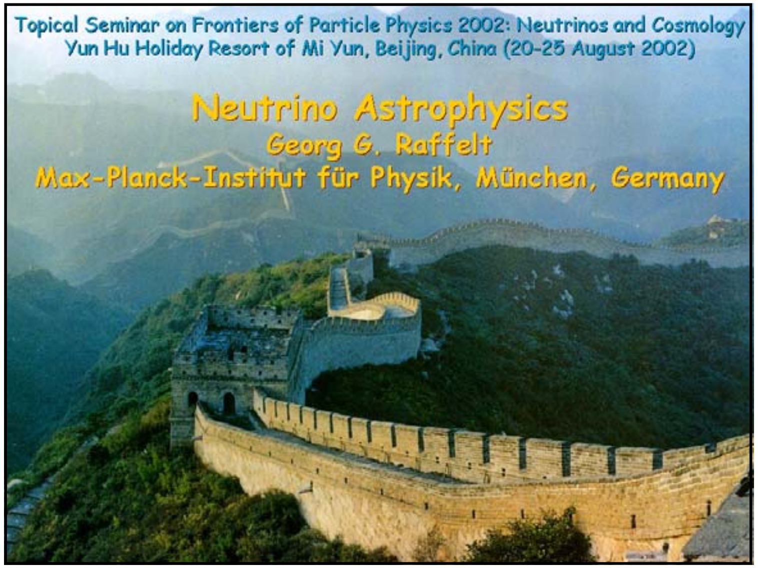
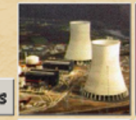
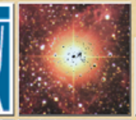


Where do Neutrinos Appear in Nature?



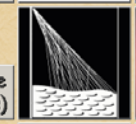
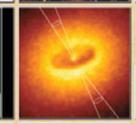

✓ Nuclear Reactors

Sun ✓



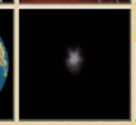
✓ Particle-Accelerators

Supernovae (Stellar Collapse) SN 1987A ✓

✓ Earth Atmosphere (Cosmic Rays)

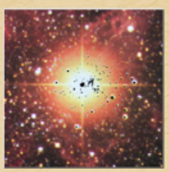
Astrophysical Accelerators Soon ?

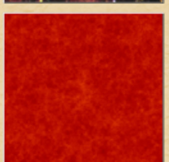
2002 ? Earth Crust (Natural Radioactivity

Cosmic Big Bang (Today 330 ν/cm^3) Indirect Evidence

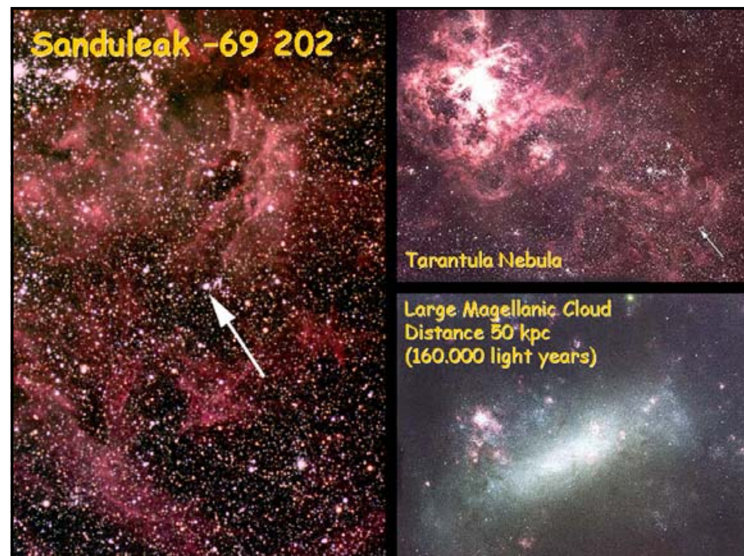
Neutrino Astrophysics

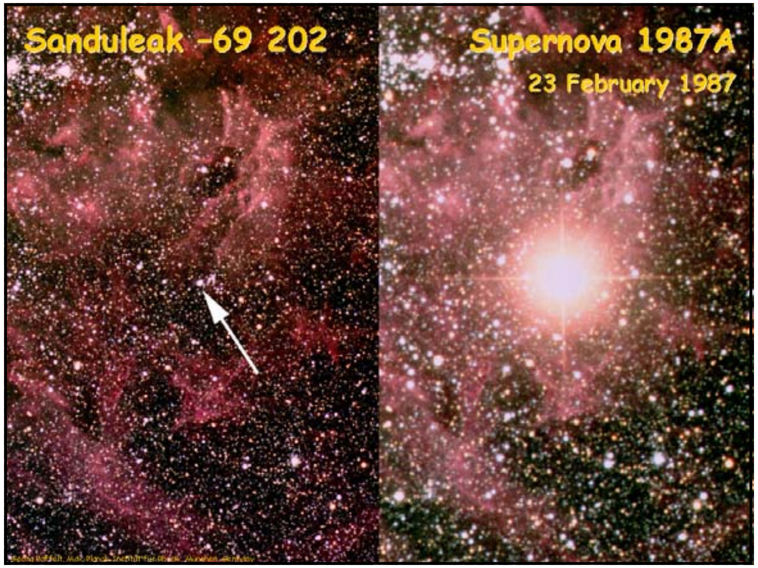


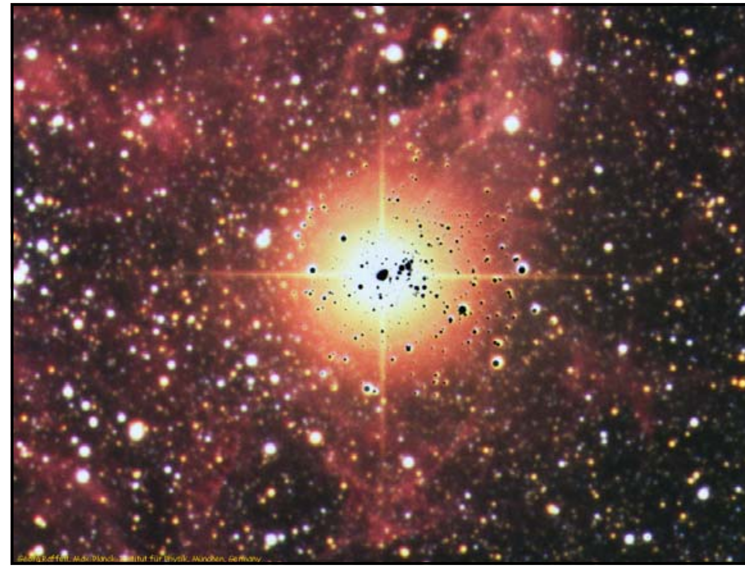
Lecture I: Physics with Supernovae

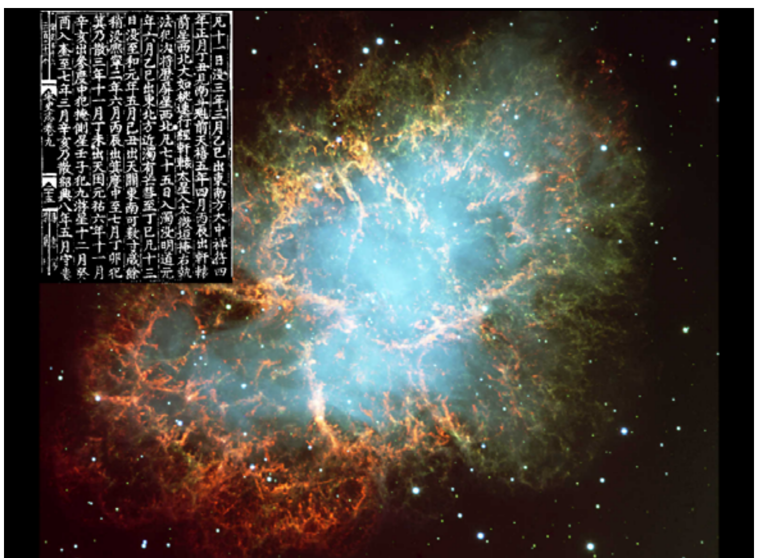


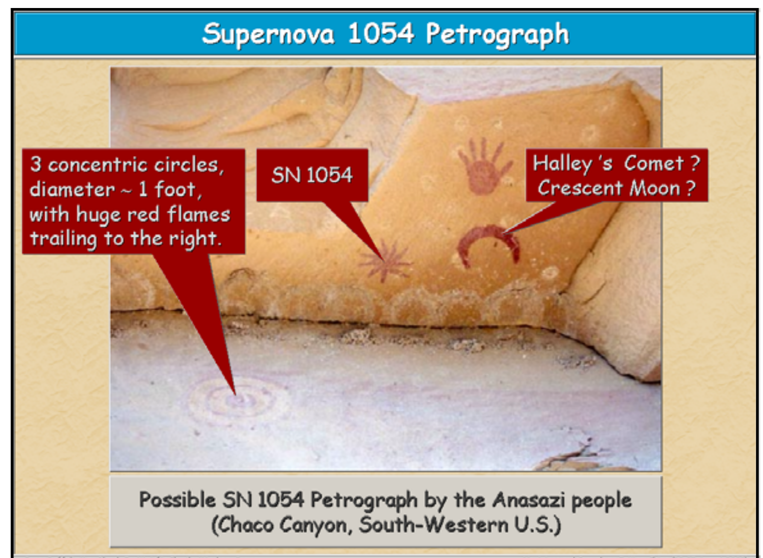
Lecture II: Cosmological Neutrinos




Sanduleak -69 202
Supernova 1987A
23 February 1987



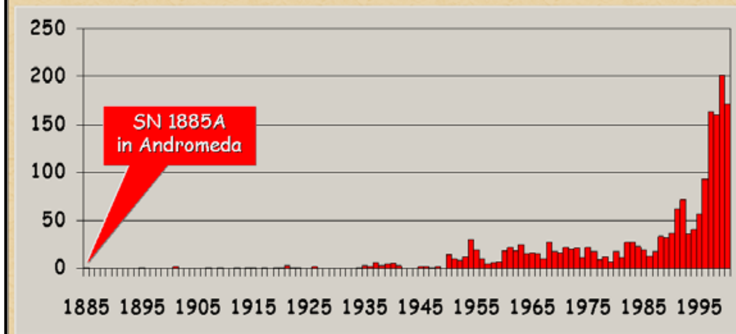



Supernova 1054 Petrograph
3 concentric circles,
diameter ~ 1 foot,
with huge red flames
trailing to the right.
SN 1054
Halley 's Comet ?
Crescent Moon ?
Possible SN 1054 Petrograph by the Anasazi people
(Chaco Canyon, South-Western U.S.)
Georg Raffelt, Max-Planck-Institut für Physik, München, Germany
Neutrinos and Cosmology, Beijing, China (20-25 August 2002)

Classification of Supernovae

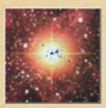
Spectral Type	Ia	Ib	Ic	II
Spectrum	No Hydrogen			Hydrogen
	Silicon	No Silicon		
		Helium	No Helium	
Physical Mechanism	Nuclear explosion of low-mass star	Core collapse of evolved massive star (may have lost its hydrogen or even helium envelope during red-giant evolution)		
Light Curve	Reproducible	Large variations		
Neutrinos	Insignificant	~ 100 × Visible energy		
Compact Remnant	None	Neutron star (typically appears as pulsar) Sometimes black hole ?		
Rate / h^2 SNu	0.36 ± 0.11	0.14 ± 0.07		0.71 ± 0.34
Observed	Total ~ 2000 as of today (nowadays ~200/year)			

Supernova Discoveries 1885-2000



Naming convention for supernovae:
SN 2000A, SN 2000B ... SN 2000Z,
SN 2000aa, SN 2000ab ... SN 2000az, SN 2000ba ... SN 2000fq

Lecture I: Physics with Supernovae



Physical Mechanism of Core-Collapse Supernovae

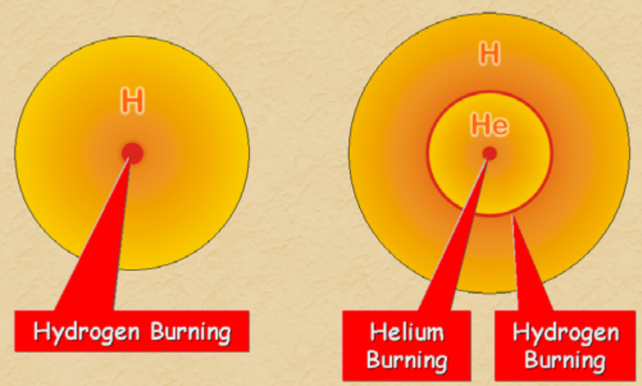
Supernova Neutrino Detection

Limits on Particle Properties

Flavor Oscillations of Supernova Neutrinos

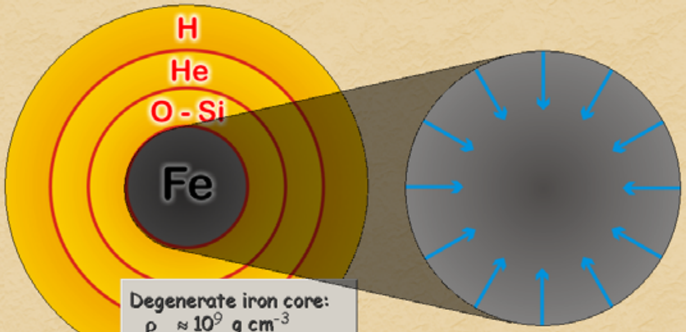
Stellar Collapse and Supernova Explosion

Main Sequence Star	Red Giant Star



Stellar Collapse and Supernova Explosion

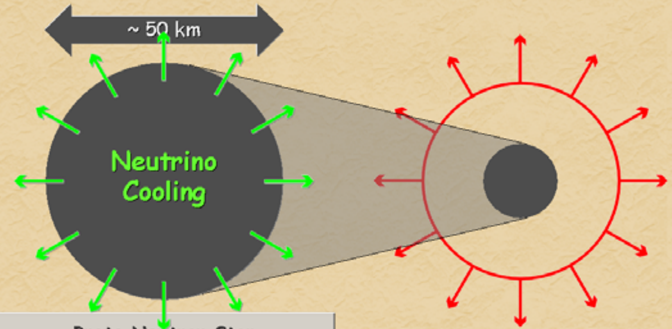
Onion Structure | Collapse (Implosion)



Degenerate iron core:
$\rho \approx 10^{9}\ \mathrm{g\ cm^{-3}}$
$T \approx 10^{10}\ \mathrm{K}$
$M_{Fe} \approx 1.5\ M_{sun}$
$R_{Fe} \approx 8000\ \mathrm{km}$

Stellar Collapse and Supernova Explosion

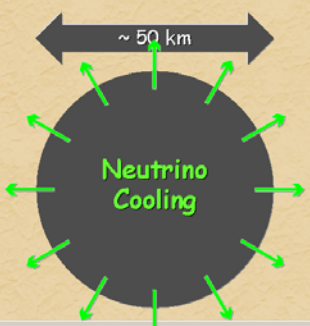
Newborn Neutron Star | Explosion



Proto-Neutron Star
$\rho \approx \rho_{nuc} = 3 \times 10^{14}\ \mathrm{g\ cm^{-3}}$
$T \approx 30\ \mathrm{MeV}$

Stellar Collapse and Supernova Explosion

Newborn Neutron Star



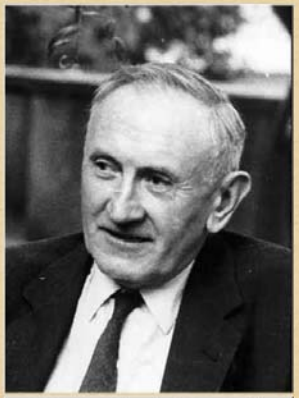
Proto-Neutron Star
$\rho \approx \rho_{nuc} = 3 \times 10^{14}\ \mathrm{g\ cm^{-3}}$
$T \approx 30\ \mathrm{MeV}$

Gravitational binding energy
$E_b \approx 3 \times 10^{53}\ \mathrm{erg} \approx 17\%\ M_{SUN}\,c^2$

This shows up as
99% Neutrinos
1% Kinetic energy of explosion
(1% of this into cosmic rays)
0.01% Photons, outshine host galaxy

Neutrino luminosity
$L_\nu \approx 3 \times 10^{53}\ \mathrm{erg}\ /\ 3\ \mathrm{sec}$
$\approx 3 \times 10^{19}\ L_{SUN}$
While it lasts, outshines the entire visible universe

Walter Baade (1893-1960)

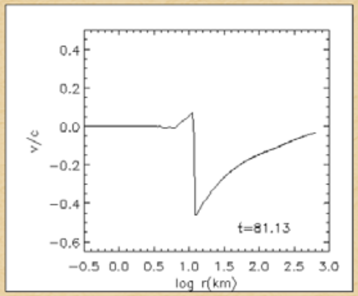


Fritz Zwicky (1898-1974)

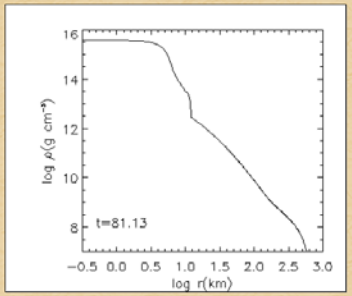
Baade and Zwicky were the first to speculate about a connection between supernova explosions and neutron-star formation [Phys. Rev. 45 (1934) 138]

Collapse and Prompt Explosion

Velocity



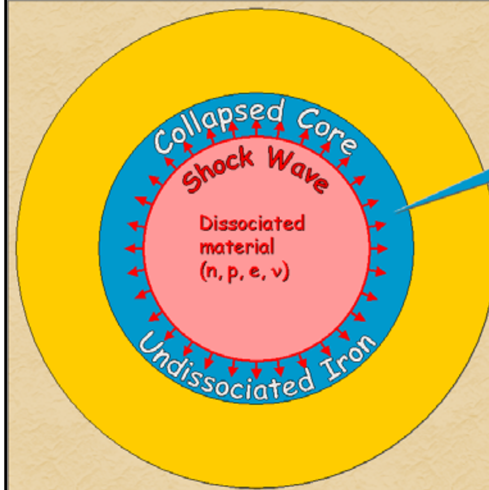
Density



Supernova explosion primarily a hydrodynamical phenomenon

Movies by J.A.Font, Numerical Hydrodynamics in General Relativity
http://www.livingreviews.org

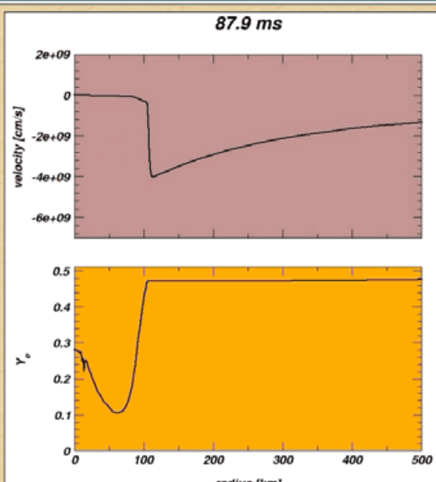
Why No Prompt Explosion?



- $0.1\ M_{sun}$ Fe has nuclear binding energy $\approx 1.7 \times 10^{51}$ erg
- Comparable to explosion energy

- Shock wave forms within the iron core
- Dissipates its energy by dissociating the remaining layer of iron

Failed Explosion

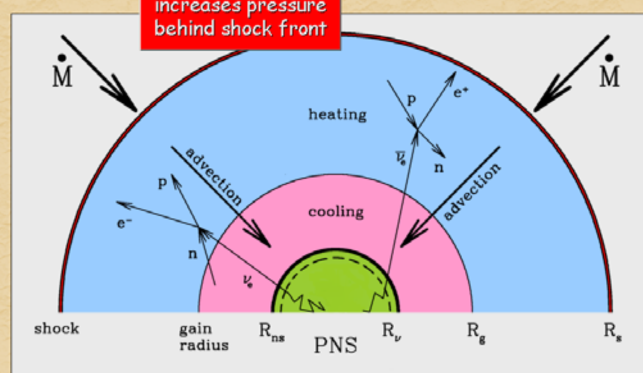


Spherically symmetric simulation of a $15\ M_{sun}$ stellar model with state-of-the-art neutrino transport

Movie courtesy of Bronson Messer, Oakridge group (2001)

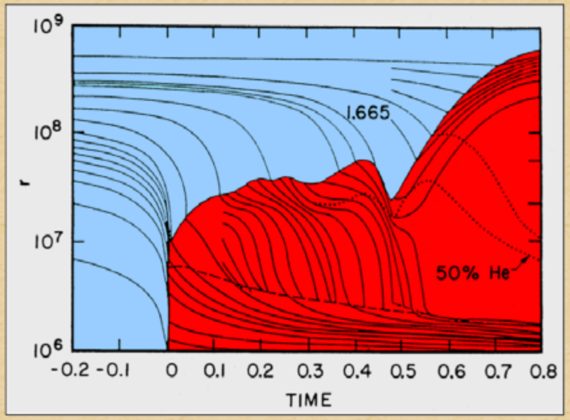
Neutrinos to the Rescue

Neutrino heating increases pressure behind shock front



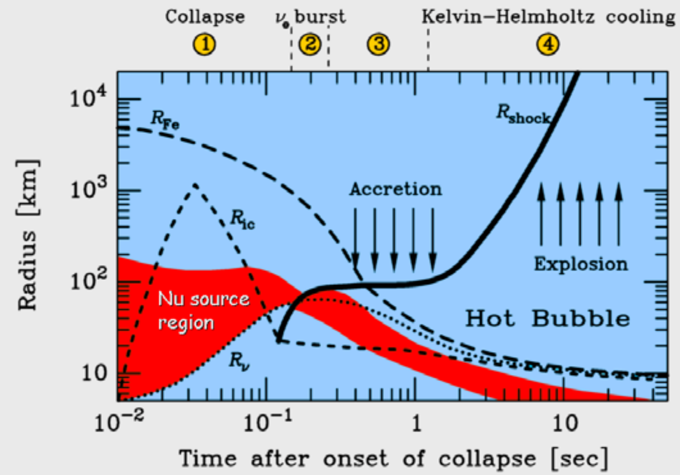
Picture adapted from Janka, astro-ph/0008432

Revival of a Stalled Supernova Shock by Neutrino Heating



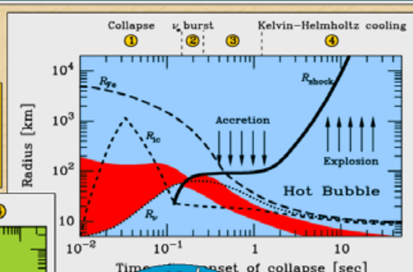
Wilson, Proc. Univ. Illinois Meeting on Numerical Astrophysics (1982)
Bethe & Wilson, ApJ 295 (1985) 14

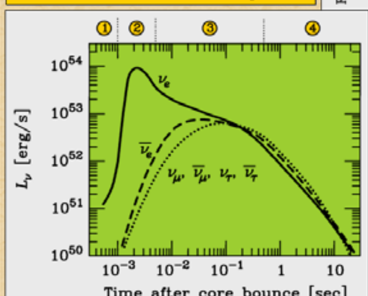
Supernova Collapse and Explosion

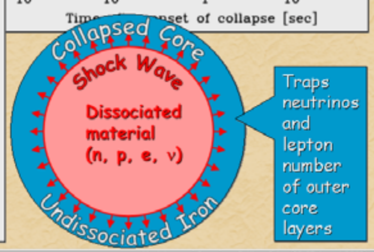


Structure of Supernova Neutrino Signal

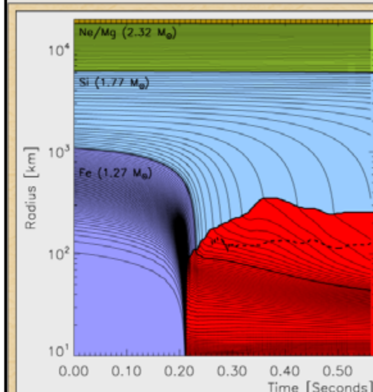
1. Collapse (infall phase)
2. Shock break out
3. Matter accretion
4. Kelvin-Helmholtz cooling



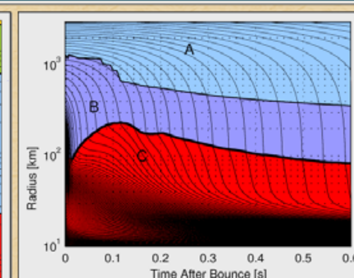




Failed Explosions in Spherical Symmetry



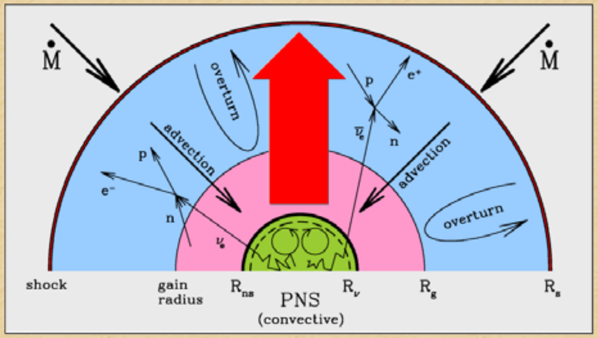
Rampp & Janka, ApJ 539 (2000) L33



Mezzacappa et al., PRL 86 (2001) 1935

Spherically symmetric (1-D) simulations with state-of-the-art neutrino transport do not explode

Novel Forms of Energy Transfer?



New particles or neutrinos with novel properties could provide a new channel of energy transfer from proto neutron star to shock wave

Must not tranfer too much energy → Limits on decaying neutrinos [Falk & Schramm, PLB 79 (1978) 511]

Shock Revival by Novel Particles?

THE ASTROPHYSICAL JOURNAL, 260:868–874, 1982 September 15

SUPERNOVAE INDUCED BY AXION-LIKE PARTICLES

DAVID N. SCHRAMM
The University of Chicago
AND
JAMES R. WILSON
Lawrence Livermore Laboratory
Received 1981 December 22; accepted 1982 April 1

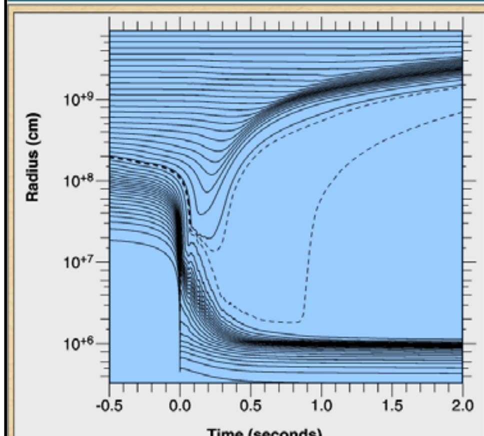
ABSTRACT

It is shown that a new type of particle which may have been seen in a recent accelerator experiment may, if truly present, provide a mechanism whereby gravitationally collapsing massive stars may eject their outer mantles and envelopes in supernova explosions of $\sim 10^{51}$ ergs while leaving the cores to form neutron star remnants. These particles are "axion-like," which means they interact semiweakly, decay to two photons with lifetimes $\sim 10^{-3}$ s, and have masses $0.15 \leq M_a \lesssim 1$ MeV. It is hoped that future accelerator searches will be able to confirm or deny the existence of these particles, the presence of which would cause a dramatic solution to the long-standing gravitational-collapse supernova problem.

Subject headings: elementary particles — nuclear reactions — stars: collapsed — stars: supernovae

Georg Raffelt, Max-Planck-Institut für Physik, München, Germany
Neutrinos and Cosmology, Beijing, China (20-25 August 2002)

Neutron-Finger Convection to the Rescue



Livermore group obtains robust delayed explosions with 1-D code of Mayle & Wilson

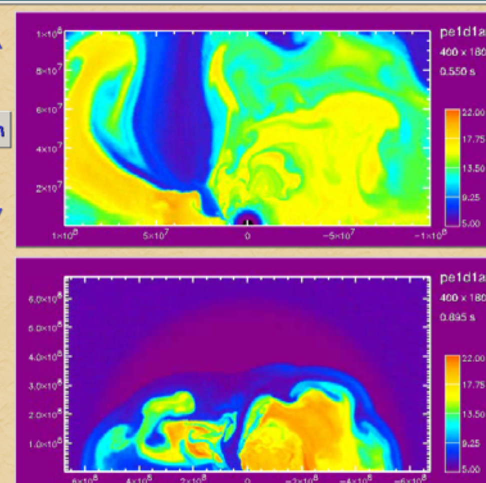
Neutrino luminosity is enhanced by "neutron finger convection" in proto neutron star

Totani, Sato, Dalhed & Wilson, ApJ 496 (1998) 216

Georg Raffelt, Max-Planck-Institut für Physik, München, Germany
Neutrinos and Cosmology, Beijing, China (20-25 August 2002)

Convection in Supernovae (2-D Simulation)

1000 km



Entropy contours

Artificially triggered explosion

Movies courtesy H.-T. Janka

Georg Raffelt, Max-Planck-Institut für Physik, München, Germany
Neutrinos and Cosmology, Beijing, China (20-25 August 2002)

Theoretical Status of Supernova Explosions

- Spherically symmetric models do not explode, even with state-of-the-art Boltzmann solvers for neutrino transport
- Delayed explosion scenario requires enhanced neutrino luminosity at early times (~ factor 2)

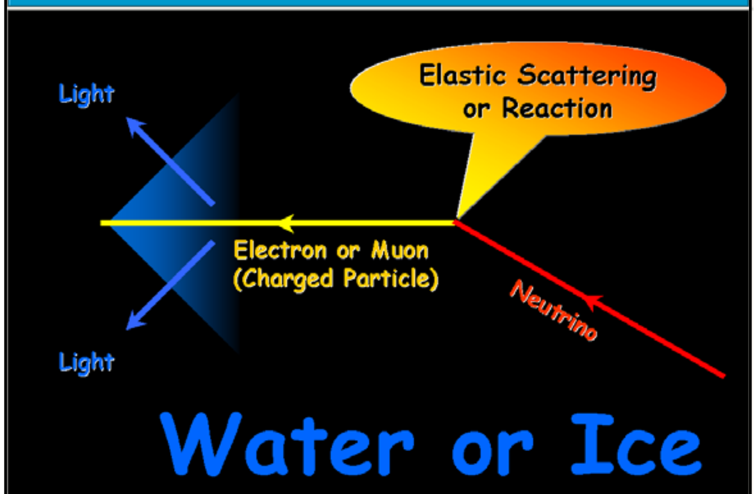
- Convection between proto neutron star (PNS) and shock wave and perhaps within PNS helps
- Next steps: 2-D and 3-D simulations self-consistently coupled with state-of-the-art neutrino transport

- Particle-physics models for new channel of energy transfer can be constructed
- Simplest neutrino flavor-oscillation scenario suppressed by large matter effects relative to small Δm

- New physical ingredients required?
- Explosion a magneto-hydrodynamical effect? (Strong B-fields and fast rotation possible)

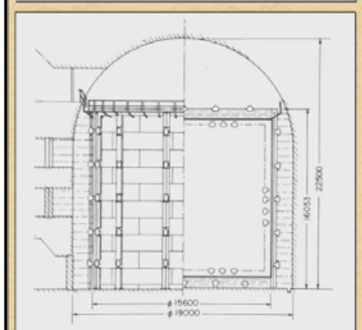
Lecture I: Physics with Supernovae

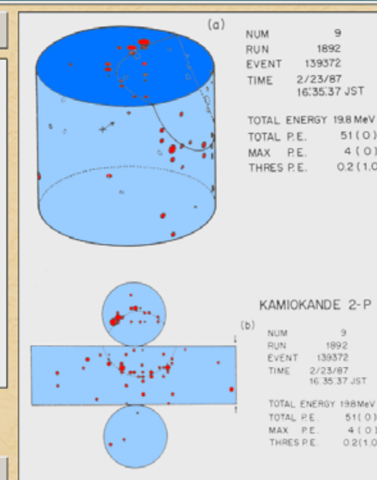
Cherenkov Effect



SN 1987A Event No.9 in Kamiokande

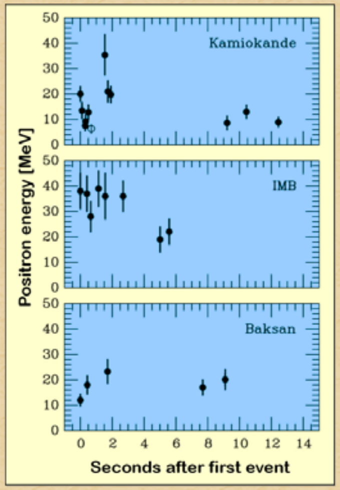
Kamiokande Detector





Hirata et al., PRD 38 (1988) 448

Neutrino Signal of Supernova 1987A



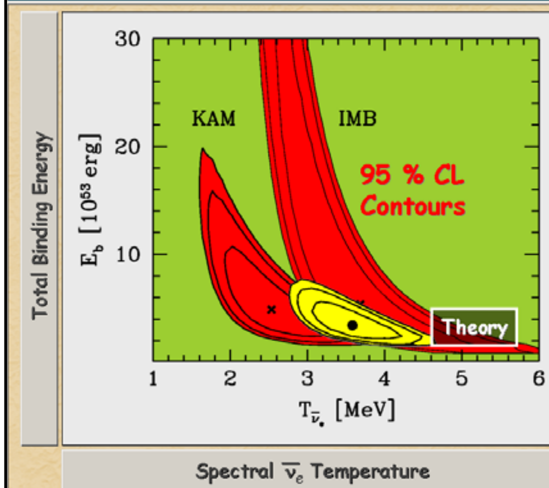
Kamiokande (Japan)
Water Cherenkov detector
Clock uncertainty ±1 min

Irvine-Michigan-Brookhaven (US)
Water Cherenkov detector
Clock uncertainty ±50 ms

Baksan Scintillator Telescope
(Soviet Union)
Clock uncertainty +2/-54 s

Within clock uncertainties, signals are contemporaneous

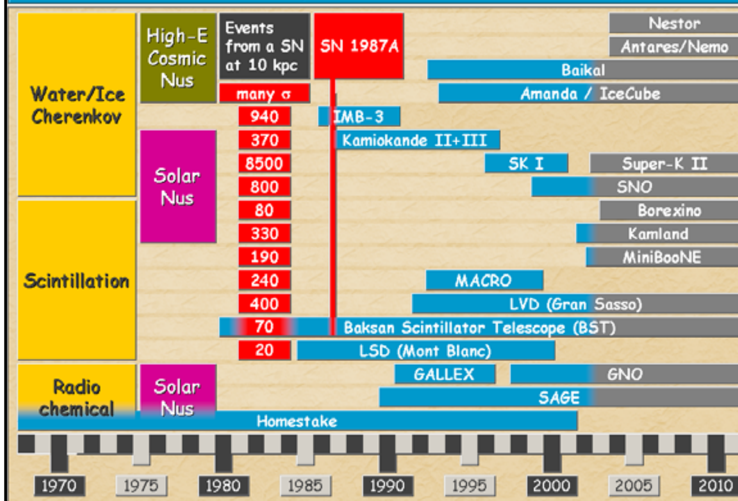
Interpreting SN 1987A Neutrinos



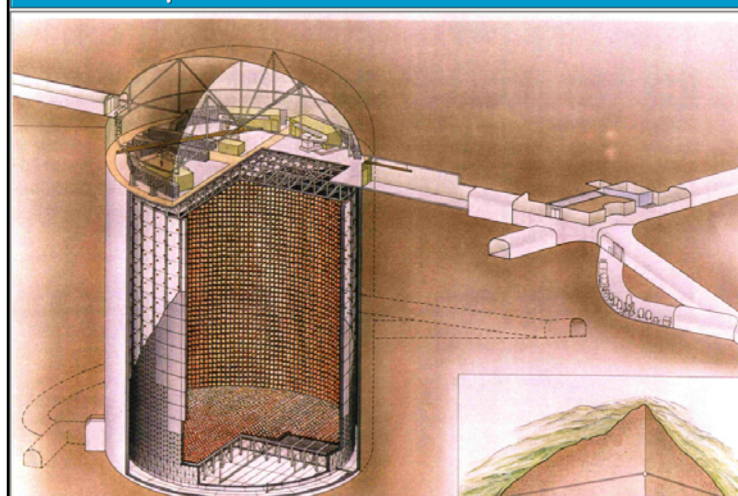
Jegerlehner, Neubig & Raffelt, PRD 54 (1996) 1194

Assume thermal spectra and equipartition of energy between the six degrees of freedom ν_e, ν_μ, ν_τ and their antiparticles

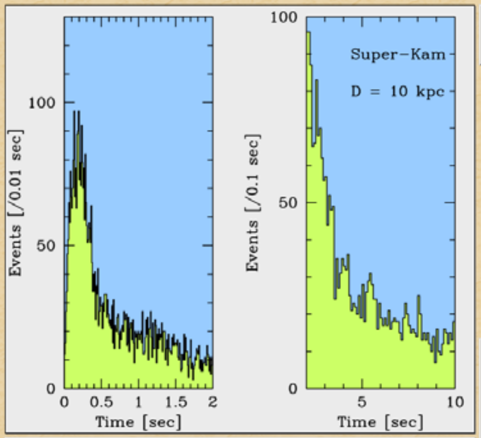
Short History of Neutrino Astronomy



Super-Kamiokande Neutrino Detector



Simulated Supernova Signal in Super-Kamiokande

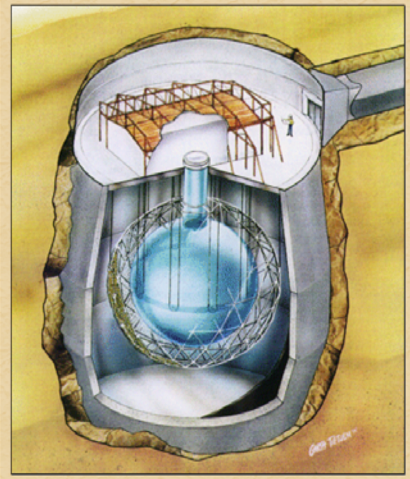


Totani, Sato, Dalhed & Wilson, ApJ 496 (1998) 216

Total of about 8300 events for t < 18 s

Monte-Carlo simulation for Super-Kamiokande signal of SN at 10 kpc, based on a numerical Livermore model

Sudbury Neutrino Observatory (SNO)

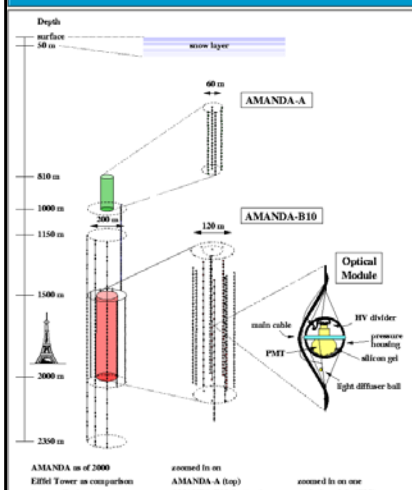


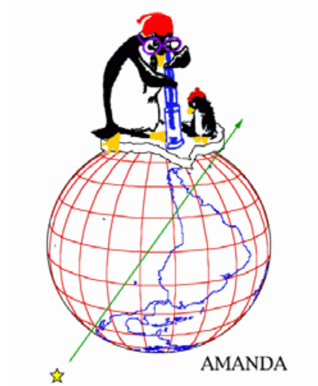
1000 tons of heavy water

Events from a SN at 10 kpc (no flavor oscillations)

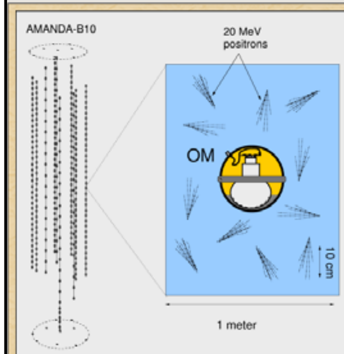
Heavy water (1 kt)	Events:
CC: $\nu_e + d \to p + p + e^-$	72
CC: $\bar\nu_e + d \to n + n + e^+$	138
NC: $\nu_e + d \to \nu_e + p + n$	30
NC: $\bar\nu_e + d \to \bar\nu_e + p + n$	32
NC: $\nu_x + d \to \nu_x + p + n$	164
Light water (1.4 kt)	**Events:**
CC: $\bar\nu_e + p \to n + e^+$	331

AMANDA – Neutrino Telescope at the Southpole

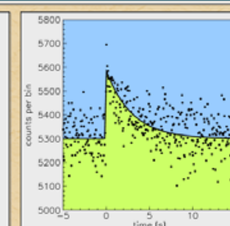




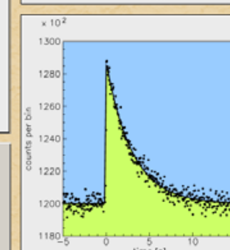
Amanda/IceCube as a Supernova Detector



Each optical module (OM) picks up Cherenkov light from its neighborhood. SN appears as correlated "noise" between OMs



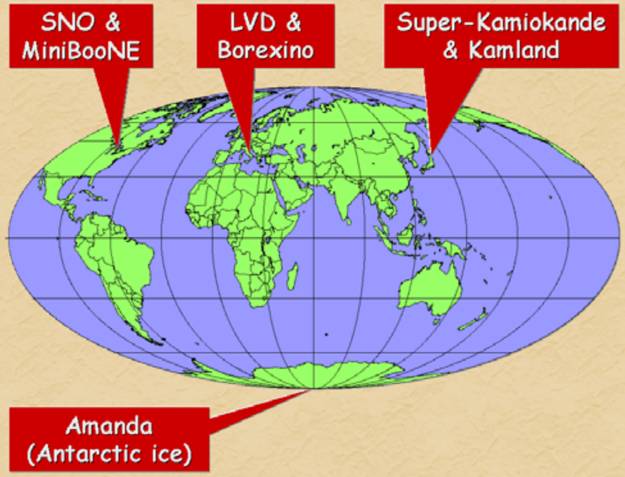
SN @ 8.5 kpc Signal in Amanda



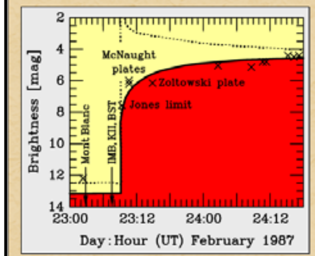
SN @ 8.5 kpc Signal in Ice Cube

Amanda Collaboration (2001)

Large Detectors for SN Neutrinos

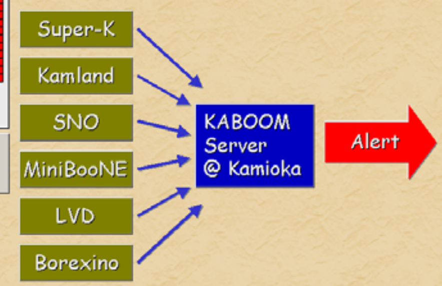


SuperNova Early Warning System (SNEWS)



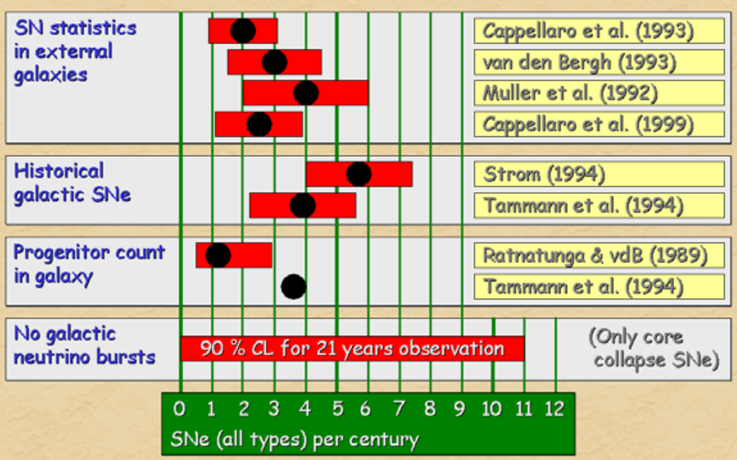
Supernova 1987A Early Light Curve

Neutrino observation can alert astronomers several hours in advance to supernova
To avoid false alarms, require alarm from at least two experiments.

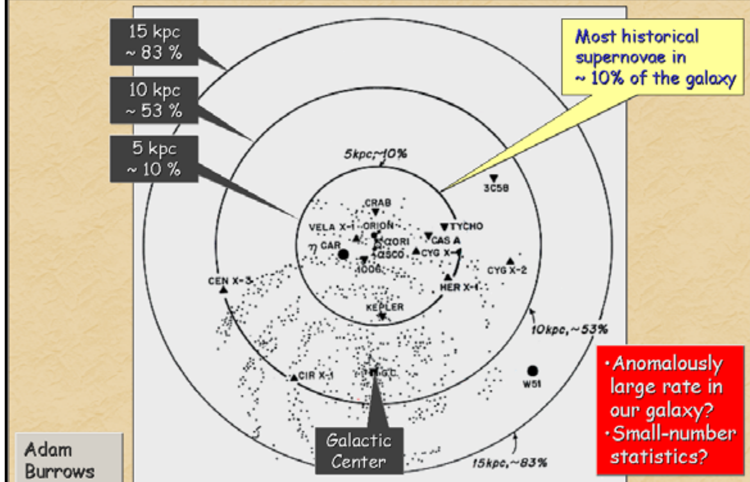


http://hep.bu.edu/~snnet

Estimates of the Galactic SN Rate

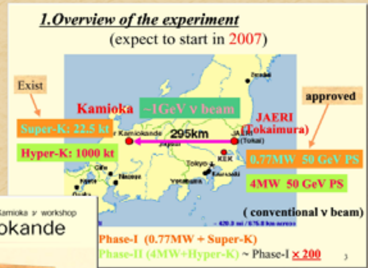


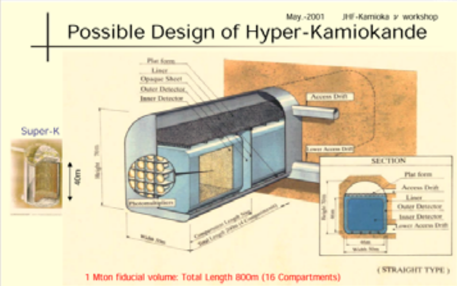
Galactic Supernova Events



The Future: A Megatonne Detector?

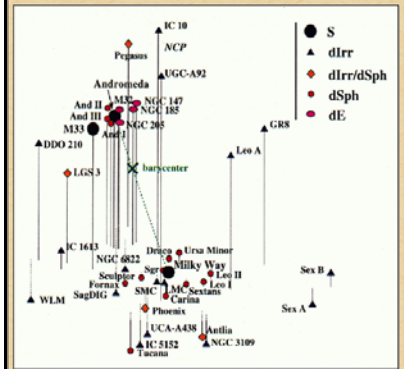
Megatonne detector motivated by
- Long baseline neutrino oscillations
- Proton decay
- Atmospheric neutrinos
- Solar neutrinos
- Supernova neutrinos
  ($\sim 10^5$ events for SN at 10 kpc)



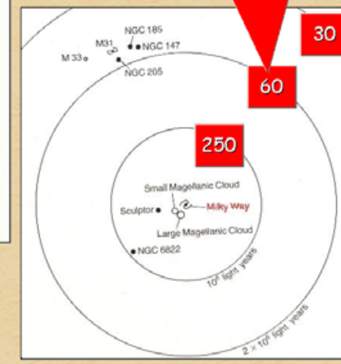


Similar discussions in
- USA (UNO project)
- Europe (Frejus Tunnel)

Local Group of Galaxies



Events in a detector with 30 × Super-K fiducial volume (8000 events in SK at 10 kpc)



Diffuse Background Flux of SN Neutrinos

1 SNu = 1 SN / 10^{10} $L_{sun,B}$ / 100 years
$L_{sun,B} = 0.54\ L_{sun} = 2 \times 10^{33}$ erg/s
$E_\nu \sim 3 \times 10^{53}$ erg per core-collapse SN

→ 1 SNu ~ 4 L_ν / $L_{\gamma,B}$
Average neutrino luminosity of galaxies ~ photon luminosity

- Photons come from nuclear energy
- Neutrinos from gravitational energy

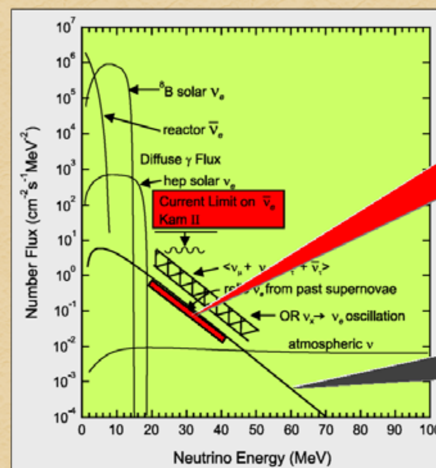
→ For galaxies, average nuclear & gravitational energy release similar

Present-day SN rate of ~ 1 SNu, extrapolated to the entire universe, corresponds to ν_e flux of ~ 1 $cm^{-2}\ s^{-1}$

Realistic flux dominated by much larger early star-formation rate
- Upper limit ~ 54 $cm^{-2}\ s^{-1}$
  [Kaplinghat et al., astro-ph/9912391]
- "Realistic estimate" ~ 10 $cm^{-2}\ s^{-1}$
  [Hartmann & Woosley, Astropart. Phys. 7 (1997) 137]

Measurement would tell us about early history of star formation

Experimental Limits on Relic SN Neutrinos



Preliminary Super-K upper limit of 39 $cm^{-2}\ s^{-1}$ for Kaplinghat et al. spectrum (Totsuka, private comm.) ~ factor 30 improvement

Upper-limit flux of Kaplinghat et al., astro-ph/9912391
Integrated 54 $cm^{-2}\ s^{-1}$

Cline, astro-ph/0103138

Lecture I: Physics with Supernovae

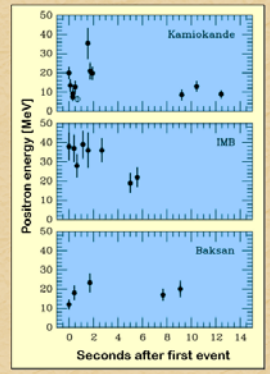
- Physical Mechanism of Core-Collapse Supernovae
- Supernova Neutrino Detection
- **Limits on Particle Properties**
- Flavor Oscillations of Supernova Neutrinos

Neutrino Mass Limits by Signal Dispersion

Time-of-flight delay of massive neutrinos

$$\Delta t = 2.57\,\mathrm{s}\left(\frac{D}{50\,\mathrm{kpc}}\right)\left(\frac{10\,\mathrm{MeV}}{E_\nu}\right)^2\left(\frac{m_\nu}{10\,\mathrm{eV}}\right)^2$$

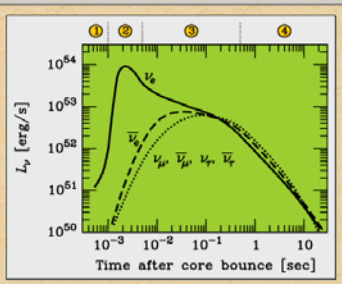
SN 1987A



Neutrino Mass Limits by Signal Dispersion

Time-of-flight delay of massive neutrinos

$$\Delta t = 2.57\,\mathrm{s}\left(\frac{D}{50\,\mathrm{kpc}}\right)\left(\frac{10\,\mathrm{MeV}}{E_\nu}\right)^2\left(\frac{m_\nu}{10\,\mathrm{eV}}\right)^2$$

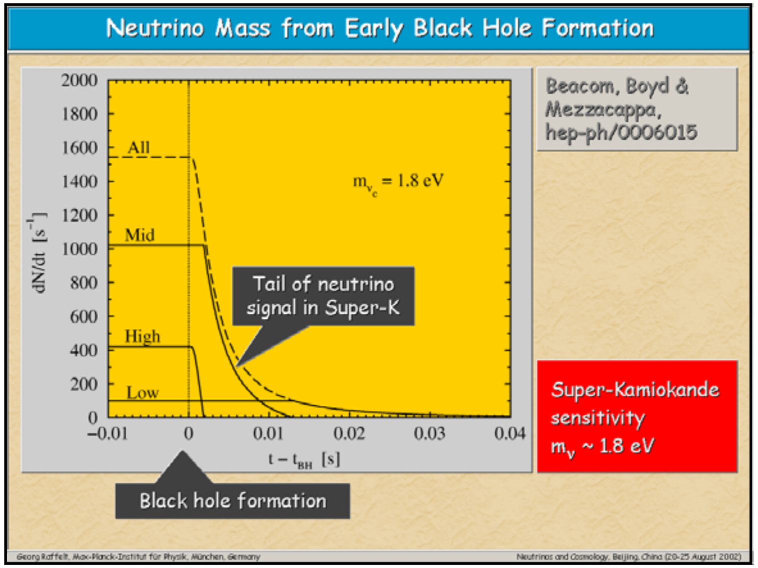
SN 1987A	$E \approx 20$ MeV, $\Delta t \approx 10$ s, $D \approx 50$ kpc Simple estimate or detailed maximum likelihood give similar results	$m_{\nu_e} < 20$ eV
Future Galactic SN (Super-K)		

Neutrino Mass Limits by Signal Dispersion

Time-of-flight delay of massive neutrinos

$$\Delta t = 2.57\,\mathrm{s}\left(\frac{D}{50\,\mathrm{kpc}}\right)\left(\frac{10\,\mathrm{MeV}}{E_\nu}\right)^2\left(\frac{m_\nu}{10\,\mathrm{eV}}\right)^2$$

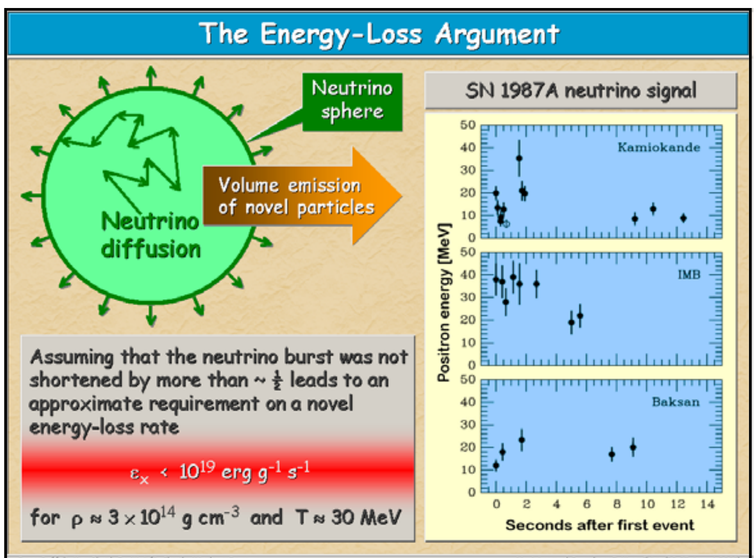
SN 1987A	$E \approx 20$ MeV, $\Delta t \approx 10$ s, $D \approx 50$ kpc Simple estimate or detailed maximum likelihood give similar results	$m_{\nu_e} < 20$ eV
Future Galactic SN (Super-K)	$D \approx 10$ kpc, Rise-time 0.01 s Sensitivity approximately [T.Totani, PRL 80 (1998) 2040]	$m_{\nu_e} \sim 3$ eV
Future SN in Andromeda (Megatonne)	$D \approx 750$ kpc, $\Delta t \approx 10$ s Sensitivity approximately	$m_{\nu_e} \sim 1$-2 eV

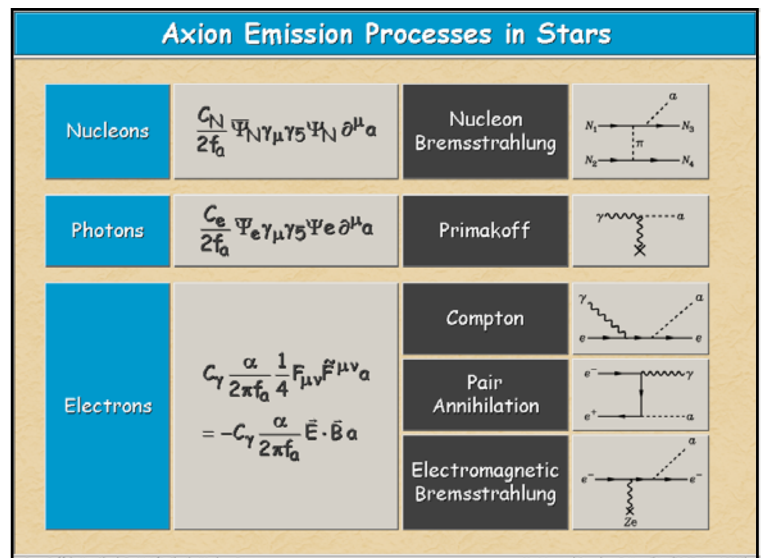


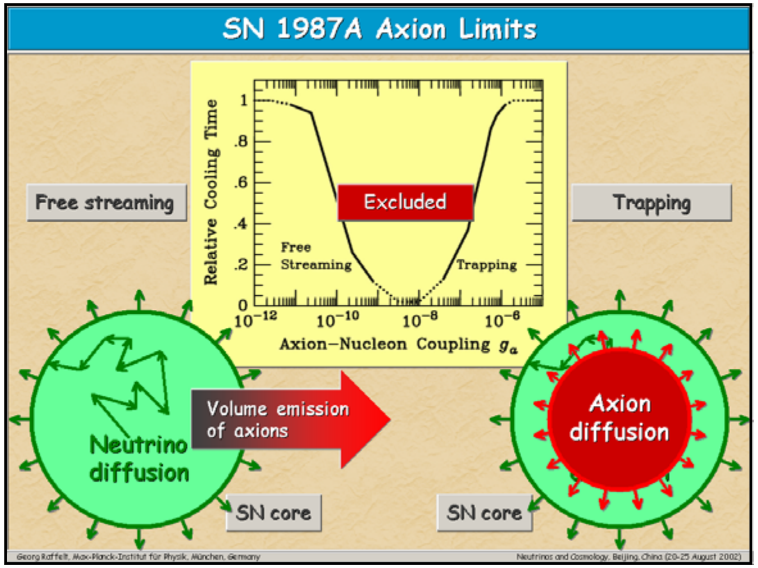
Neutrino Mass Limits and Future Sensitivity

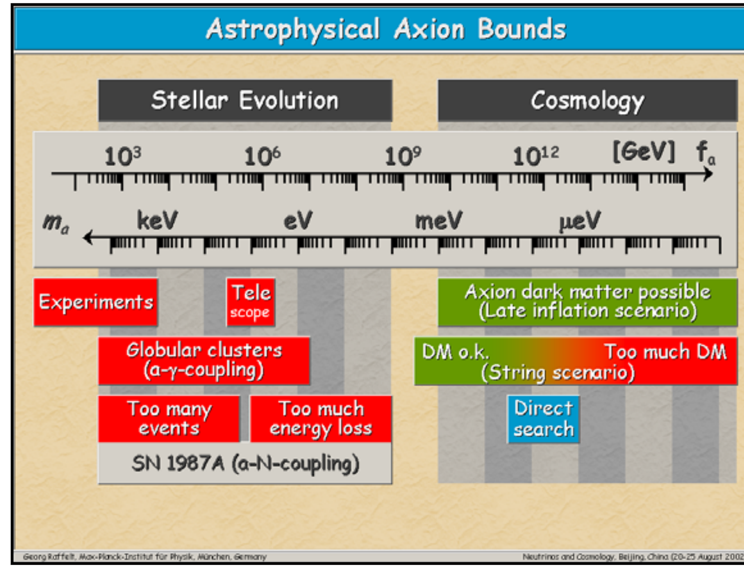
Tritium endpoint	Mainz/Troitsk	2.5 eV
	KATRIN	0.3 eV
Supernova Nus Time-of-flight	SN 1987A	20 eV
	Super-Kamiokande	3 eV
	with black hole	2 eV
	with gravity waves	1 eV
Cosmic structure	2dF Redshift Survey	0.8 eV
	Sloan Digital Sky Survey	0.3 eV

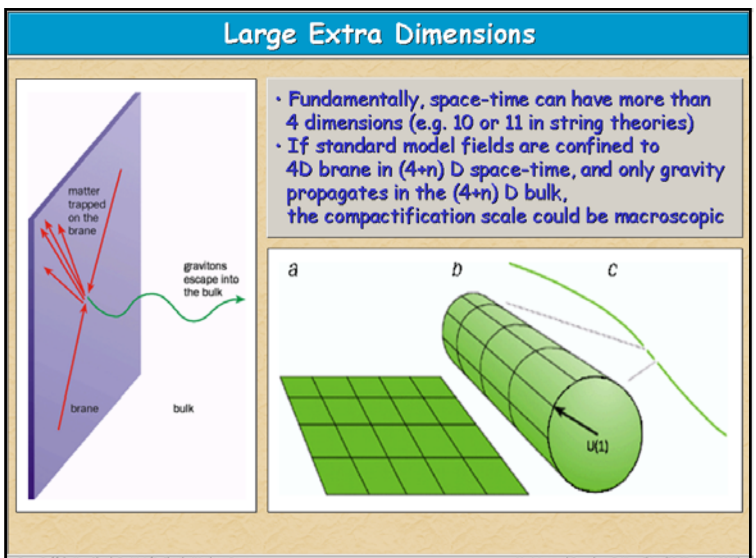
- Assume 3 mass eigenstates with very small mass differences as indicated by atmospheric and solar neutrinos
- The cosmological limit refers to $m_\nu = \Sigma\, m_\nu/3$

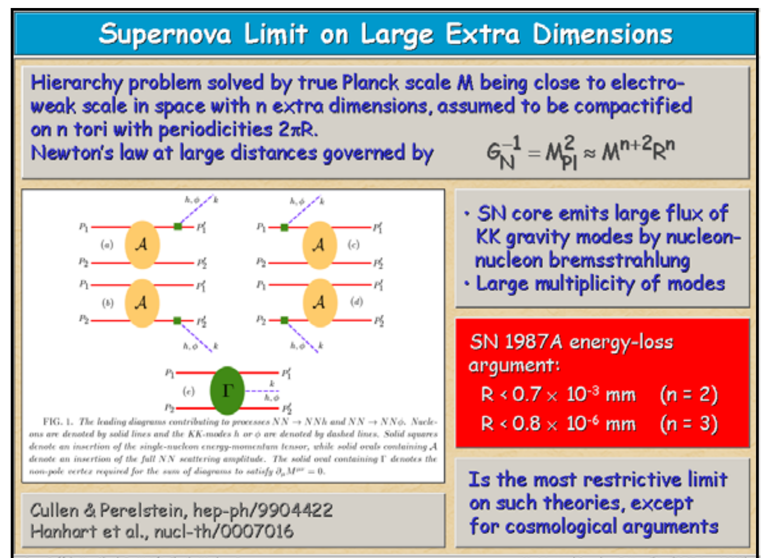




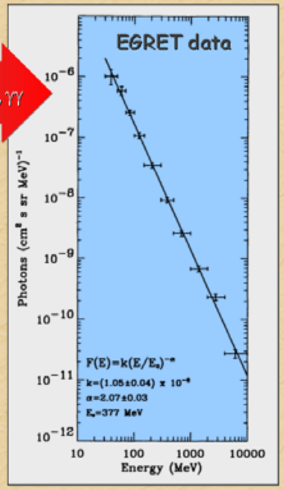

SN 1987A Axion Limits
Free streaming
Excluded
Trapping
Relative Cooling Time
1
.8
.6
.4
.2
0
Free Streaming
Trapping
10^{-12} 10^{-10} 10^{-8} 10^{-6}
Axion–Nucleon Coupling g_a
Volume emission of axions
Neutrino diffusion
Axion diffusion
SN core
SN core
Georg Raffelt, Max-Planck-Institut für Physik, München, Germany
Neutrinos and Cosmology, Beijing, China (20-25 August 2002)


Astrophysical Axion Bounds
Stellar Evolution
Cosmology
10^3 10^6 10^9 10^{12} [GeV] f_a
m_a keV eV meV μeV
Experiments
Tele scope
Axion dark matter possible (Late inflation scenario)
Globular clusters (a-γ-coupling)
DM o.k.
Too much DM
(String scenario)
Too many events
Too much energy loss
Direct search
SN 1987A (a-N-coupling)
Georg Raffelt, Max-Planck-Institut für Physik, München, Germany
Neutrinos and Cosmology, Beijing, China (20-25 August 2002)


Large Extra Dimensions
matter trapped on the brane
gravitons escape into the bulk
brane
bulk
• Fundamentally, space-time can have more than 4 dimensions (e.g. 10 or 11 in string theories)
• If standard model fields are confined to 4D brane in (4+n) D space-time, and only gravity propagates in the (4+n) D bulk, the compactification scale could be macroscopic
a
b
c
U(1)
Georg Raffelt, Max-Planck-Institut für Physik, München, Germany
Neutrinos and Cosmology, Beijing, China (20-25 August 2002)


Supernova Limit on Large Extra Dimensions
Hierarchy problem solved by true Planck scale M being close to electro-weak scale in space with n extra dimensions, assumed to be compactified on n tori with periodicities 2πR.
Newton's law at large distances governed by $G_N^{-1} = M_{Pl}^2 \approx M^{n+2} R^n$
• SN core emits large flux of KK gravity modes by nucleon-nucleon bremsstrahlung
• Large multiplicity of modes
SN 1987A energy-loss argument:
R < 0.7 × 10^{-3} mm (n = 2)
R < 0.8 × 10^{-6} mm (n = 3)
Is the most restrictive limit on such theories, except for cosmological arguments
Cullen & Perelstein, hep-ph/9904422
Hanhart et al., nucl-th/0007016
Georg Raffelt, Max-Planck-Institut für Physik, München, Germany
Neutrinos and Cosmology, Beijing, China (20-25 August 2002)

Improved Limits on Large Extra Dimensions



- From all SNe in the universe, KK decay contributes to diffuse cosmic γ-rays in 100 MeV range
- EGRET data & conservative estimate of SN rate:
  < 1 % of SN energy into KK gravitons
  i.e. 0.01 of SN 1987A cooling limit

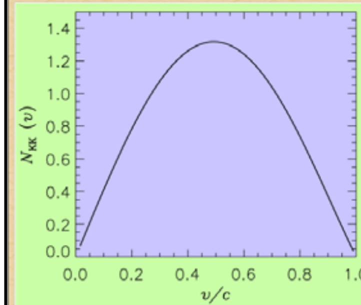
Our new limits $R < 0.1 \times 10^{-3}$ mm (n = 2)
$R < 0.2 \times 10^{-6}$ mm (n = 3)

Hannestad & Raffelt, hep-ph/0103201

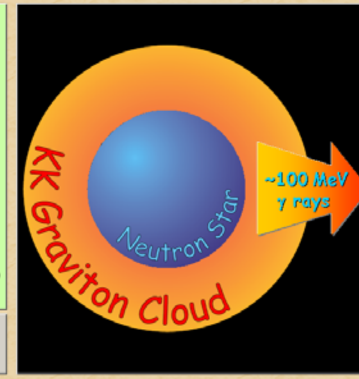
KK Graviton Retention by Neutron Star

$$v_{escape} = \sqrt{\frac{2GM}{R}} \approx 0.6c$$

- Neutron stars retain 50-60% of KK gravitons in a halo
- Emits γ rays by KK decays



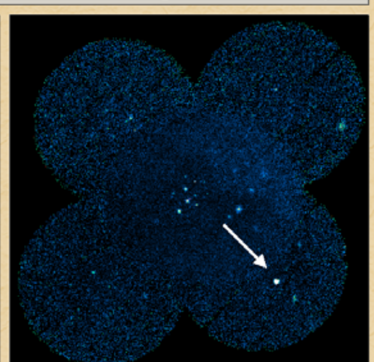
Velocity distribution of KK-gravitons emitted by supernova



Nearby Neutron Star RX J185635-3754

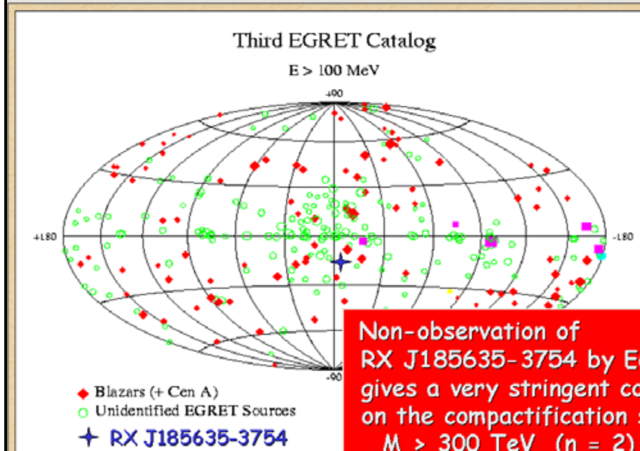
D = 61 pc (closest known neutron star), Age ~ 1.2×10^6 yr

HST Image
(Walter & Matthews 1997)



ROSAT Image
(Walter, Wolk & Neuhauser 1996)

Third EGRET Catalog (Hartmann et al. 1999)

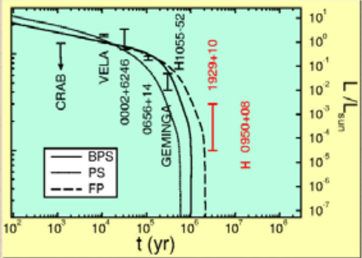


Non-observation of RX J185635-3754 by EGRET gives a very stringent constraint on the compactification scale:
M > 300 TeV (n = 2)
M > 20 TeV (n = 3)

Neutron Star Excess Heat

- Neutron stars retain 50-60% of KK gravitons in a halo
- Emits γ rays by KK decays

Neutron star cooling calculations vs. observations (Pavlov, Stringfellow & Cordova 1996, Larson & Link 1999)



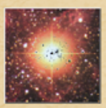
To avoid excess heating by KK decay
M > 1600 TeV (n = 2)
M > 60 TeV (n = 3)

Summary of Limits on Large Extra Dimensions

		M^{min} [TeV]		R^{max} [mm]	
		n = 2	n = 3	n = 2	n = 3
Laboratory experiments		0.6	-	0.2	-
SN 1987A neutrino signal		30	3	7×10^{-4}	8×10^{-7}
EGRET	Cosmic SNe	84	7	8×10^{-5}	2×10^{-7}
	Cas A	73	7	1×10^{-4}	2×10^{-7}
	PSR J0953+0755	300	19	8×10^{-6}	4×10^{-8}
	RX J185635-3754	454	27	3×10^{-6}	2×10^{-8}
Excess heat PSR J0953+0755		1680	60	2×10^{-7}	5×10^{-9}

Hannestad & Raffelt, PRL 88 (2002) 071301 [hep-ph/0103201]

Lecture I: Physics with Supernovae

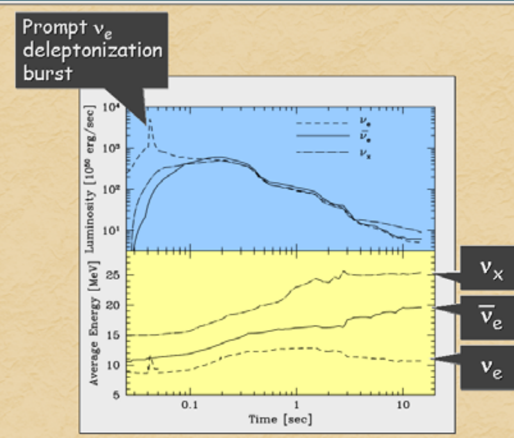

Physical Mechanism of Core-Collapse Supernovae


Supernova Neutrino Detection

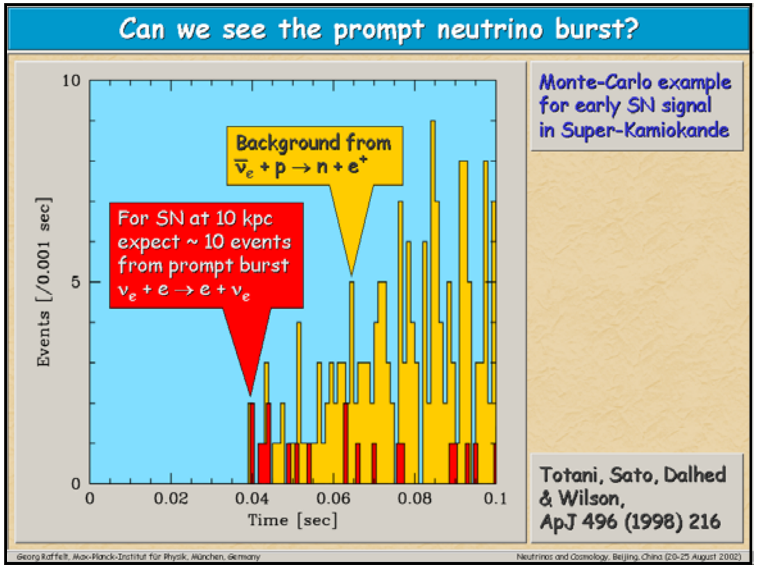

Limits on Particle Properties

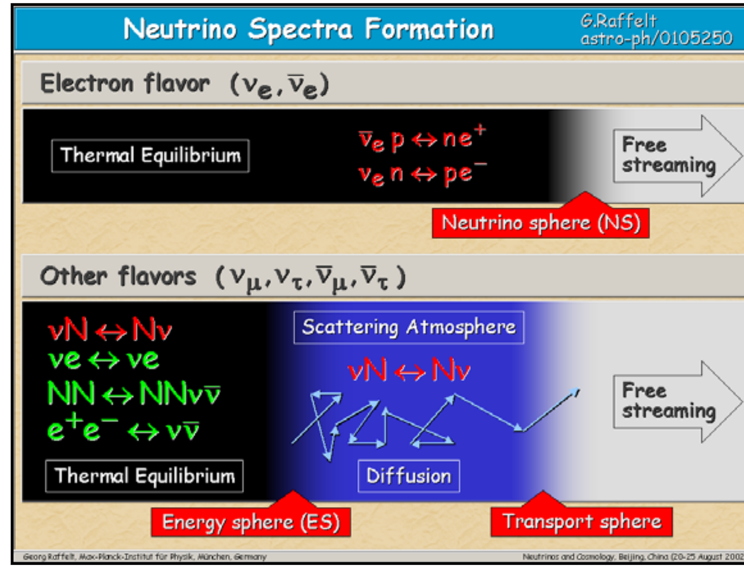

Flavor Oscillations of Supernova Neutrinos

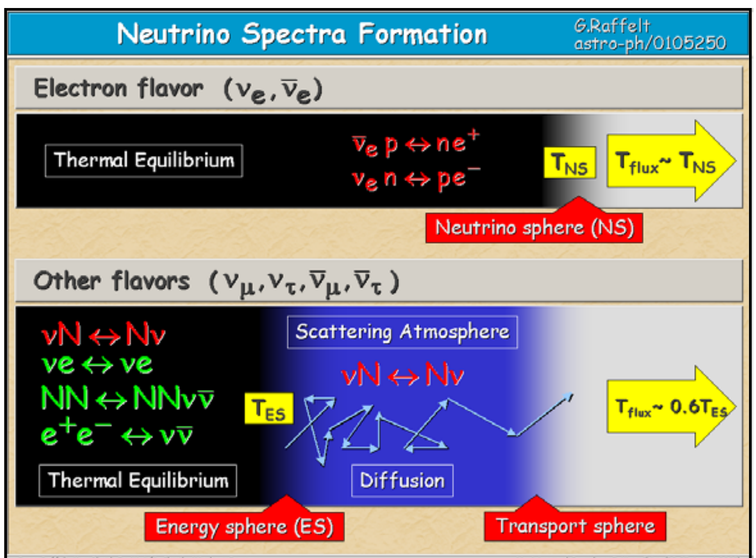
Flavor-Dependent Fluxes and Spectra

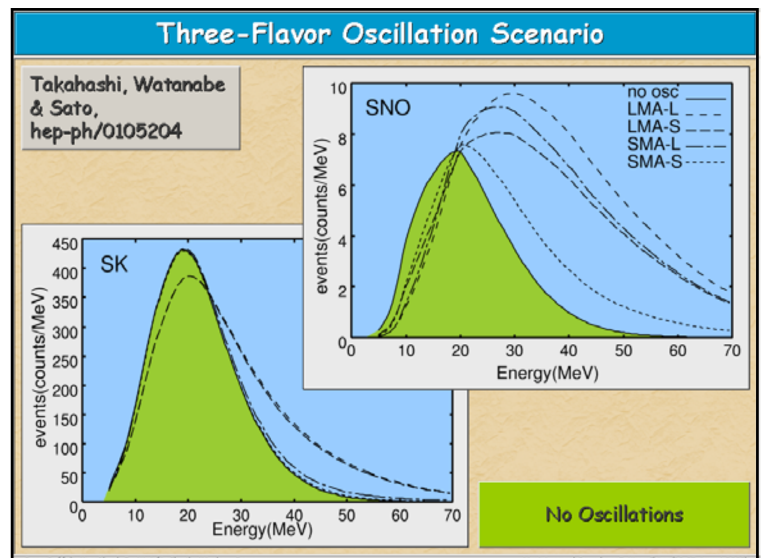


Numerical model of Totani, Sato, Dalhed & Wilson, ApJ 496 (1998) 216


Can we see the prompt neutrino burst?
Monte-Carlo example for early SN signal in Super-Kamiokande
Background from $\bar{\nu}_e + p \to n + e^+$
For SN at 10 kpc expect ~ 10 events from prompt burst $\nu_e + e \to e + \nu_e$
Events [/0.001 sec]
Time [sec]
0
0.02
0.04
0.06
0.08
0.1
5
10
Totani, Sato, Dalhed & Wilson, ApJ 496 (1998) 216
Georg Raffelt, Max-Planck-Institut für Physik, München, Germany
Neutrinos and Cosmology, Beijing, China (20-25 August 2002)

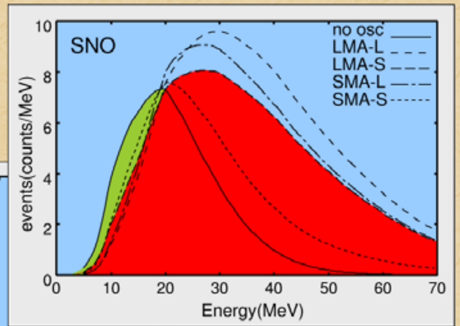

Neutrino Spectra Formation
G.Raffelt astro-ph/0105250
Electron flavor ($\nu_e, \bar{\nu}_e$)
Thermal Equilibrium
$\bar{\nu}_e\, p \leftrightarrow n e^+$
$\nu_e\, n \leftrightarrow p e^-$
Free streaming
Neutrino sphere (NS)
Other flavors ($\nu_\mu, \nu_\tau, \bar{\nu}_\mu, \bar{\nu}_\tau$)
$\nu N \leftrightarrow N \nu$
$\nu e \leftrightarrow \nu e$
$NN \leftrightarrow NN \nu\bar{\nu}$
$e^+ e^- \leftrightarrow \nu\bar{\nu}$
Thermal Equilibrium
Scattering Atmosphere
$\nu N \leftrightarrow N \nu$
Diffusion
Free streaming
Energy sphere (ES)
Transport sphere
Georg Raffelt, Max-Planck-Institut für Physik, München, Germany
Neutrinos and Cosmology, Beijing, China (20-25 August 2002)

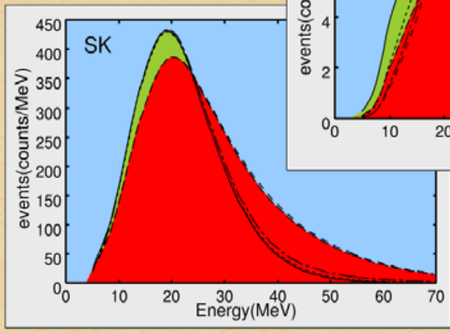

Neutrino Spectra Formation
G.Raffelt astro-ph/0105250
Electron flavor ($\nu_e, \bar{\nu}_e$)
Thermal Equilibrium
$\bar{\nu}_e\, p \leftrightarrow n e^+$
$\nu_e\, n \leftrightarrow p e^-$
T_{NS}
$T_{flux} \sim T_{NS}$
Neutrino sphere (NS)
Other flavors ($\nu_\mu, \nu_\tau, \bar{\nu}_\mu, \bar{\nu}_\tau$)
$\nu N \leftrightarrow N \nu$
$\nu e \leftrightarrow \nu e$
$NN \leftrightarrow NN \nu\bar{\nu}$
$e^+ e^- \leftrightarrow \nu\bar{\nu}$
Thermal Equilibrium
T_{ES}
Scattering Atmosphere
$\nu N \leftrightarrow N \nu$
Diffusion
$T_{flux} \sim 0.6 T_{ES}$
Energy sphere (ES)
Transport sphere
Georg Raffelt, Max-Planck-Institut für Physik, München, Germany
Neutrinos and Cosmology, Beijing, China (20-25 August 2002)


Three-Flavor Oscillation Scenario
Takahashi, Watanabe & Sato, hep-ph/0105204
SNO
no osc
LMA-L
LMA-S
SMA-L
SMA-S
events(counts/MeV)
Energy(MeV)
SK
events(counts/MeV)
Energy(MeV)
No Oscillations
Georg Raffelt, Max-Planck-Institut für Physik, München, Germany
Neutrinos and Cosmology, Beijing, China (20-25 August 2002)

Three-Flavor Oscillation Scenario

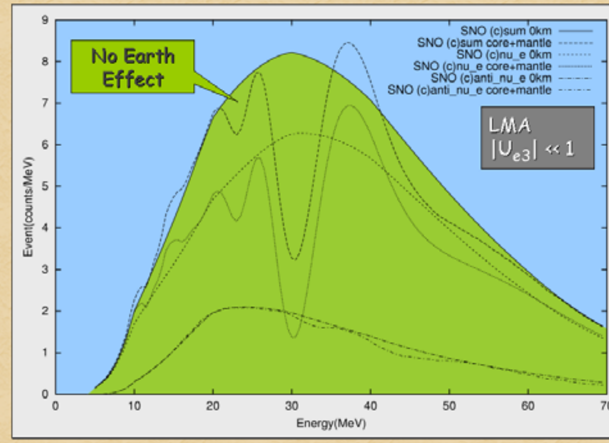
Takahashi, Watanabe & Sato, hep-ph/0105204





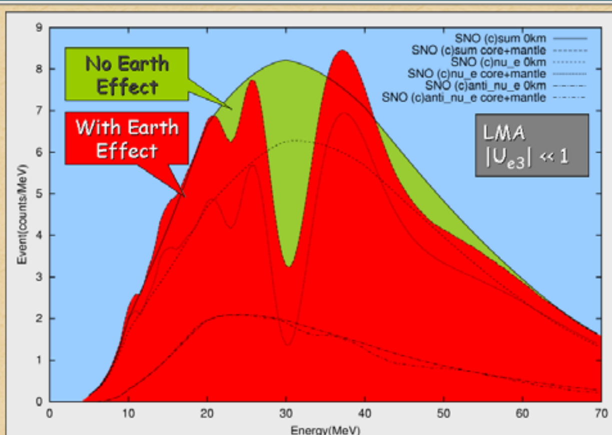
Oscillations with LMA and $|U_{e3}| \ll 1$

Earth Effect at SNO



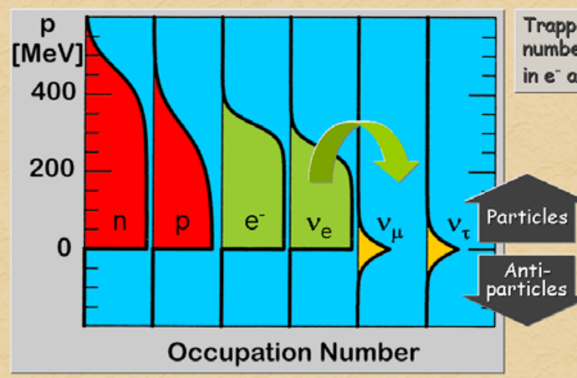
Takahashi, Watanabe & Sato, hep-ph/0012354

Earth Effect at SNO



Takahashi, Watanabe & Sato, hep-ph/0012354

Degenerate Fermi Seas in a Supernova Core



Trapped lepton number is stored in e^- and ν_e

In true thermal equilibrium with flavor mixing only <u>one</u> chemical potential for charged leptons and <u>one</u> for neutrinos

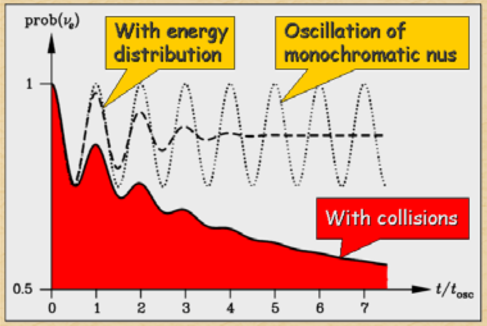
Time scale to achieve flavor equilibrium?

Flavor Relaxation in a Supernova Core

Neutrinos suffer collisions in a medium that can interrupt the coherence of flavor oscillations: The flavor content is "measured" and oscillations start from scratch from the "collapsed state".

Average oscillation probability $\frac{1}{2}\sin^2(2\Theta)$
Collision rate ~ damping rate Γ

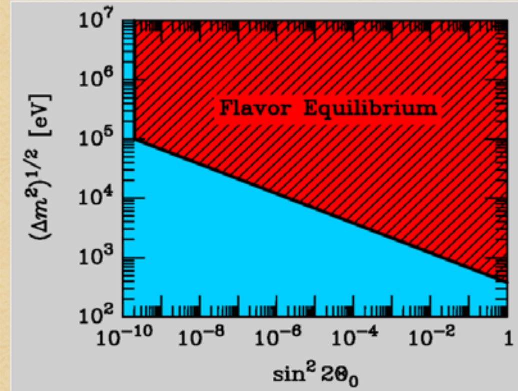
Conversion rate $\frac{1}{2}\sin^2(2\Theta)\,\Gamma$



Θ is the mixing angle in the medium.
In a SN core, the weak potential corresponds to Δm ~ 10-100 keV

Vast suppression of flavor conversion for sub-eV masses.

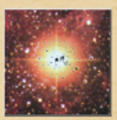
Flavor Conversion in a Supernova Core



Within ~ 1 sec flavor equilibrium is achieved between ν_e and ν_μ or ν_τ

Suppression of mixing angle by medium effects responsible for flavor-lepton number conservation in a supernova core

Conclusions of Lecture I



- Core-collapse supernova explosions probably explained by neutrino-driven delayed explosion mechanism
- But thus far no working numerical standard model
- Convection key to successful explosion?

High-statistics observation of a galactic SN is
- Crucial for empirical study of core-collapse event
- Not sensitive to sub-eV neutrino masses
- May differentiate between some mixing scenarios

If neutrino mixing parameters in currently favored regions
- Neutrino flavor oscillations not important for SN physics
- But crucial for detector signal interpretation
- Sterile nus and/or dipole moments can have strong effects

- Particle emission by supernova cores continues to provide most restrictive limits on various theories (axions, r.h. neutrinos, extra dimensions, ...)
- High-statistics observation would put these on firm grounds