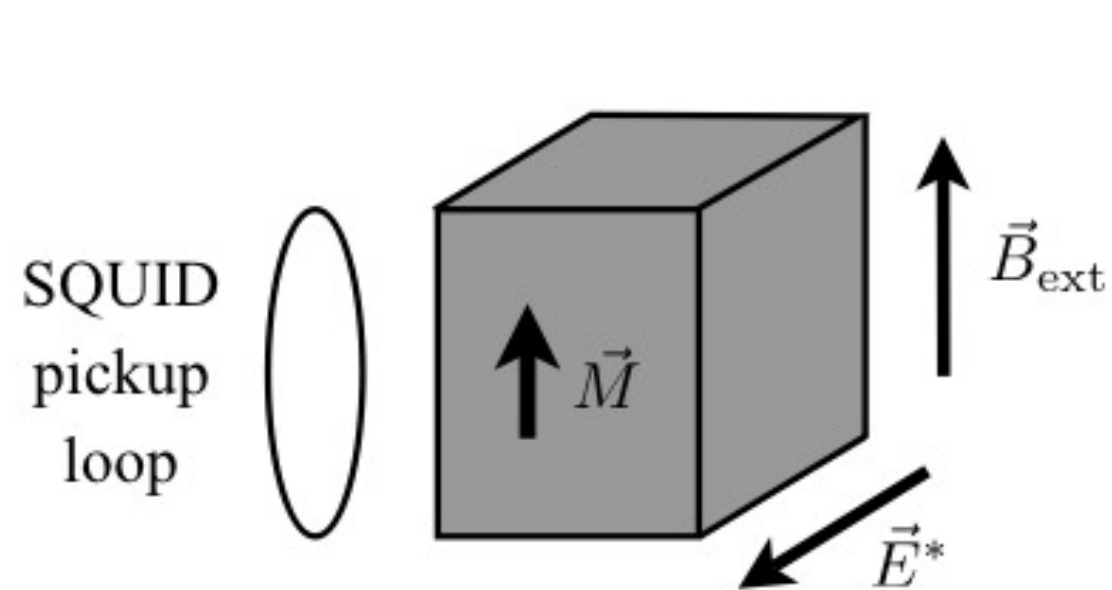
$$\text{magnetic signal} \propto np \epsilon_s d E^* T_2$$

number density nuclear spin polarization Schiff suppression nuclear spin coherence time

- EDM + Large E-fields in PbTiO₃
- Mainz (D. Budker's group) & Berkeley
- B-field, coherence time, sensitivity to $m < \text{neV}$
- Phase I starts in 2016, Phase II physics results
- Mass range limited by B-field strength

CASPER : Spin precession

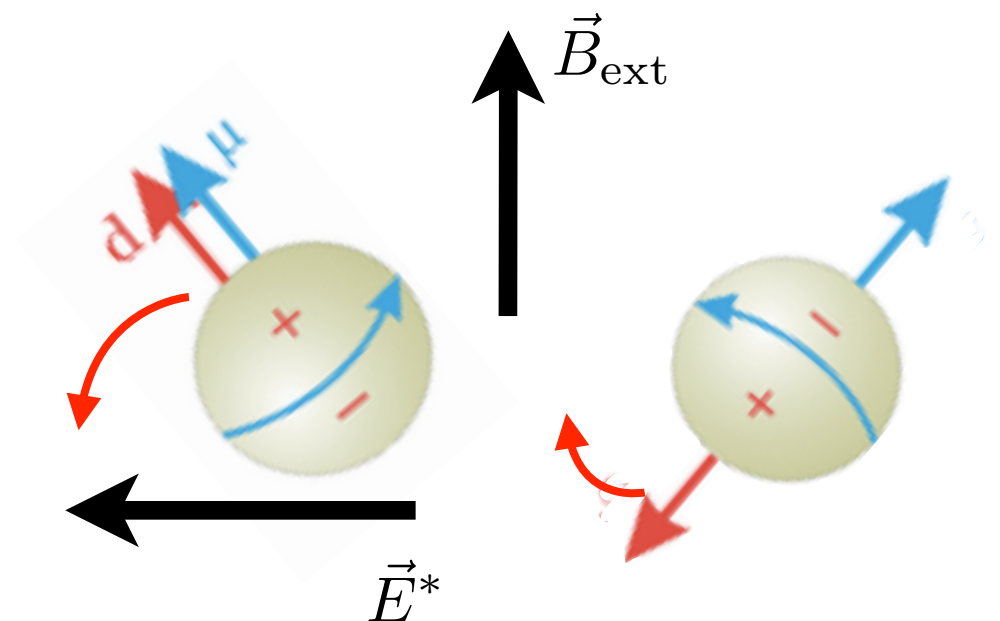
Mainz, Berkeley



$$\text{magnetic signal} \propto np \epsilon_s d E^* T_2$$

number density $\rightarrow$ n
 nuclear spin polarization $\rightarrow$ p
 Schiff suppression $\rightarrow$ ϵ_s
 nuclear spin coherence time $\rightarrow$ T_2

- EDM + Large E-fields in PbTiO₃
- Mainz (D. Budker's group) & Berkeley
- B-field, coherence time, sensitivity to $m < \text{neV}$
- Phase I starts in 2016, Phase II physics results
- Mass range limited by B-field strength

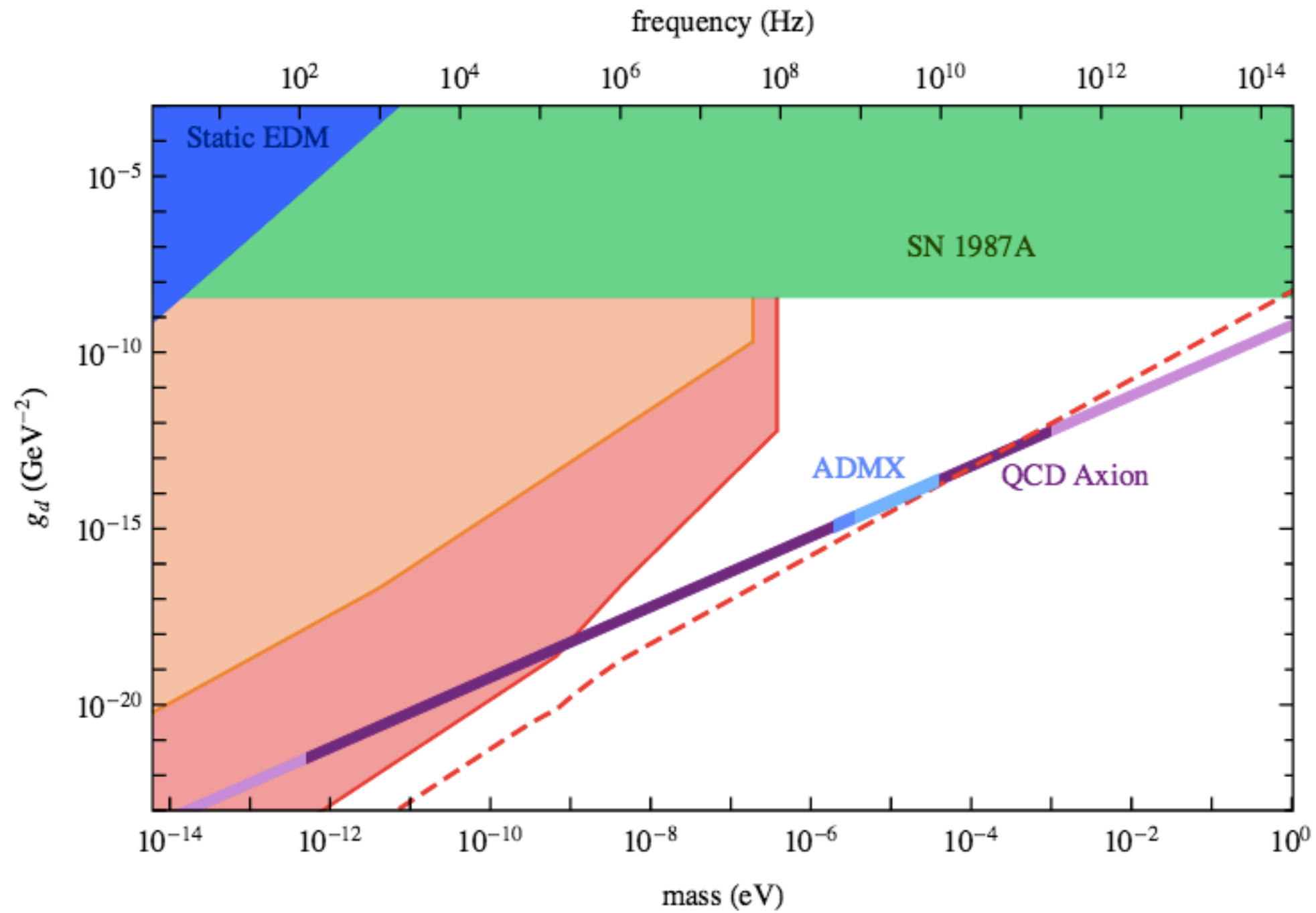


Oscillating EDM, effects add up, transverse magnetisation grows if

$$m_a = \omega = \mu |\vec{B}_{\text{ext}}|$$

CASPER : Spin precession

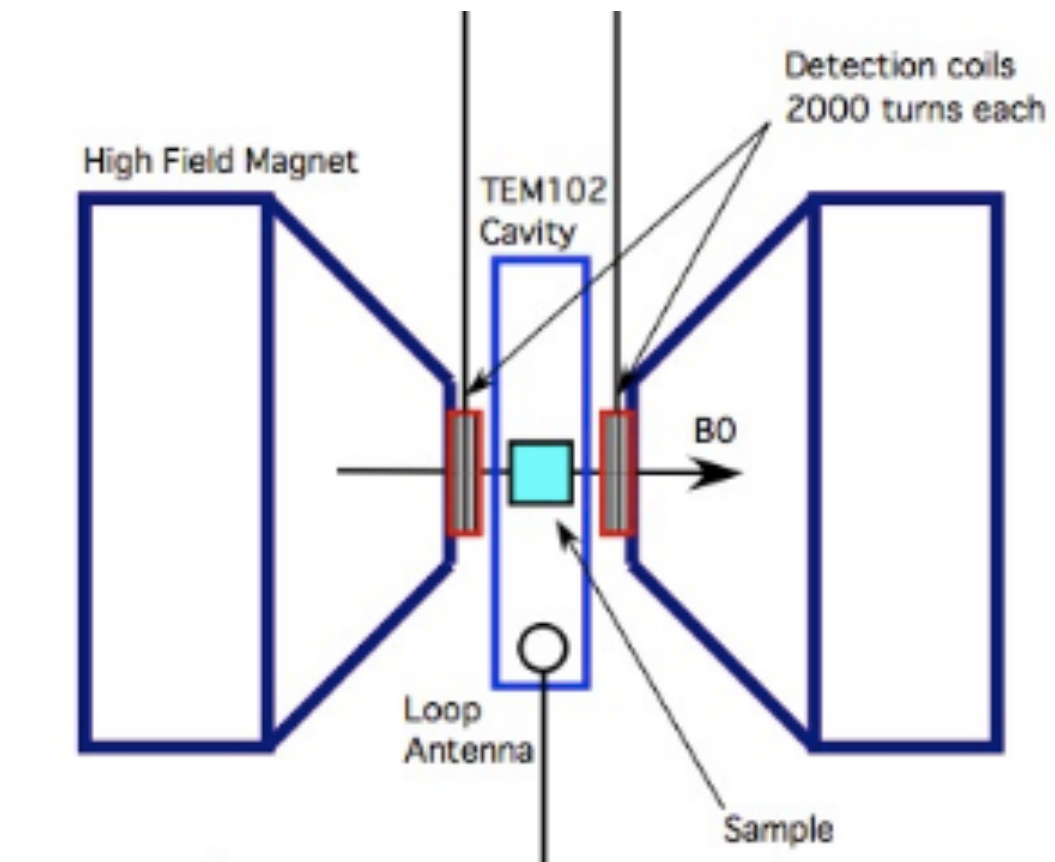
Mainz, Berkeley



- Electron coupling in the non-relativistic limit, Electron spin - axion “wind”

$$\mathcal{L}_{ae} = C_{ae} [\bar{e} \gamma^\mu \gamma_5 e] \partial_\mu \theta \quad \rightarrow \quad C_{ae} (\nabla \theta) \cdot \sigma_e \mu_B \quad \sim \quad C_{ae} m_a \underbrace{\langle \vec{v} \rangle \theta}_{\text{Effective Magnetic field}} \cdot \sigma_e \mu_B$$

Effective Magnetic field



- Use Electron Spin Resonance (similar to NMR but with electrons) $\omega = \mu_B \vec{B}_{\text{ext}} = m_a$
- Bohr magneton much larger, smaller B-fields required for large axion mass
- Short coherence times, radiation damping (R+D)
- HF detection ? Use non-linearity and search for LF oscillations $\omega \sim \mu_B |\vec{B}_{\text{ext}}| - m_a$

Axion DM : A developing picture

